

# Measuring the Neutrino Flux in Segments along the Galactic Plane with IceCube

## The IceCube Collaboration

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Gamma-ray emission from the plane of the Milky Way is understood as partly originating from the interaction of cosmic rays with the interstellar medium. The same interaction is expected to produce a corresponding flux of neutrinos. In 2023, IceCube reported the first observation of this galactic neutrino flux at  $4.5\sigma$  confidence level. The analysis relied on neutrino flux predictions – based on gamma ray observations – to model the expected neutrino emission from the galactic plane. Three signal hypotheses describing different possible spatial and energy distributions were tested, where the single free parameter in each test was the normalization of the neutrino flux. We present first results of an analysis that can improve the characterization of Galactic neutrino emission by dividing the galactic plane into segments in galactic longitude. An unbinned maximum-likelihood analysis is used that can fit the spectral index and the flux normalization separately in each segment. While gamma ray telescopes can not differentiate between hadronic and leptonic emission, neutrino production must come from hadronic processes. Measuring a spectral index can further help to understand the contribution of unresolved neutrino sources inside the galactic plane. This work uses a full-sky cascade dataset and provides model-independent insight into the variation of the neutrino flux and energy distribution from different regions of the galactic plane.

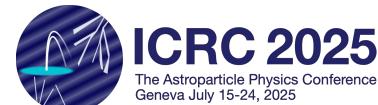
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## 1. Neutrino Flux from the Milky Way

For many forms of astronomy, the Milky Way is the most prominent source in the sky. It has been observed throughout the electromagnetic spectrum from radio waves to gamma rays. The gamma rays emitted in the Galactic Plane (GP) are understood to be partially produced by cosmic rays interacting with the interstellar medium. In these hadronic interactions, charged ( $\pi^+, \pi^-$ ) and neutral ( $\pi^0$ ) pions are created approximately at a ratio of 1:1:1. The neutral pions decay into two photons ( $\pi^0 \rightarrow \gamma\gamma$ ), producing a diffuse gamma ray emission in the GP, which has been observed beyond 100 TeV by the LHAASO [1] Tibet AS $\gamma$  [2] and HAWC [3] observatories. The charged pions decay into neutrinos ( $\pi^+ \rightarrow \mu^+ \nu_\mu \rightarrow e^+ \nu_\mu \bar{\nu}_\mu \nu_e$ ). Two years ago, the IceCube collaboration announced the first observation of neutrino emission from the GP at  $4.5\sigma$  confidence [4]. This discovery-oriented analysis relied on templates, which model the spatial and energy distribution of the diffuse neutrino emission using the measurement of diffuse gamma ray emission at lower energies from Fermi-LAT [5, 6]. The shape of the spectrum was fixed to the model prediction, and the single free fit parameter was the normalization of the spectrum. We present a new analysis method, which does not rely on a specific model assumption, but instead allows fitting a single power law in different regions of the GP. This method provides a flux normalization and spectral index in each of the chosen segments of the GP.

The IceCube observatory consists of 5160 digital optical modules, each equipped with a photomultiplier tube, frozen inside a volume of one cubic kilometer of glacial ice 1.5 kilometers below the geographic South Pole. When a high-energy neutrino interacts with the ice, it produces charged secondary particles, which emit Cherenkov light. The Cherenkov light can be used to reconstruct the direction and energy of the primary neutrino. IceCube records two main event topologies: “tracks” consist of Cherenkov light emitted by a muon along a line. Muons are created in charged-current  $W^\pm$  boson interactions between muon neutrinos and atomic nuclei in the ice. Tracks provide the best angular resolution ( $< 1^\circ$ ), but suffer from a background of atmospheric muons in the southern sky, which occur one million times more frequently than any neutrino interaction observable by IceCube. “Cascades” occur when any flavor of neutrino interacts inside the detector volume via a neutral current  $Z^0$  boson interaction or electron neutrinos  $\nu_e$  and tau neutrinos  $\nu_\tau$  interact via any interaction type. The light from these cascades radiates almost spherically and the angular resolution is thus poorer ( $\sim 5^\circ$  [7]). Cascades are essential to gain sensitivity in the southern sky, since they are distinguishable from down going atmospheric muons.

The observation of the GP was enabled by novel machine learning techniques used to select and reconstruct cascades [8, 9], which were used to increase the previous cascade selection by a factor of 20 to approximately 60 000 events. The same 60 000 cascade events were used to perform the analysis presented here.

## 2. Segmenting the Signal Hypothesis

This work uses the same unbinned maximum likelihood technique as previous IceCube neutrino source searches [10, 11]. The likelihood

$$\mathcal{L}(n_s, \gamma) = \prod_{i=1}^N \left[ \frac{n_s}{N} S_i(E_i, \delta_i, \alpha_i, \sigma_i, \gamma) + \frac{N - n_s}{N} B_i(E_i, \delta_i) \right] \quad (1)$$

is evaluated for each event  $i$  with its reconstructed quantities energy  $E_i$ , declination  $\delta_i$ , right ascension  $\alpha_i$  and angular uncertainty  $\sigma_i$ . The number of signal neutrinos  $n_s$  relative to the total number of neutrino events  $N$  and the spectral index of the signal probability density function (PDF)  $\gamma$  is optimized to maximize the likelihood of the data. The signal PDFs  $S_i$  for each event is constructed using Monte Carlo (MC) simulation, while the background PDFs  $B_i$  are derived in a data-driven way. Previously, it was assumed that in a given declination band the data is dominated by background events and averaging over right ascension gives a good estimation of the background PDF. For the given dataset the contribution from the GP cannot be neglected any longer. Thus, the right ascension average of the signal PDF  $\tilde{S}_i$  is subtracted from the averaged data PDF  $\tilde{D}_i$  using the method described in Reference [12], which was also used in the 2023 GP analysis [4]. This gives the signal subtracted background PDF

$$B_i^{\text{SigSub}}(E_i, \delta_i, \gamma) = \frac{N}{N - n_s} \left( \tilde{D}(\delta_i, E_i) - \frac{n_s}{N} \tilde{S}_i(\delta_i, E_i, \sigma_i, \gamma) \right). \quad (2)$$

Substituting the last term in Equation (1) with Equation (2) results in the Likelihood used in this work.

The analysis presented here introduces additional parameters into the likelihood by *segmenting* the signal PDF  $S_i$  into  $M$  parts

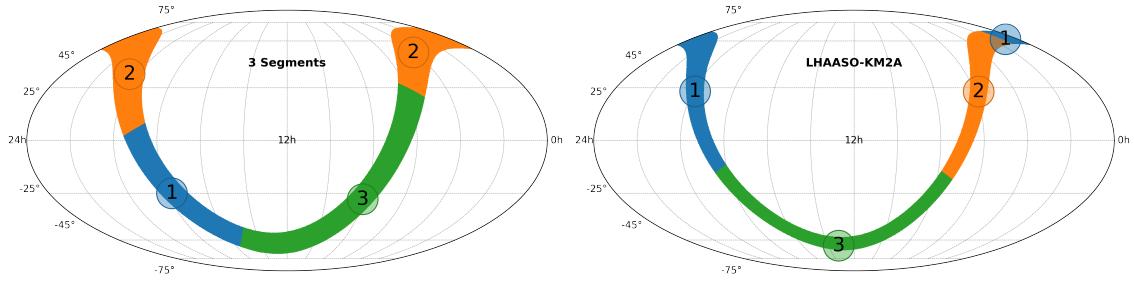
$$S_i(E_i, \delta_i, \alpha_i, \sigma_i, w_1, \dots, w_M, \gamma_1, \dots, \gamma_M) = \sum_{k=1}^M w_k S_i^k(E_i, \delta_i, \alpha_i, \sigma_i, \gamma_k) \quad (3)$$

where  $w_1, \dots, w_M$  are the weights assigned to each segment's PDF  $S_i^k$  and  $\gamma_k$  their spectral indices. In order to make  $S_i$  a PDF when segmenting, the weights must be normalized  $\sum_k w_k = 1$  and positive  $w_k \geq 0$ , which is implemented via a softmax-transformation. The spatial shape of the segments and their number can be freely chosen. The choice will be called a “segmentation scheme” here. In this work we consider spatially disjunct, homogeneously emitting regions, aligning with the galactic plane in longitude. They are presented in the following sections. Equivalently, with Equation (3), the likelihood Equation (1) can be written as

$$\mathcal{L}(n_1, \dots, n_M, \gamma_1, \dots, \gamma_M) = \prod_{i=1}^N \left[ \frac{1}{N} \sum_{k=1}^M n_k S_i^k(E_i, \delta_i, \alpha_i, \sigma_i, \gamma_k) + \frac{N - n_s}{N} B_i^{\text{SigSub}}(E_i, \delta_i, \gamma_k) \right] \quad (4)$$

where  $n_k = w_k n_s$  is the number of signal neutrinos observed from a given segment  $k$ . The total number of signal neutrinos can then be obtained by summing the contribution of all segments  $n_s = \sum_k n_k$ . The likelihood-ratio test statistic

$$\Lambda = 2 \ln \frac{\mathcal{L}(n_k = \hat{n}_k, \gamma_k = \hat{\gamma}_k)}{\mathcal{L}(n_s = 0)} \quad (5)$$



(a) The 3 segments “generic” segmentation scheme. (b) Scheme aligning with the LHAASO GP analysis region.

**Figure 1:** Two of the eight used segmentation schemes of the galactic plan shown in equatorial coordinates in the Mollweide projection.

is defined by the ratio of the likelihood where all parameters  $n_k, \gamma_k$  ( $k = 1, \dots, M$ ) are optimized to maximize the likelihood, compared to the null hypothesis of no neutrino emission  $n_s = 0$ .

## 2.1 Segmentation Schemes

Overall, eight different segmentation schemes are used to analyze the data. This will result in eight test statistics. First, three “generic” segmentation schemes are defined, which are chosen to obtain insights into the changing flux and spectrum at different galactic longitudes. All three generic segmentation schemes have segments that extend in galactic latitude from  $-8^\circ$  to  $8^\circ$ , which means they have a height of  $16^\circ$ . This aligns with the inner galactic analysis region of the Fermi-LAT telescope in Reference [5], but is also a value which fits this analysis. Since the cascades used in this analysis have an average angular uncertainty of  $7^\circ$  [4], choosing a smaller height than around  $14^\circ$  does not improve the sensitivity of this analysis. The *first generic segmentation scheme* divides the galaxy into an inner region from galactic longitude  $-60^\circ$  to  $60^\circ$  and *one outer region*  $-180^\circ$  to  $-60^\circ$ ,  $60^\circ$  to  $180^\circ$ . The *second generic segmentation scheme* divides the galaxy into an inner region from  $-40^\circ$  to  $40^\circ$ , a left arm from  $-180^\circ$  to  $-40^\circ$  as well as a right arm from  $40^\circ$  to  $180^\circ$ . This “3 Segments scheme” is visualized in Figure 1a in equatorial coordinates. Then, to stress the analysis method, a segmentation scheme with six equal-size, adjacent segments is defined, each spanning a width in galactic longitude of  $60^\circ$ . The central segment of that scheme is centered around the galactic center from  $-30^\circ$  to  $30^\circ$  in galactic longitude. An overview of the different generic segmentation schemes along with their unblinding results is provided in Table 1.

In order to compare IceCube’s GP results in a more model-independent way than previously possible, five segmentation schemes aligning with analysis regions of high energy  $\gamma$  ray observatories are defined. These are HAWC [3], H.E.S.S. [13], LHAASO[1, 14] and Tibet AS $\gamma$  [2], an overview is provided in Table 2. The two analysis regions of Tibet AS $\gamma$  overlap and must therefore be split into two different segmentation schemes. The chosen segmentation schemes are constructed by having the first segment(s) aligning with the analysis region(s) of the experiments and the last segment is the remaining GP. This is illustrated in Figure 1b for the LHAASO segmentation scheme in equatorial coordinates. There, Segment 3 covers the part of the GP which is not observed by LHAASO. The obtained results can then be compared to a neutrino equivalent flux of the  $\gamma$ -fluxes. Under simple assumptions they evaluate to  $E_\nu^2 \frac{dN_\nu^{\text{All-Flavor}}}{dE_\nu} = {}^3/2 E_\gamma^2 \frac{dN_\gamma}{dE_\gamma}$  and  $E_\gamma = 2E_\nu$  [15, 16].

**Table 1:** Overview over the generic segmentation schemes.  $l_{\min/\max}$  is the minimal/maximal galactic longitude of the given segment. All segments are in between  $\pm 8^\circ$  of galactic latitude  $b$ . The summed number of signal neutrinos  $n_s$  from all segments and the significance of the analysis is also given.

Scheme	Segment $k$	$l_{\min}$	$l_{\max}$	Significance	$n_s = \sum_k n_k$
2 Segments	1	$-60^\circ$	$60^\circ$	$3.81 \sigma$	643
	2	$60^\circ$	$-60^\circ$		
3 Segments	1	$-40^\circ$	$40^\circ$	$3.84 \sigma$	643
	2	$40^\circ$	$180^\circ$		
	3	$-180^\circ$	$-40^\circ$		
6 Segments	1	$-30^\circ$	$30^\circ$	$3.58 \sigma$	808
	2	$30^\circ$	$90^\circ$		
	3	$90^\circ$	$150^\circ$		
	4	$150^\circ$	$-150^\circ$		
	5	$-150^\circ$	$-90^\circ$		
	6	$-90^\circ$	$-30^\circ$		

### 3. Results

This analysis is designed to gain a deeper understanding of the previously reported  $4.5\sigma$  observation of the GP. Additional fit parameters are thus introduced, allowing for a more differentiated measurement at the cost of global significance.

Despite the expected decrease in rejection power of the null hypothesis, the  $p$ -value can still be computed with the standard methodology of comparing the obtained test statistic  $\Lambda$  Equation (5), against the distribution of test statistics of background trials. Background trials are constructed by randomizing the right-ascension  $\alpha$  of each event. The unblinded significances along their best fit segment-wise number of signal neutrinos  $n_k$  are presented in Tables 1 and 2. The most significant segmentation scheme turned out to be the one with 3 Segments with a value of  $p = 6.1 \times 10^{-5}$  ( $3.84\sigma$ ). This result follows what is expected from this analysis method. When injecting simulation using the best fit value of the  $\pi^0$  template from IceCube’s previous GP template search [4], the median obtained significance is  $3.57\sigma$  for this analysis.

The background distribution of  $\Lambda$  for each segmentation scheme is very different from one-another, due to the different number of free parameters. In order to obtain a global  $p$ -value of all analyses, a trials correction must be performed. This is done with correlated background trials, in which for one given background scramble all 8 analyses are performed and the smallest  $p$ -value among them is saved. The smallest unblinded  $p$ -value can then be compared to the distribution of smallest  $p$ -values on background trials. For this analysis this gives an overall value of  $p = 2.97 \times 10^{-4}$  ( $3.43\sigma$ ), which is a factor of 4.9 higher than the best local  $p$ -value. This is better than applying the Šidák correction [17], which would result in a trials factor of  $(1 - (1 - p)^8) \approx 8p$ .

**Table 2:** Overview over the 5 segmentation schemes used to align with the galactic analysis regions of  $\gamma$  ray telescopes. The region along the GP not covered by the  $\gamma$  ray experiments is analysed as well, and corresponds to the latitude range not covered in the table. For LHAASO with one segmentation scheme two regions are analysed and both are provided in the table above each other. The significance of the scheme is given, as well as the sum of the best fit number of signal neutrinos from all segments  $n_s$ . The significance and number of neutrinos include also the segments not shown in the table and thus the whole range of galactic longitude.

Experiment	$l_{\min}$	$l_{\max}$	$ b _{\max}$	Significance	$n_s = \sum_k n_k$
HAWC [3]	43°	73°	5°	3.01 $\sigma$	664
H.E.S.S [13]	-110°	65°	3°	3.06 $\sigma$	569
LHAASO [1, 14]	15°	125°	5°	2.73 $\sigma$	371
	125°	-55°	5°		
Tibet AS $\gamma$ I [2]	25°	100°	5°	2.96 $\sigma$	660
Tibet AS $\gamma$ II [2]	50°	200°	5°	2.98 $\sigma$	626

It shows that the analyses are, as expected, highly correlated.

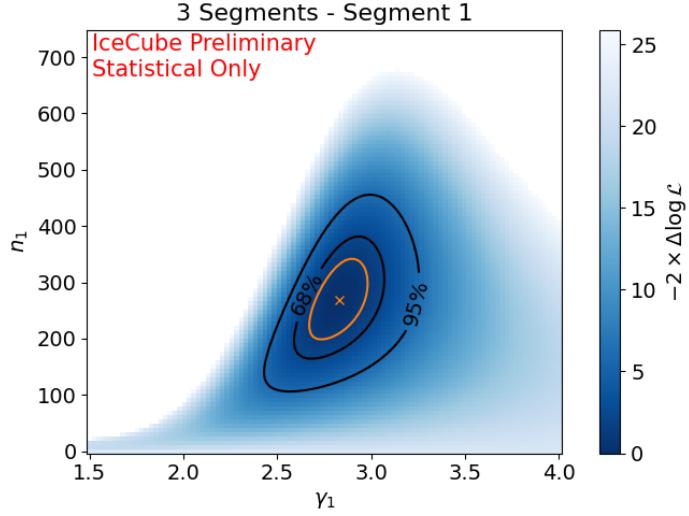
### 3.1 Likelihood Contours

To obtain uncertainty intervals or contours on the obtained parameters the profile likelihood method can be used. In practice for this analysis, it means that one can evaluate the parameters  $n_k$  and  $\gamma_k$  of one segment through a grid of points, while re-maximizing the likelihood by varying all other parameters. Since the segments are spatially disjunct and large compared to the angular uncertainty, the segments are nearly perfectly uncorrelated among each other, so profiling over the other segments makes little difference. The likelihood contour for the inner galaxy is shown in Figure 2 and the two contours for the arms of the galaxy are shown in Figure 3. The confidence contours are drawn by assuming Wilks' theorem [18]. For convenience, the contours assuming one degree of freedom in the  $\chi^2$  distribution are shown in orange, which enables a direct extraction of the profiled uncertainties per-parameter.

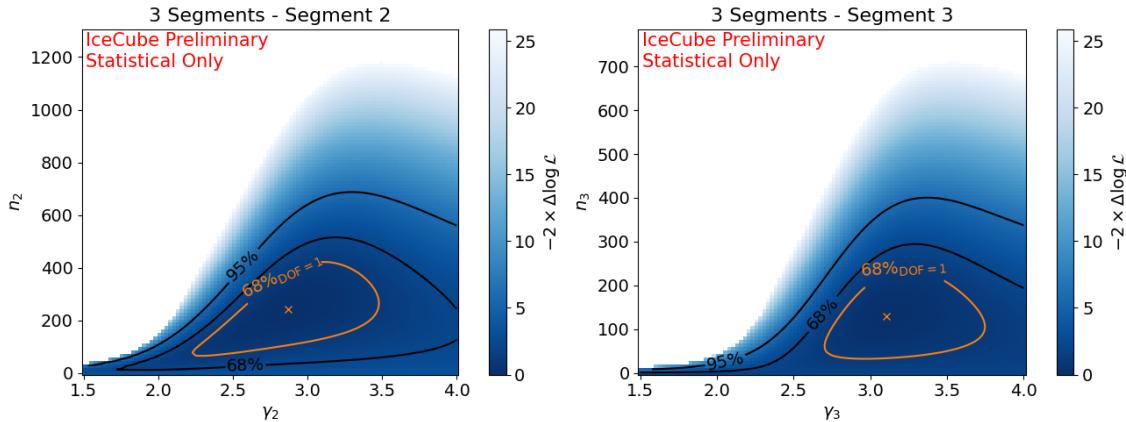
## 4. Conclusion

This work presents a first application of a new method to characterize neutrino emission from the GP. Applied to the same dataset used in the original observation of the GP [4], it can provide confidence regions of the flux normalization and the spectral index of a simple unbroken power law, as shown in Figures 2 and 3. The first measurement of the spectral index  $\gamma$  of the neutrino flux in the inner galaxy is shown in Figure 2. The presented method does not rely on a fine-grained template of the spatial distribution of neutrino emission from the GP, but rather makes the simple choice of assuming uniform emission in each segment. While the outer regions yield larger contours, they are expected to shrink with updated datasets.

In the future, this analysis can be applied to updated datasets using updated reconstructions, combined event topologies and more data. The data-driven estimation of the background distribution



**Figure 2:** The  $n_k$ - $\gamma_k$  likelihood contour of the inner galaxy ( $|l| < 40^\circ$ ,  $|b| < 8^\circ$  “Segment 1”) in the 3-segments generic segmentation scheme, visualized in blue in Figure 1a. The black 68 % and 95 % contours are drawn assuming Wilks’ theorem assuming 2 degrees of freedom, the orange 68 % contour is drawn assuming Wilks’ theorem with 1 degree of freedom. The orange cross is the best fit point. The shown contours only show statistical uncertainty, no systematics treatment is included here.



**Figure 3:** The  $n_k$ - $\gamma_k$  likelihood contours of the left and right arm of the galaxy (Segment 2 and 3) in the 3-segments generic segmentation scheme, visualized in Figure 1a. The black 68 % and 95 % contours are drawn assuming Wilks’ theorem assuming 2 degrees of freedom, the orange 68 % contour is drawn assuming Wilks’ theorem with 1 degree of freedom. The orange cross is the best fit point. The shown contours only show statistical uncertainty, no systematics treatment is included here.

has the advantage that it does not rely on modeling of the atmospheric neutrino flux, which comes with a large uncertainty. Measuring a spectral index of the galactic neutrino emission enables to enhance the comparison between gamma ray prediction or direct results and in turn helps understanding the origin of the galactic neutrino flux better. It is a crucial step in differentiating hadronic and leptonic emission from gamma rays and allows estimating contributions from neutrino sources in the GP.

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