

Neutrino Target-of-Opportunity Observations with Space-based and Suborbital Optical Cherenkov Detectors

Tonia M. Venters^{a,*}, Mary Hall Reno^b and John F. Krizmanic^{a,c,d} on behalf of the POEMMA and JEM-EUSO Collaborations

(a complete list of authors can be found at the end of the proceedings)

^aNASA Goddard Space Flight Center,
Astrophysics Science Division, Greenbelt, Maryland, USA

^bUniversity of Iowa,
Department of Physics and Astronomy, Iowa City, Iowa, USA

^cUniversity of Maryland, Baltimore County,
Center for Space Sciences and Technology,
Baltimore, Maryland, USA

^dCenter for Research and Exploration in Space Science & Technology
E-mail: tonia.m.venters@nasa.gov, mary-hall-reno@uiowa.edu

Cosmic-ray accelerators capable of reaching ultra-high energies are expected to also produce very-high energy neutrinos via hadronic interactions within the source or its surrounding environment. Many of the candidate astrophysical source classes are either transient in nature or exhibit flaring activity. Using the Earth as a neutrino converter, suborbital and space-based optical Cherenkov detectors, such as POEMMA and EUSO-SPB2, will be able to detect upward-moving extensive air showers induced by decaying tau-leptons generated from cosmic tau neutrinos with energies ~ 10 PeV and above. Both EUSO-SPB2 and POEMMA will be able to quickly repoint, enabling rapid response to astrophysical transient events. We calculate the transient sensitivity and sky coverage for both EUSO-SPB2 and POEMMA, accounting for constraints imposed by the Sun and the Moon on the observation time. We also calculate both detectors' neutrino horizons for a variety of modeled astrophysical neutrino fluences. We find that both EUSO-SPB2 and POEMMA will achieve transient sensitivities at the level of modeled neutrino fluences for nearby sources. We conclude with a discussion of the prospects of each mission detecting at least one transient event for various modeled astrophysical neutrino sources.

37th International Cosmic Ray Conference (ICRC 2021)
July 12th – 23rd, 2021
Online – Berlin, Germany

*Presenter

1. Introduction

Transient phenomena underpinned both of the recent landmark multimessenger events—the coincident neutrino and electromagnetic detection of a flare in blazar TXS0506+056 [1] and the coincident gravitational wave and electromagnetic detection of GW170817, a binary neutron star merger accompanied by a short gamma-ray burst [2, 3]. These events solidified the central role of Target-of-Opportunity (ToO) observations in the emerging multimessenger landscape for the coming era and beyond.

Many astrophysical transient sources are suspected of producing high-energy and very-high energy neutrinos, including tidal disruption events [*e.g.*, 4], compact object mergers [*e.g.*, 5, 6], gamma-ray bursts [*e.g.*, 7, 8], blazar flares [*e.g.*, 9], and possibly others. In typical astrophysical models, neutrinos¹ are primarily produced through the decay of charged pions and secondary muons generated through interactions of accelerated cosmic rays with matter and radiation within the source [*e.g.*, 10]. The expected at-source flavor ratio from pion decay ($\nu_e : \nu_\mu : \nu_\tau$)_{src}, is nearly (1 : 2 : 0); though, neutrino oscillations drive the observed flavor ratio towards (1 : 1 : 1)_⊕. This provides a channel through which indirect cosmic ray experiments may observe transient astrophysical neutrino sources by detecting upward-going extensive air showers (EASs) induced by decaying τ -leptons produced by ν_τ interactions in the Earth.

In this work, we focus on the unique contributions to multimessenger studies of astrophysical transients from the proposed space-based Probe of Extreme MultiMessenger Astrophysics (POEMMA) [11] and the upcoming second-generation Extreme Universe Space Observatory on a Super-Pressure Balloon (EUSO-SPB2) [12]. Both missions will monitor the Earth’s atmosphere in search of fluorescence² signals from ultra-high energy cosmic rays and neutrinos above 20 EeV and optical Cherenkov signals from upward-going EASs of tau neutrinos above 10 PeV. In order to detect these signals, both missions must operate during dark, moonless night conditions. While it is possible that POEMMA and EUSO-SPB2 could generate their own transient alerts, this work focuses on their capabilities for following up alerts generated by other multimessenger facilities. Table 1 provides a select list of relevant observatory specifications for POEMMA and EUSO-SPB2.

The POEMMA observatory design consists of two telescopes on board individual satellites orbiting the Earth in tandem with a period of ~ 95 min. Each spacecraft will be equipped with avionics that will allow them to quickly repoint (at a rate of 90° in ~ 500 s) in the direction of a transient source and track it through the tau neutrino detection region (up to 18° below the Earth’s limb). With these design features, POEMMA will have access to the entire dark sky within the time scale of one orbit and full-sky coverage within a few months. POEMMA is proposed for launch in 2028 and has a target mission lifetime of 3–5 years.

EUSO-SPB2 is a suborbital payload experiment that will be a pathfinder for POEMMA in many respects. It will feature a similar, though simplified, optical design for validating the technology for detecting optical Cherenkov signals from Earth-skimming tau neutrinos. It will also have a ToO operation mode that will enable pointing to a transient source by tilting the telescope in elevation and rotating the entire gondola in azimuth. The expected launch for EUSO-SPB2 is in Spring 2023 from Wanaka, New Zealand, with a target flight duration of ~ 100 days.

¹In this work, we do not distinguish between neutrinos and antineutrinos.

²The fluorescence detection capabilities of POEMMA and EUSO-SPB2 are discussed in other proceedings.

Observatory Specification	POEMMA	EUSO-SPB2
Altitude (h)	525 km	33 km
Field-of-View ($\Delta\phi \times \Delta\alpha$)	$30^\circ \times 9^\circ$	$12.8^\circ \times 6.4^\circ$
Max. Angle from Limb	18°	6.4°
Detection threshold	20 (40) phot m^{-2}	200 phot m^{-2}
Max. sky coverage	100%	74%

Table 1: Select list of observatory specifications for POEMMA and EUSO-SPB2.

2. ν_τ Acceptance and Transient Sensitivity Including Sun and Moon Constraints

The ν_τ acceptance and the transient sensitivity depend on the instantaneous effective area to point sources (see Fig. 1), which is defined as the area subtended on the ground by the Cherenkov cone given by [13, 14],

$$A_{\text{Ch}} \simeq \pi(v - s)^2 \times (\theta_{\text{Ch}}^{\text{eff}})^2, \quad (1)$$

where v is the distance between the detector at altitude h and the spot on the ground along the line-of-sight to a point source in the ν_τ detection region, s is the decay length for the τ -lepton, and $\theta_{\text{Ch}}^{\text{eff}}$ is the effective Cherenkov angle for upward-going EASs. The line-of-sight is viewed at an angle $\alpha(t)$ from detector nadir, corresponding to an elevation angle β_v with respect to the spot on the ground. The effective Cherenkov angle depends on the altitude at which the τ -lepton decays, the shower energy, and elevation angle of the τ -lepton emerging from the ground, β_{tr} [15]. Since the neutrino sources in question are distant point sources, we take $\beta_{\text{tr}} \simeq \beta_v$. The instantaneous acceptance to tau neutrinos with energy E_{ν_τ} is given by

$$A(\alpha(t), E_{\nu_\tau}) \simeq \int dP_{\text{obs}}(E_{\nu_\tau}, \beta_v, s) A_{\text{Ch}}(s), \quad (2)$$

where P_{obs} is the observation probability. P_{obs} is related to the probability for a τ -lepton to emerge at angle β_v (*i.e.*, the tau exit probability, p_{exit}), the τ -lepton decay distribution (p_{dec}), and the detection probability (p_{det}) by

$$P_{\text{obs}} = \iint p_{\text{exit}}(E_\tau | E_{\nu_\tau}, \beta_v) p_{\text{dec}}(s') p_{\text{det}}(E_\tau, \theta_{\text{Ch}}^{\text{eff}}, \beta_v, s') ds' dE_\tau. \quad (3)$$

For p_{exit} , we used the results from [15], which performed detailed simulations of ν_τ propagation through the Earth [see also 16]. [15] also performed detailed simulations propagating Cherenkov signals from upward-going EASs through the atmosphere to a space-based detector flying at altitude $h = 525$ km; we used these results to determine $\theta_{\text{Ch}}^{\text{eff}}$ and p_{det} for POEMMA. We performed new Cherenkov simulations to determine $\theta_{\text{Ch}}^{\text{eff}}$ and p_{det} for EUSO-SPB2 [see also 17].

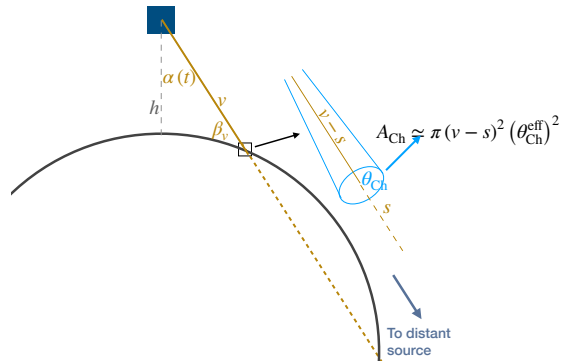


Figure 1: ν_τ detection geometry for point sources.

As the source moves through the detection region, the effective area and the ν_τ acceptance change with time. In determining the transient sensitivity, we compute the time-averaged ν_τ acceptance,

$$\langle A(E_{\nu_\tau}) \rangle = \frac{1}{T_0} \int_{T_0}^{t_0+T_0} dt A(\alpha(t), E_{\nu_\tau}) . \quad (4)$$

The all-flavor transient sensitivity is then given by

$$\text{Sensitivity} = \frac{2.44 \times 3}{\ln(10) \times E_{\nu_\tau} \times \langle A(E_{\nu_\tau}) \rangle} , \quad (5)$$

where we have taken the 90% unified confidence upper limit [18] per decade in energy ($2.44/\ln(10)$). In all scenarios considered here, we assume that both instruments can slew over 360° in azimuth.

In this work, we calculate separate estimates for the transient sensitivity for long-duration and short-duration events. For short-duration events, the relevant time scale for computing the time-averaged ν_τ acceptance is the observation time, which we take to be $\sim 10^3$ s. For long-duration events (those lasting $\gtrsim 1$ day), the relevant time scale is one that allows for adequate sampling of the variation in the effective area as the source moves through the detection region. For POEMMA, this time scale is $T_0 \simeq T_{\text{orb}} = 95$ minutes. For EUSO-SPB2, we set T_0 equal to the duration of the balloon flight in order to more robustly account for the constraints due to the Sun and the Moon (as discussed below). In this work, we consider times of 30 days and 100 days for a typical balloon flight and the target flight duration, respectively.

During long-duration events (time scales $\gtrsim 1$ day), POEMMA's spacecraft propulsion systems can maneuver its satellites closer together (separation ~ 25 km) in order to observe the source within overlapping instrument light pools (denoted ToO-stereo configuration). This allows for the use of time coincidence between the two instruments, enabling better rejection of the night-sky air glow background and lowering the energy threshold for neutrino detection. For events shorter than a day, the two POEMMA telescopes will observe the source independently in separate light pools (denoted ToO-dual configuration); this configuration doubles the effective area at the expense of observing with a higher energy threshold.

Both POEMMA and EUSO-SPB2 must operate in dark, moonless night conditions. We account for the effects of the Sun and the Moon depending on the scenario under consideration. For long-duration events, the time scales are sufficiently long that the Sun and the Moon are guaranteed to constrain the observation period. For space-based instruments, the reduction in exposure depends on the sky location of the source and the relative positions of the Sun, the Moon, the Earth, and position of the satellite in its orbit, resulting in a large number of possible configurations. For POEMMA, the average duty cycle³ ranges between 0.2–0.4 over the sky [13] [see also 19]; for our calculations of the POEMMA's long-duration transient sensitivity, we adopt a value of 30%.

For EUSO-SPB2, we perform a more detailed calculation of the average ν_τ acceptance, accounting for the reduction in observation time due to the Sun and the Moon. For the Sun, we set the acceptance equal to zero during those time periods in which the elevation of the Sun was greater than 18° below the horizon as viewed by EUSO-SPB2 (elevation angle of -5.8°). For the Moon, we calculate the average intensity of moonlight, including direct and scattered components, accounting for the

³Averaged over seven precession periods of POEMMA's orbital plane (7×54.3 days $\simeq 380$ days).

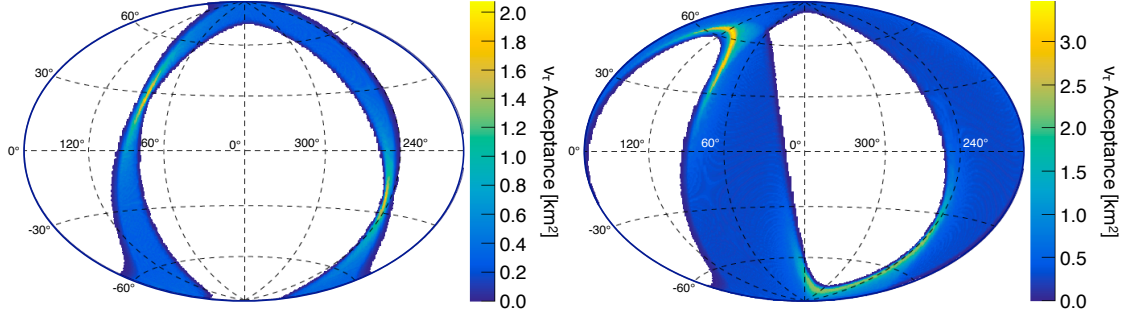


Figure 2: Sky plots of the ν_τ acceptance at $10^{8.5}$ GeV averaged over a period of a day when the Moon is close to full (left) and close to new (right).

phase. To that end, we used libRadtran, a radiative transfer library [20], to calculate the average intensity of sunlight (I_\odot) for an instrument flying at 33 km, assuming an Earth albedo of ~ 0.4 [21] and an atmospheric profile consistent with midlatitude winter and accounting for aerosols. The moonlight intensity is found by rescaling the solar intensity,

$$I_{\text{Moon}} = 0.07I_\odot \left(\frac{R_{\text{Moon}}}{d_{\oplus-\text{Moon}}} \right)^2 10^{-0.4(0.026|i|+4.0 \times 10^{-9}i^4)}, \quad (6)$$

where R_{Moon} is the radius of the Moon, $d_{\oplus-\text{Moon}}$ is the distance between the Earth and the Moon, and the last multiplicative factor accounts for the variation in intensity due to the phase of the Moon with i being the lunar phase angle in degrees [22].⁴ For the constraint due to the Moon, we set the acceptance equal to zero during those time periods in which I_{Moon} is greater than 10% of the estimated intensity of the night sky air glow ($\sim 18801 \text{ phot m}^{-2} \text{ ns}^{-1} \text{ sr}^{-1}$ integrated from 300–1000 nm to cover the wavelength range of the optical Cherenkov telescope). The impact of the Sun and the Moon can be seen in Figure 2, which plots the ν_τ acceptance averaged over a period of a day for two different phases of the Moon.

For short-duration events for both missions, the time scale of the observation considered here (10^3 s) is sufficiently short that observations may occur without interference from the Sun or the Moon. In this work, we consider short-duration scenarios in which observations are not constrained by the Sun and the Moon. We also assume that event occurs when the source is in the optimal position for ν_τ detection (*i.e.*, when source just dips below the Earth’s limb). As such, our short-duration scenarios should be considered “best-case” scenarios for neutrino ToO observations.

The left panels of Figures 3 and 4 show the long-duration all-flavor transient sensitivities for POEMMA and EUSO-SPB2, respectively. The sensitivities are presented as purple bands to demonstrate the variation across the sky. For comparison, modeled all-flavor neutrino fluences for a binary neutron star (BNS) merger [6] at various time scales and distances are also plotted. Figures 3 and 4 also provide sky plots of the time-averaged ν_τ acceptance at $10^{8.5}$ GeV to indicate the sky coverage for each instrument. The left panels of Figures 5 and 6 show the best-case all-flavor transient sensitivities to short-duration events for POEMMA and EUSO-SPB2, respectively. For comparison, modeled all-flavor neutrino fluences for various phases of a short gamma-ray burst [8] at various distances are also plotted. Figures 5 and 6 also provide sky plots of the time-averaged ν_τ

⁴Note that $0^\circ \leq i \leq 180^\circ$, where $i = 0^\circ$ corresponds to a full moon and $i = 180^\circ$ corresponds to a new moon.

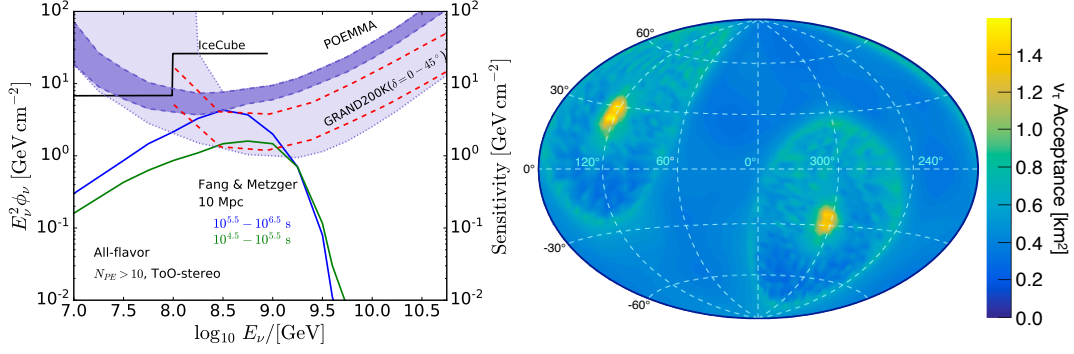


Figure 3: Left: All-flavor long-duration transient sensitivity band for POEMMA. For comparison, upper limits from IceCube for neutrino searches for GW170817 [23] and sensitivity projections for GRAND200k [24] are also plotted. Right: Sky plot of the time-averaged ν_τ acceptance at $10^{8.5}$ GeV.

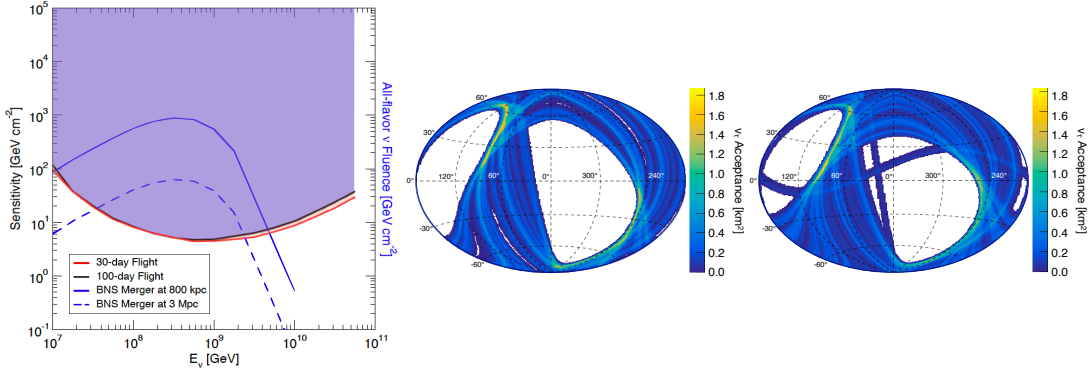


Figure 4: Left: All-flavor long-duration transient sensitivity band for EUSO-SPB2. Middle: Sky plot of the ν_τ acceptance at $10^{8.5}$ GeV for EUSO-SPB2 averaged over 30 days. Right: Sky plot of the ν_τ acceptance at $10^{8.5}$ GeV for EUSO-SPB2 averaged over 100 days.

acceptance at $10^{8.5}$ GeV assuming the best-case short-duration scenario. Figures 3–6 demonstrate that both POEMMA and EUSO-SPB2 will achieve transient sensitivities at the level of modeled neutrino fluences for nearby sources.

3. Prospects of Detecting Neutrinos for Candidate Transient Astrophysical Sources

In order to determine the probability for detecting at least one ToO from a given source class, we must calculate the neutrino horizon, defined as the redshift, z_{hor} , at which a given instrument can expect to detect one ν_τ event. The expected number of neutrino events from an astrophysical source at redshift z is given by

$$N_{\text{ev}} = \int \phi_{\nu_\tau}(E_\nu, z) \langle A(E_\nu) \rangle dE_\nu, \quad (7)$$

where $\phi_{\nu_\tau}(E_\nu, z)$ is the observed single-flavor neutrino fluence [13]. With z_{hor} , we can then find the cosmological volume in which POEMMA or EUSO-SPB2 would be able to detect tau neutrinos. Multiplying by the cosmological event rate for the source population taken from the literature gives the expected rate of ToOs, r , for the source class. Modeling the occurrence of transient events as a Poisson process, we can determine the probability of POEMMA or EUSO-SPB2 detecting at least

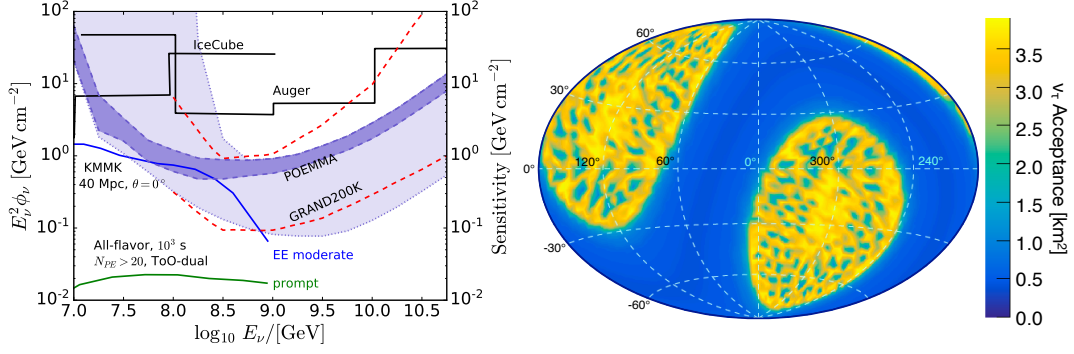


Figure 5: Left: Best-case all-flavor short-duration transient sensitivity band for POEMMA. For comparison, upper limits from IceCube and Auger for neutrino searches for GW170817 [23] and sensitivity projections for GRAND200K [24] are also plotted. Right: Sky plot of the time-averaged ν_τ acceptance at $10^{8.5}$ GeV for POEMMA assuming the best-case short scenario.

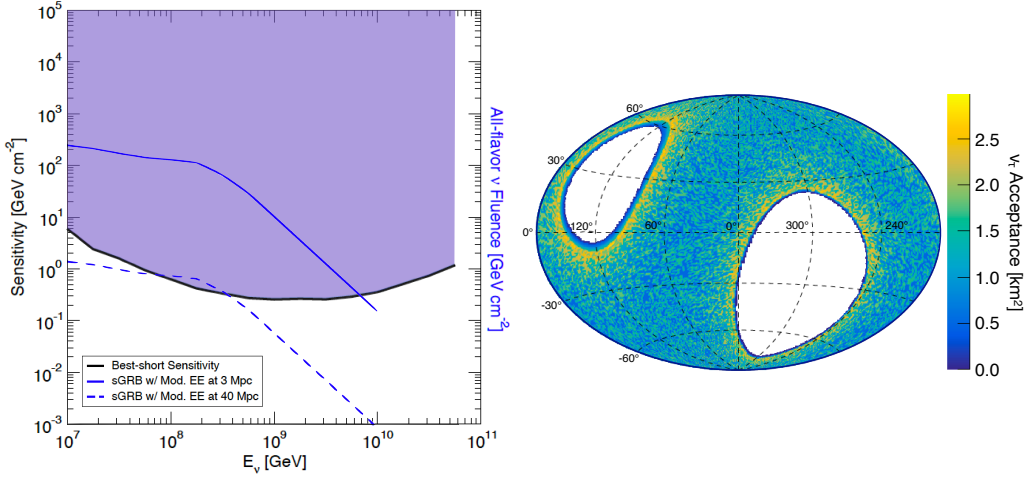


Figure 6: Left: Best-case all-flavor short-duration transient sensitivity band for EUSO-SPB2. Right: Sky plot of the time-averaged ν_τ acceptance at $10^{8.5}$ GeV for EUSO-SPB2 assuming the best-case short scenario.

one ToO as a function of time, t ,

$$P(\geq 1\text{ToO}) = 1 - P(0) = 1 - e^{-rt}. \quad (8)$$

Table 2 provides the neutrino horizon luminosity distances for POEMMA and EUSO-SPB2 for various neutrino source models. For these models, both POEMMA and EUSO-SPB2 will be able to detect neutrinos out to neighboring galaxies and beyond. The figure at right provides the Poisson probabilities of POEMMA detecting at least one ToO as a function of mission operation time for the various models. The most promising source classes for POEMMA are BNS mergers, binary black hole mergers, and tidal disruption events.

Acknowledgements – This work is supported by NASA grants 80NSSC19K0626 at the University of Maryland, Baltimore County, 80NSSC19K0460 at the Colorado School of Mines, 80NSSC19K0484 at the University of Iowa, and 80NSSC19K0485 at the University of Utah, 80NSSC18K0464 at Lehman College, and under proposal 17-APRA17-0066 at NASA/GSFC and JPL. The conceptual design of POEMMA was supported by NASA Probe Mission Concept Study grant NNX17AJ82G for the 2020 Decadal Survey Planning. Contributors to this work were supported in part by NASA awards 16-APROBES16-0023, NNX17AJ82G, NNX13AH54G, 80NSSC18K0246, and 80NSSC18K0473.

Source Class	EUSO-SPB2 ν_τ Horizon	POEMMA ν_τ Horizon
TDE $M_{\text{BH}} = 5 \times 10^6 M_\odot$ [4]	9 Mpc	128 Mpc
TDE Base Scenario [4]	4.5 Mpc	69 Mpc
BBH merger – Low Fluence [5]	6 Mpc	43 Mpc
BBH merger – High Fluence [5]	19 Mpc	137 Mpc
BNS merger [6]	2.3 Mpc	16 Mpc
sGRB w/ Mod. Ext. Emission [8]	25 Mpc	90 Mpc

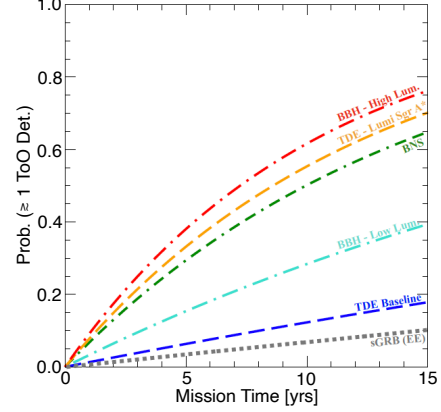


Table 2: Left: Table of horizon distances for EUSO-SPB2 and POEMMA for various astrophysical neutrino models. Right: Poisson probability of POEMMA detecting at least one ToO versus mission time for several modeled source classes. Probabilities for EUSO-SPB2 are $\lesssim 1\%$. Plot reproduced from [13].

References

- [1] M. G. Aartsen et al., *Science* **361**, eaat1378 (2018), [1807.08816](#).
- [2] B. P. Abbott et al., *Phys. Rev. Lett.* **119**, 161101 (2017), [1710.05832](#).
- [3] B. P. Abbott et al., *Astrophys. J.* **848**, L13 (2017), [1710.05834](#).
- [4] C. Lunardini and W. Winter, *Phys. Rev.* **D95**, 123001 (2017), [1612.03160](#).
- [5] K. Kotera and J. Silk, *Astrophys. J.* **823**, L29 (2016), [1602.06961](#).
- [6] K. Fang and B. D. Metzger, *The Astrophysical Journal* **849**, 153 (2017), ISSN 1538-4357.
- [7] E. Waxman and J. N. Bahcall, *Phys. Rev. Lett.* **78**, 2292 (1997), [astro-ph/9701231](#).
- [8] S. S. Kimura, K. Murase, P. Mészáros, and K. Kiuchi, *Astrophys. J.* **848**, L4 (2017), [1708.07075](#).
- [9] E. Waxman and J. N. Bahcall, *Phys. Rev. D* **59**, 023002 (1999), [hep-ph/9807282](#).
- [10] P. Lipari, M. Lusignoli, and D. Meloni, *Phys. Rev.* **D75**, 123005 (2007), [0704.0718](#).
- [11] A. Olinto et al., *Journal of Cosmology and Astroparticle Physics* **2021**, 007 (2021).
- [12] J. Eser, A. V. Olinto, L. Wiencke, et al., in *37th International Cosmic Ray Conference (these proceedings)* (2021).
- [13] T. M. Venters, M. H. Reno, J. F. Krizmanic, et al., *Phys. Rev. D* **102**, 123013 (2020).
- [14] M. H. Reno et al., in *37th International Cosmic Ray Conference (these proceedings)* (2021).
- [15] M. H. Reno, J. F. Krizmanic, and T. M. Venters, *Phys. Rev. D* **100**, 063010 (2019), [1902.11287](#).
- [16] S. Patel, M. H. Reno, et al., in *37th International Cosmic Ray Conference (these proceedings)* (2021).
- [17] A. L. Cummings, R. Aloisio, and J. F. Krizmanic, *Phys. Rev. D* **103**, 043017 (2021), [2011.09869](#).
- [18] G. J. Feldman and R. D. Cousins, *Phys. Rev. D* **57**, 3873 (1998), [physics/9711021](#).
- [19] C. Guépin, F. Sarazin, J. Krizmanic, et al., *JCAP* **03**, 021 (2019), [1812.07596](#).
- [20] C. Emde et al., *Geoscientific Model Development* **9**, 1647 (2016).
- [21] R. Peyrou-Lauga, in *47th International Conference on Environmental Systems* (2017), p. 142.
- [22] K. Krisciunas and B. E. Schaefer, *Publ. Astron. Soc. Pacif.* **103**, 1033 (1991).
- [23] A. Albert et al., *Astrophys. J.* **850**, L35 (2017), [1710.05839](#).
- [24] J. Alvarez-Muñiz et al. (GRAND) (2018), [1810.09994](#).

POEMMA Collaboration

A. V. Olinto,¹ J. Krizmanic,^{2,3} J. H. Adams,⁴ R. Aloisio,⁵ L. A. Anchordoqui,⁶ A. Anzalone,^{7,8} M. Bagheri,⁹ D. Barghini,¹⁰ M. Battisti,¹⁰ D. R. Bergman,¹¹ M. E. Bertaina,¹⁰ P. F. Bertone,¹² F. Bisconti,¹³ M. Bustamante,¹⁴ F. Cafagna,¹⁵ R. Caruso,^{16,8} M. Casolino,^{17,18} K. Černý,¹⁹ M. J. Christl,¹² A. L. Cummings,⁵ I. De Mitri,⁵ R. Diesing,¹ R. Engel,²⁰ J. Eser,¹ K. Fang,²¹ F. Fenu,¹⁰ G. Filippatos,²² E. Gazda,⁹ C. Guepin,²³ A. Haungs,²⁰ E. A. Hays,² E. G. Judd,²⁴ P. Klimov,²⁵ V. Kungel,²² E. Kuznetsov,⁴ Š. Mackovjak,²⁶ D. Mandát,²⁷ L. Marcelli,¹⁸ J. McEnery,² G. Medina-Tanco,²⁸ K.-D. Merenda,²² S. S. Meyer,¹ J. W. Mitchell,² H. Miyamoto,¹⁰ J. M. Nachtman,²⁹ A. Neronov,³⁰ F. Oikonomou,³¹ Y. Onel,²⁹ G. Osteria,³² A. N. Otte,⁹ E. Parizot,³³ T. Paul,⁶ M. Pech,²⁷ J. S. Perkins,² P. Picozza,^{18,34} L.W. Piotrowski,³⁵ Z. Plebaniak,¹⁰ G. Prévôt,³³ P. Reardon,⁴ M. H. Reno,²⁹ M. Ricci,³⁶ O. Romero Matamala,⁹ F. Sarazin,²² P. Schovánek,²⁷ V. Scotti,^{32,37} K. Shinozaki,³⁸ J. F. Soriano,⁶ F. Stecker,² Y. Takizawa,¹⁷ R. Ulrich,²⁰ M. Unger,²⁰ T. M. Venters,² L. Wiencke,²² D. Winn,²⁹ R. M. Young,¹² M. Zotov²

¹The University of Chicago, Chicago, IL, USA ²NASA Goddard Space Flight Center, Greenbelt, MD, USA ³Center for Space Science & Technology, University of Maryland, Baltimore County, Baltimore, MD, USA ⁴University of Alabama in Huntsville, Huntsville, AL, USA ⁵Gran Sasso Science Institute, L'Aquila, Italy ⁶City University of New York, Lehman College, NY, USA ⁷Istituto Nazionale di Astrofisica INAF-IASF, Palermo, Italy ⁸Istituto Nazionale di Fisica Nucleare, Catania, Italy ⁹Georgia Institute of Technology, Atlanta, GA, USA ¹⁰Universita' di Torino, Torino, Italy ¹¹University of Utah, Salt Lake City, Utah, USA ¹²NASA Marshall Space Flight Center, Huntsville, AL, USA ¹³Istituto Nazionale di Fisica Nucleare, Turin, Italy ¹⁴Niels Bohr Institute, University of Copenhagen, DK-2100 Copenhagen, Denmark ¹⁵Istituto Nazionale di Fisica Nucleare, Bari, Italy ¹⁶Universita' di Catania, Catania Italy ¹⁷RIKEN, Wako, Japan ¹⁸Istituto Nazionale di Fisica Nucleare, Section of Roma Tor Vergata, Italy ¹⁹Joint Laboratory of Optics, Faculty of Science, Palacký University, Olomouc, Czech Republic ²⁰Karlsruhe Institute of Technology, Karlsruhe, Germany ²¹Kavli Institute for Particle Astrophysics and Cosmology, Stanford University, Stanford, CA94305, USA ²²Colorado School of Mines, Golden, CO, USA ²³Department of Astronomy, University of Maryland, College Park, MD, USA ²⁴Space Sciences Laboratory, University of California, Berkeley, CA, USA ²⁵Skobel'syn Institute of Nuclear Physics, Lomonosov Moscow State University, Moscow, Russia ²⁶Institute of Experimental Physics, Slovak Academy of Sciences, Kosice, Slovakia ²⁷Institute of Physics of the Czech Academy of Sciences, Prague, Czech Republic ²⁸Instituto de Ciencias Nucleares, UNAM, CDMX, Mexico ²⁹University of Iowa, Iowa City, IA, USA ³⁰University of Geneva, Geneva, Switzerland ³¹Institut for fysikk, NTNU, Trondheim, Norway ³²Istituto Nazionale di Fisica Nucleare, Napoli, Italy ³³Université de Paris, CNRS, Astroparticule et Cosmologie, F-75013 Paris, France ³⁴Universita di Roma Tor Vergata, Italy ³⁵Faculty of Physics, University of Warsaw, Warsaw, Poland ³⁶Istituto Nazionale di Fisica Nucleare - Laboratori Nazionali di Frascati, Frascati, Italy ³⁷Universita' di Napoli Federico II, Napoli, Italy ³⁸National Centre for Nuclear Research, Lodz, Poland

JEM-EUSO Collaboration

G. Abdellaoui^{ah}, S. Abef^q, J.H. Adams Jr.^{pd}, D. Allard^{cb}, G. Alonso^{md}, L. Anchordoqui^{pe}, A. Anzalone^{eh,ed}, E. Arnone^{ek,el}, K. Asano^{fe}, R. Attallah^{ac}, H. Attoui^{aa}, M. Ave Pernas^{mc}, M. Bagheri^{ph}, J. Baláz^{la}, M. Bakiri^{aa}, D. Barghini^{el,ek}, S. Bartocci^{ei,ej}, M. Battisti^{ek,el}, J. Bayer^{dd}, B. Beldjilali^{ah}, T. Belenguer^{mb}, N. Belkhalifa^{aa}, R. Bellotti^{ea,eb}, A.A. Belov^{kb}, K. Benmessai^{aa}, M. Bertaina^{ek,el}, P.F. Bertone^{pf}, P.L. Biermann^{db}, F. Bisconti^{el,ek}, C. Blaksley^{ft}, N. Blanc^{oa}, S. Blin-Bondil^{ca,cb}, P. Bobik^{la}, M. Bogomilov^{ba}, K. Bolmgren^{na}, E. Bozzo^{ob}, S. Briz^{pb}, A. Bruno^{eh,ed}, K.S. Caballero^{hd}, F. Cafagna^{ea}, G. Cambi^{ei,ej}, D. Campana^{ef}, J.-N. Capdevielle^{cb}, F. Capel^{de}, A. Caramete^{ja}, L. Caramete^{ja}, P. Carlson^{na}, R. Caruso^{ec,ed}, M. Casolino^{ft,ei}, C. Cassardo^{ek,el}, A. Castellina^{ek,em}, O. Catalano^{eh,ed}, A. Cellino^{ek,em}, K. Černý^{bb}, M. Chikawa^{fc}, G. Chiritoi^{ja}, M.J. Christl^{pf}, R. Colalillo^{ef,eg}, L. Conti^{en,ei}, G. Cotto^{ek,el}, H.J. Crawford^{pa}, R. Cremonini^{el}, A. Creusot^{cb}, A. de Castro González^{pb}, C. de la Taille^{ca}, L. del Peral^{mc}, A. Diaz Damian^{cc}, R. Diesing^{pb}, P. Dinaucourt^{ca}, A. Djakonow^{ia}, T. Djemil^{ac}, A. Ebersoldt^{db}, T. Ebisuzaki^{ft}, J. Eser^{pb}, F. Fenu^{ek,el}, S. Fernández-González^{ma}, S. Ferrarese^{ek,el}, G. Filippatos^{pc}, W.I. Finch^{pc}, C. Fornaro^{en,ei}, M. Fouka^{ab}, A. Franceschi^{ee}, S. Franchini^{md}, C. Fuglesang^{na}, T. Fujii^{fg}, M. Fukushima^{fe}, P. Galeotti^{ek,el}, E. García-Ortega^{ma}, D. Gardiol^{ek,em}, G.K. Garipov^{kb}, E. Gascón^{ma}, E. Gazda^{ph}, J. Genci^{lb}, A. Golzio^{ek,el}, C. González Alvarado^{mb}, P. Gorodetzky^{ft}, A. Green^{pc}, F. Guarino^{ef,eg}, C. Guépin^{pl}, A. Guzmán^{dd}, Y. Hachisu^{ft}, A. Haungs^{db}, J. Hernández Carretero^{mc}, L. Hulet^{pc}, D. Ikeda^{fe}, N. Inoue^{fn}, S. Inoue^{ft}, F. Isgrò^{ef,eg}, Y. Itow^{fk}, T. Jammer^{dc}, S. Jeong^{gb}, E. Joven^{me}, E.G. Judd^{pa}, J. Jochum^{dc}, F. Kajino^{ff}, T. Kajino^{fi}, S. Kallí^{af}, I. Kaneko^{ft}, Y. Karadzhov^{ba}, M. Kasztelan^{ia}, K. Katahira^{ft}, K. Kawai^{ft}, Y. Kawasaki^{ft}, A. Kedadra^{aa}, H. Khales^{aa}, B.A. Khrenov^{kb}, Jeong-Sook Kim^{ga}, Soon-Wook Kim^{ga}, M. Kleifges^{db}, P.A. Klimov^{pd}, D. Kolev^{ba}, I. Kreykenbohm^{da}, J.F. Krizmanic^{pf,pk}, K. Królik^{ia}, V. Kungel^{pc}, Y. Kurihara^{fs}, A. Kusenko^{fr,pe}, E. Kuznetsov^{pd}, H. Lahmar^{aa}, F. Lakhdari^{ag}, J. Licandro^{me}, L. López Campano^{ma}, F. López Martínez^{pb}, S. Mackovjak^{la}, M. Mahdi^{aa}, D. Mandát^{bc}, M. Manfrin^{ek,el}, L. Marcelli^{ei}, J.L. Marcos^{ma}, W. Marszał^{ia}, Y. Martín^{me}, O. Martínez^{hc}, K. Mase^{fa}, R. Matev^{ba}, J.N. Matthews^{pg}, N. Mebarki^{ad}, G. Medina-Tanco^{ha}, A. Meshnikov^{db}, A. Merino^{ma}, M. Mese^{ef,eg}, J. Meseguer^{md}, S.S. Meyer^{pb}, J. Mimouni^{ad}, H. Miyamoto^{ek,el}, Y. Mizumoto^{fi}, A. Monaco^{ea,eb}, J.A. Morales de los Ríos^{mc}, M. Mastafa^{pd}, S. Nagataki^{ft}, S. Naitamor^{ab}, T. Napolitano^{ee}, J. M. Nachtman^{pi}, A. Neronov^{ob,cb}, K. Nomoto^{fr}, T. Nonaka^{fe}, T. Ogawa^{ft}, S. Ogio^{fl}, H. Ohmori^{ft}, A.V. Olinto^{pb}, Y. Onel^{pi}, G. Osteria^{ef}, A.N. Otte^{ph}, A. Pagliaro^{eh,ed}, W. Painter^{db}, M.I. Panasyuk^{kb}, B. Panico^{ef}, E. Parizot^{cb}, I.H. Park^{gb}, B. Pastircak^{la}, T. Paul^{pe}, M. Pech^{bb}, I. Pérez-Grande^{md}, F. Peretto^{ef}, T. Peter^{oc}, P. Picozza^{ei,ej,ft}, S. Pindado^{md}, L.W. Piotrowski^{pb}, S. Piraino^{dd}, Z. Plebaniak^{ek,el,ia}, A. Pollini^{oa}, E.M. Popescu^{ja}, R. Prevet^{ef,eg}, G. Prévôt^{cb}, H. Prieto^{mc}, M. Przybylak^{ia}, G. Puehlhofer^{dd}, M. Putis^{la}, P. Reardon^{pd}, M.H. Reno^{pi}, M. Reyes^{me}, M. Ricci^{ee}, M.D. Rodríguez Frías^{mc}, O.F. Romero Matamala^{ph}, F. Ronga^{ee}, M.D. Sabau^{mb}, G. Saccà^{ec,ed}, G. Sáez Cano^{mc}, H. Sagawa^{fe}, Z. Sahnoun^{ab}, A. Saito^{fg}, N. Sakaki^{ft}, H. Salazar^{hc}, J.C. Sanchez Balanzar^{ha}, J.L. Sánchez^{ma}, A. Santangelo^{dd}

A. Sanz-Andrés^{md}, M. Sanz Palomino^{mb}, O.A. Saprykin^{kc}, F. Sarazin^{pc}, M. Sato^{fo}, A. Scagliola^{ea,eb}, T. Schanz^{dd}, H. Schieler^{db}, P. Schovánek^{bc}, V. Scotti^{ef,es}, M. Serra^{me}, S.A. Sharakin^{kb}, H.M. Shimizu^{ff}, K. Shinozaki^{ia}, J.F. Soriano^{pe}, A. Sotgiu^{ei,ej}, I. Stan^{ja}, I. Strharsky^{la}, N. Sugiyama^{ff}, D. Supanitsky^{ha}, M. Suzuki^{fm}, J. Szabelski^{ia}, N. Tajima^{ft}, T. Tajima^{ft}, Y. Takahashi^{fo}, M. Takeda^{fe}, Y. Takizawa^{ft}, M.C. Talai^{ac}, Y. Tameda^{fp}, C. Tenzer^{dd}, S.B. Thomas^{pg}, O. Tibolla^{he}, L.G. Tkachev^{ka}, T. Tomida^{fh}, N. Tone^{ft}, S. Toscano^{ob}, M. Traïche^{aa}, Y. Tsunesada^{fl}, K. Tsuno^{ft}, S. Turriziani^{ft}, Y. Uchihori^{fb}, O. Vaduvescu^{me}, J.F. Valdés-Galicia^{ha}, P. Vallania^{ek,em}, L. Valore^{ef,es}, G. Vankova-Kirilova^{ba}, T.M. Venters^{pj}, C. Vigorito^{ek,el}, L. Villaseñor^{hb}, B. Vlcek^{mc}, P. von Ballmoos^{cc}, M. Vrabel^{lb}, S. Wada^{ft}, J. Watanabe^{fi}, J. Watts Jr.^{pd}, R. Weigand Muñoz^{ma}, A. Weindl^{db}, L. Wiencke^{pc}, M. Wille^{da}, J. Wilms^{da}, D. Winn^{pm}, T. Yamamoto^{ff}, J. Yang^{gb}, H. Yano^{fm}, I.V. Yashin^{kb}, D. Yonetoku^{fd}, S. Yoshida^{fa}, R. Young^{pf}, I.S. Zgura^{ja}, M.Yu. Zotov^{kb}, A. Zuccaro Marchi^{ft}

^{aa}Centre for Development of Advanced Technologies (CDTA), Algiers, Algeria, ^{ab}Dept. Astronomy, Centre Res. Astronomy, Astrophysics and Geophysics (CRAAG), Algiers, Algeria, ^{ac}LPR at Dept. of Physics, Faculty of Sciences, University Badji Mokhtar, Annaba, Algeria, ^{ad}Lab. of Math. and Sub-Atomic Phys. (LPMPs), Univ. Constantine I, Constantine, Algeria, ^{af}Department of Physics, Faculty of Sciences, University of M'sila, M'sila, Algeria, ^{ag}Research Unit on Optics and Photonics, UROP-CDTA, Sétif, Algeria, ^{ah}Telecom Lab., Faculty of Technology, University Abou Bekr Belkaid, Tlemcen, Algeria, ^{ba}St. Kliment Ohridski University of Sofia, Bulgaria, ^{bb}Joint Laboratory of Optics, Faculty of Science, Palacký University, Olomouc, Czech Republic, ^{bc}Institute of Physics of the Czech Academy of Sciences, Prague, Czech Republic, ^{ca}Omega, Ecole Polytechnique, CNRS/IN2P3, Palaiseau, France, ^{cb}Université de Paris, CNRS, AstroParticule et Cosmologie, F-75013 Paris, France, ^{cc}IRAP, Université de Toulouse, CNRS, Toulouse, France, ^{da}ECAP, University of Erlangen-Nuremberg, Germany, ^{db}Karlsruhe Institute of Technology (KIT), Germany, ^{dc}Experimental Physics Institute, Kepler Center, University of Tübingen, Germany, ^{dd}Institute for Astronomy and Astrophysics, Kepler Center, University of Tübingen, Germany, ^{de}Technical University of Munich, Munich, Germany, ^{ea}Istituto Nazionale di Fisica Nucleare - Sezione di Bari, Italy, ^{eb}Universita' degli Studi di Bari Aldo Moro and INFN - Sezione di Bari, Italy, ^{ec}Dipartimento di Fisica e Astronomia "Ettore Majorana", Università di Catania, Italy, ^{ed}Istituto Nazionale di Fisica Nucleare - Sezione di Catania, Italy, ^{ee}Istituto Nazionale di Fisica Nucleare - Laboratori Nazionali di Frascati, Italy, ^{ef}Istituto Nazionale di Fisica Nucleare - Sezione di Napoli, Italy, ^{es}Università di Napoli Federico II - Dipartimento di Fisica "Ettore Pancini", Italy, ^{eh}INAF - Istituto di Astrofisica Spaziale e Fisica Cosmica di Palermo, Italy, ^{ei}Istituto Nazionale di Fisica Nucleare - Sezione di Roma Tor Vergata, Italy, ^{ej}Università di Roma Tor Vergata - Dipartimento di Fisica, Roma, Italy, ^{ek}Istituto Nazionale di Fisica Nucleare - Sezione di Torino, Italy, ^{el}Dipartimento di Fisica, Università di Torino, Italy, ^{em}Osservatorio Astrofisico di Torino, Istituto Nazionale di Astrofisica, Italy, ^{en}Uninettuno University, Rome, Italy, ^{fa}Chiba University, Chiba, Japan, ^{fb}National Institutes for Quantum and Radiological Science and Technology (QST), Chiba, Japan, ^{fc}Kindai University, Higashi-Osaka, Japan, ^{fd}Kanazawa University, Kanazawa, Japan, ^{fe}Institute for Cosmic Ray Research, University of Tokyo, Kashiwa, Japan, ^{ff}Konan University, Kobe, Japan, ^{fg}Kyoto University, Kyoto, Japan, ^{fh}Shinshu University, Nagano, Japan, ^{fi}National Astronomical Observatory, Mitaka, Japan, ^{fj}Nagoya University, Nagoya, Japan, ^{fk}Institute for Space-Earth Environmental Research, Nagoya University, Nagoya, Japan, ^{fl}Graduate School of Science, Osaka City University, Japan, ^{fm}Institute of Space and Astronautical Science/JAXA, Sagami-hara, Japan, ^{fn}Saitama University, Saitama, Japan, ^{fo}Hokkaido University, Sapporo, Japan, ^{fp}Osaka Electro-Communication University, Neyagawa, Japan, ^{fq}Nihon University Chiyoda, Tokyo, Japan, ^{fr}University of Tokyo, Tokyo, Japan, ^{fs}High Energy Accelerator Research Organization (KEK), Tsukuba, Japan, ^{ft}RIKEN, Wako, Japan, ^{ga}Korea Astronomy and Space Science Institute (KASI), Daejeon, Republic of Korea, ^{gb}Sungkyunkwan University, Seoul, Republic of Korea, ^{ha}Universidad Nacional Autónoma de México (UNAM), Mexico, ^{hb}Universidad Michoacana de San Nicolás de Hidalgo (UMSNH), Morelia, Mexico, ^{hc}Benemérita Universidad Autónoma de Puebla (BUAP), Mexico, ^{hd}Universidad Autónoma de Chiapas (UNACH), Chiapas, Mexico, ^{he}Centro Mesoamericano de Física Teórica (MCTP), Mexico, ^{ia}National Centre for Nuclear Research, Lodz, Poland, ^{ib}Faculty of Physics, University of Warsaw, Poland, ^{ja}Institute of Space Science ISS, Magurele, Romania, ^{ka}Joint Institute for Nuclear Research, Dubna, Russia, ^{kb}Skobeltsyn Institute of Nuclear Physics, Lomonosov Moscow State University, Russia, ^{kc}Space Regatta Consortium, Korolev, Russia, ^{la}Institute of Experimental Physics, Kosice, Slovakia, ^{lb}Technical University Kosice (TUKE), Kosice, Slovakia, ^{ma}Universidad de León (ULE), León, Spain, ^{mb}Instituto Nacional de Técnica Aeroespacial (INTA), Madrid, Spain, ^{mc}Universidad de Alcalá (UAH), Madrid, Spain, ^{md}Universidad Politécnica de Madrid (UPM), Madrid, Spain, ^{me}Instituto de Astrofísica de Canarias (IAC), Tenerife, Spain, ^{na}KTH Royal Institute of Technology, Stockholm, Sweden, ^{oa}Swiss Center for Electronics and Microtechnology (CSEM), Neuchâtel, Switzerland, ^{ob}ISDC Data Centre for Astrophysics, Versoix, Switzerland, ^{oc}Institute for Atmospheric and Climate Science, ETH Zürich, Switzerland, ^{pa}Space Science Laboratory, University of California, Berkeley, CA, USA, ^{pb}University of Chicago, IL, USA, ^{pc}Colorado School of Mines, Golden, CO, USA, ^{pd}University of Alabama in Huntsville, Huntsville, AL, USA, ^{pe}Lehman College, City University of New York (CUNY), NY, USA, ^{pf}NASA Marshall Space Flight Center, Huntsville, AL, USA, ^{pg}University of Utah, Salt Lake City, UT, USA, ^{ph}Georgia Institute of Technology, USA, ^{pi}University of Iowa, Iowa City, IA, USA, ^{pj}NASA Goddard Space Flight Center, Greenbelt, MD, USA, ^{pk}Center for Space Science & Technology, University of Maryland, Baltimore County, Baltimore, MD, USA, ^{pl}Department of Astronomy, University of Maryland, College Park, MD, USA, ^{pm}Fairfield University, Fairfield, CT, USA