
















# Time-integrated Southern-sky Neutrino Source Searches with 10 yr of IceCube Starting-track Events at Energies Down to 1 TeV

R. Abbasi<sup>1</sup>, M. Ackermann<sup>2</sup>, J. Adams<sup>3</sup>, S. K. Agarwalla<sup>4,66</sup>, J. A. Aguilar<sup>5</sup>, M. Ahlers<sup>6</sup>, J.M. Alameddine<sup>7</sup>, N. M. Amin<sup>8</sup>, K. Andeen<sup>9</sup>, C. Argüelles<sup>10</sup>, Y. Ashida<sup>11</sup>, S. Athanasiadou<sup>2</sup>, S. N. Axani<sup>8</sup>, R. Babu<sup>12</sup>, X. Bai<sup>13</sup>, A. Balagopal V.<sup>4</sup>, M. Baricevic<sup>4</sup>, S. W. Barwick<sup>14</sup>, S. Bash<sup>15</sup>, V. Basu<sup>4</sup>, R. Bay<sup>16</sup>, J. J. Beatty<sup>17,18</sup>, J. Becker Tjus<sup>19,67</sup>, J. Beise<sup>20</sup>, C. Bellenghi<sup>15</sup>, S. BenZvi<sup>21</sup>, D. Berley<sup>22</sup>, E. Bernardini<sup>23</sup>, D. Z. Besson<sup>24</sup>, E. Blaufuss<sup>22</sup>, L. Bloom<sup>25</sup>, S. Blot<sup>2</sup>, F. Bontempo<sup>26</sup>, J. Y. Book Motzkin<sup>10</sup>, C. Boscolo Meneguolo<sup>23</sup>, S. Böser<sup>27</sup>, O. Botner<sup>20</sup>, J. Böttcher<sup>28</sup>, J. Braun<sup>4</sup>, B. Brinson<sup>29</sup>, Z. Brissou-Tsavoussis<sup>30</sup>, J. Brostean-Kaiser<sup>2</sup>, L. Brusa<sup>28</sup>, R. T. Burley<sup>31</sup>, D. Butterfield<sup>4</sup>, M. A. Campana<sup>32</sup>, I. Caracas<sup>27</sup>, K. Carloni<sup>10</sup>, J. Carpio<sup>33,34</sup>, S. Chattopadhyay<sup>4,66</sup>, N. Chau<sup>5</sup>, Z. Chen<sup>35</sup>, D. Chirkin<sup>4</sup>, S. Choi<sup>36,37</sup>, B. A. Clark<sup>22</sup>, A. Coleman<sup>20</sup>, P. Coleman<sup>28</sup>, G. H. Collin<sup>38</sup>, A. Connolly<sup>17,18</sup>, J. M. Conrad<sup>38</sup>, R. Corley<sup>11</sup>, D. F. Cowen<sup>39,40</sup>, C. De Clercq<sup>41</sup>, J. J. DeLaunay<sup>25</sup>, D. Delgado<sup>10</sup>, S. Deng<sup>28</sup>, A. Desai<sup>4</sup>, P. Desiati<sup>4</sup>, K. D. de Vries<sup>41</sup>, G. de Wasseige<sup>42</sup>, T. DeYoung<sup>12</sup>, A. Diaz<sup>38</sup>, J. C. Díaz-Vélez<sup>4</sup>, P. Dierichs<sup>28</sup>, M. Dittmer<sup>43</sup>, A. Domi<sup>44</sup>, L. Draper<sup>11</sup>, H. Dujmovic<sup>4</sup>, D. Durnford<sup>45</sup>, K. Dutta<sup>27</sup>, M. A. DuVernois<sup>4</sup>, T. Ehrhardt<sup>27</sup>, L. Eidenschink<sup>15</sup>, A. Eimer<sup>44</sup>, P. Eller<sup>15</sup>, E. Ellinger<sup>46</sup>, S. El Mentawi<sup>28</sup>, D. Elsässer<sup>7</sup>, R. Engel<sup>26,47</sup>, H. Erpenbeck<sup>4</sup>, W. Esmail<sup>43</sup>, J. Evans<sup>22</sup>, P. A. Evenson<sup>8</sup>, K. L. Fan<sup>22</sup>, K. Fang<sup>4</sup>, K. Farrag<sup>48</sup>, A. R. Fazely<sup>49</sup>, A. Fedynitch<sup>50</sup>, N. Feigl<sup>51</sup>, S. Fiedlschuster<sup>44</sup>, C. Finley<sup>52</sup>, L. Fischer<sup>2</sup>, D. Fox<sup>39</sup>, A. Franczkowiak<sup>19</sup>, S. Fukami<sup>2</sup>, P. Fürst<sup>28</sup>, J. Gallagher<sup>53</sup>, E. Ganster<sup>28</sup>, A. Garcia<sup>10</sup>, M. Garcia<sup>8</sup>, G. Garg<sup>4,66</sup>, E. Genton<sup>10,42</sup>, L. Gerhardt<sup>54</sup>, A. Ghadimi<sup>25</sup>, C. Girard-Carillo<sup>27</sup>, C. Glaser<sup>20</sup>, T. Glüsenskamp<sup>20,44</sup>, J. G. Gonzalez<sup>8</sup>, S. Goswami<sup>33,34</sup>, A. Granados<sup>12</sup>, D. Grant<sup>55</sup>, S. J. Gray<sup>22</sup>, S. Griffin<sup>4</sup>, S. Griswold<sup>21</sup>, K. M. Groth<sup>6</sup>, D. Guevel<sup>4</sup>, C. Günther<sup>28</sup>, P. Gutjahr<sup>7</sup>, C. Ha<sup>56</sup>, C. Haack<sup>44</sup>, A. Hallgren<sup>20</sup>, L. Halve<sup>28</sup>, F. Halzen<sup>4</sup>, L. Hamacher<sup>28</sup>, H. Hamdaoui<sup>35</sup>, M. Ha Minh<sup>15</sup>, M. Handt<sup>28</sup>, K. Hanson<sup>4</sup>, J. Hardin<sup>38</sup>, A. A. Harnisch<sup>12</sup>, P. Hatch<sup>30</sup>, A. Haungs<sup>26</sup>, J. Häusser<sup>28</sup>, K. Helbing<sup>46</sup>, J. Hellrung<sup>19</sup>, J. Hermannsgabner<sup>28</sup>, L. Heuermann<sup>28</sup>, N. Heyer<sup>20</sup>, S. Hickford<sup>46</sup>, A. Hidvegi<sup>52</sup>, C. Hill<sup>48</sup>, G. C. Hill<sup>31</sup>, R. Hmaid<sup>48</sup>, K. D. Hoffman<sup>22</sup>, S. Hori<sup>4</sup>, K. Hoshina<sup>4,68</sup>, M. Hostert<sup>10</sup>, W. Hou<sup>26</sup>, T. Huber<sup>26</sup>, K. Hultqvist<sup>52</sup>, M. Hünnefeld<sup>4</sup>, R. Hussain<sup>4</sup>, K. Hymon<sup>7,50</sup>, A. Ishihara<sup>48</sup>, W. Iwakiri<sup>48</sup>, M. Jacquart<sup>4</sup>, S. Jain<sup>4</sup>, O. Janik<sup>44</sup>, M. Jansson<sup>36</sup>, M. Jeong<sup>11</sup>, M. Jin<sup>10</sup>, B. J. P. Jones<sup>57</sup>, N. Kamp<sup>10</sup>, D. Kang<sup>26</sup>, W. Kang<sup>36</sup>, X. Kang<sup>32</sup>, A. Kappes<sup>43</sup>, D. Kappesser<sup>27</sup>, L. Kardum<sup>7</sup>, T. Karg<sup>2</sup>, M. Karl<sup>15</sup>, A. Karle<sup>4</sup>, A. Katil<sup>45</sup>, U. Katz<sup>44</sup>, M. Kauer<sup>4</sup>, J. L. Kelley<sup>4</sup>, M. Khanal<sup>11</sup>, A. Khattee Zathul<sup>4</sup>, A. Kheirandish<sup>33,34</sup>, J. Kiryluk<sup>35</sup>, S. R. Klein<sup>16,54</sup>, Y. Kobayashi<sup>48</sup>, A. Kochocki<sup>12</sup>, R. Koirala<sup>8</sup>, H. Kolanoski<sup>51</sup>, T. Kontrimas<sup>15</sup>, L. Köpke<sup>27</sup>, C. Kopper<sup>44</sup>, D. J. Koskinen<sup>6</sup>, P. Koundal<sup>8</sup>, M. Kowalski<sup>2,51</sup>, T. Kozynets<sup>6</sup>, N. Krieger<sup>19</sup>, J. Krishnamoorthi<sup>4,66</sup>, T. Krishnan<sup>10</sup>, K. Kruijswijk<sup>42</sup>, E. Krupczak<sup>12</sup>, A. Kumar<sup>2</sup>, E. Kun<sup>19</sup>, N. Kurahashi<sup>32</sup>, N. Lad<sup>2</sup>, C. Lagunas Gualda<sup>15</sup>, M. Lamoureux<sup>42</sup>, M. J. Larson<sup>22</sup>, F. Lauber<sup>46</sup>, J. P. Lazar<sup>42</sup>, K. Leonard DeHolton<sup>40</sup>, A. Leszczyńska<sup>8</sup>, J. Liao<sup>29</sup>, M. Lincetto<sup>19</sup>, Y. T. Liu<sup>40</sup>, M. Liubarska<sup>45</sup>, C. Love<sup>32</sup>, L. Lu<sup>4</sup>, F. Lucarelli<sup>58</sup>, W. Luszczyk<sup>17,18</sup>, Y. Lyu<sup>16,54</sup>, J. Madsen<sup>4</sup>, E. Magnus<sup>41</sup>, K. B. M. Mahn<sup>12</sup>, Y. Makino<sup>4</sup>, E. Manao<sup>15</sup>, S. Mancina<sup>23</sup>, A. Mand<sup>4</sup>, W. Marie Sainte<sup>4</sup>, I. C. Mariş<sup>5</sup>, S. Marka<sup>59</sup>, Z. Marka<sup>59</sup>, M. Marsee<sup>25</sup>, I. Martinez-Soler<sup>10</sup>, R. Maruyama<sup>60</sup>, F. Mayhew<sup>12</sup>, F. McNally<sup>61</sup>, J. V. Mead<sup>6</sup>, K. Meagher<sup>4</sup>, S. Mechbal<sup>2</sup>, A. Medina<sup>18</sup>, M. Meier<sup>48</sup>, Y. Merckx<sup>41</sup>, L. Merten<sup>19</sup>, J. Mitchell<sup>49</sup>, T. Montaruli<sup>58</sup>, R. W. Moore<sup>45</sup>, Y. Morii<sup>48</sup>, R. Morse<sup>4</sup>, M. Moulai<sup>4</sup>, T. Mukherjee<sup>26</sup>, R. Naab<sup>2</sup>, M. Nakos<sup>4</sup>, U. Naumann<sup>46</sup>, J. Necker<sup>2</sup>, A. Negi<sup>57</sup>, L. Neste<sup>52</sup>, M. Neumann<sup>43</sup>, H. Niederhausen<sup>12</sup>, M. U. Nisa<sup>12</sup>, K. Noda<sup>48</sup>, A. Noell<sup>28</sup>, A. Novikov<sup>8</sup>, A. Obertacke Pollmann<sup>48</sup>, V. O'Dell<sup>4</sup>, A. Olivás<sup>22</sup>, R. Orsoe<sup>15</sup>, J. Osborn<sup>4</sup>, E. O'Sullivan<sup>20</sup>, V. Palusova<sup>27</sup>, H. Pandya<sup>8</sup>, N. Park<sup>30</sup>, G. K. Parker<sup>57</sup>, V. Parrish<sup>12</sup>, E. N. Paudel<sup>8</sup>, L. Paul<sup>13</sup>, C. Pérez de los Heros<sup>20</sup>, T. Pernice<sup>2</sup>, J. Peterson<sup>4</sup>, A. Pizzuto<sup>4</sup>, M. Plum<sup>13</sup>, A. Pontén<sup>20</sup>, Y. Popovych<sup>27</sup>, M. Prado Rodriguez<sup>4</sup>, B. Pries<sup>12</sup>, R. Procter-Murphy<sup>22</sup>, G. T. Przybylski<sup>54</sup>, L. Pyras<sup>11</sup>, C. Raab<sup>42</sup>, J. Rack-Helleis<sup>27</sup>, N. Rad<sup>2</sup>, M. Ravn<sup>20</sup>, K. Rawlins<sup>62</sup>, Z. Rechav<sup>4</sup>, A. Rehman<sup>8</sup>, E. Resconi<sup>15</sup>, S. Reusch<sup>2</sup>, W. Rhode<sup>7</sup>, B. Riedel<sup>4</sup>, A. Rifaie<sup>46</sup>, E. J. Roberts<sup>31</sup>, S. Robertson<sup>16,54</sup>, S. Rodan<sup>36,37</sup>, M. Rongen<sup>44</sup>, A. Rosted<sup>48</sup>, C. Rott<sup>11,36</sup>, T. Ruhe<sup>7</sup>, L. Ruohan<sup>15</sup>, D. Ryckbosch<sup>63</sup>, I. Safa<sup>4</sup>, J. Saffer<sup>47</sup>, D. Salazar-Gallegos<sup>12</sup>, P. Sampathkumar<sup>26</sup>, A. Sandrock<sup>46</sup>, M. Santander<sup>25</sup>, S. Sarkar<sup>45</sup>, S. Sarkar<sup>64</sup>, J. Savelberg<sup>28</sup>, P. Savina<sup>4</sup>, P. Schaile<sup>15</sup>, M. Schaufel<sup>28</sup>, H. Schieler<sup>26</sup>, S. Schindler<sup>44</sup>, L. Schlickmann<sup>27</sup>, B. Schlüter<sup>43</sup>, F. Schlüter<sup>5</sup>, N. Schmeisser<sup>46</sup>, T. Schmidt<sup>22</sup>, J. Schneider<sup>44</sup>, F. G. Schröder<sup>8,26</sup>, L. Schumacher<sup>44</sup>, S. Schwirn<sup>28</sup>, S. Sclafani<sup>22</sup>, D. Seckel<sup>8</sup>, L. Seen<sup>4</sup>, M. Seikh<sup>24</sup>, M. Seo<sup>36</sup>, S. Seunarine<sup>65</sup>, P. Sevlé Myhr<sup>42</sup>, R. Shah<sup>32</sup>, S. Shefali<sup>47</sup>, N. Shimizu<sup>48</sup>, M. Silva<sup>4</sup>, B. Skrzypek<sup>16</sup>, B. Smithers<sup>57</sup>, R. Snihur<sup>4</sup>, J. Soedingrekso<sup>7</sup>, A. Søgaard<sup>6</sup>, D. Soldin<sup>11</sup>, P. Soldin<sup>28</sup>, G. Sommani<sup>19</sup>, C. Spannfellner<sup>15</sup>, G. M. Spiczak<sup>65</sup>, C. Spiering<sup>2</sup>, J. Stachurska<sup>63</sup>, M. Stamatikos<sup>18</sup>, T. Stanev<sup>8</sup>, T. Stezelberger<sup>54</sup>, T. Stürwald<sup>46</sup>, T. Stuttard<sup>6</sup>, G. W. Sullivan<sup>22</sup>, I. Taboada<sup>29</sup>, S. Ter-Antonyan<sup>49</sup>, A. Terliuk<sup>15</sup>, M. Thiesmeyer<sup>4</sup>, W. G. Thompson<sup>10</sup>, J. Thwaites<sup>4</sup>, S. Tilav<sup>8</sup>, K. Tollefson<sup>12</sup>, C. Tönnis<sup>36</sup>, S. Toscano<sup>5</sup>, D. Tosi<sup>4</sup>, A. Trettin<sup>2</sup>, M. A. Unland Elorrieta<sup>43</sup>, A. K. Upadhyay<sup>4,66</sup>, K. Upshaw<sup>49</sup>, A. Vaidyanathan<sup>9</sup>, N. Valtonen-Mattila<sup>20</sup>, J. Vandenbroucke<sup>4</sup>, N. van Eijndhoven<sup>41</sup>, D. Vannerom<sup>38</sup>, J. van Santen<sup>2</sup>, J. Vara<sup>43</sup>, F. Varsi<sup>47</sup>, J. Veitch-Michaelis<sup>4</sup>, M. Venugopal<sup>26</sup>, M. Vereecken<sup>42</sup>, S. Vergara Carrasco<sup>3</sup>, S. Verpoest<sup>8</sup>, D. Veske<sup>59</sup>, A. Vijal<sup>22</sup>, C. Walck<sup>52</sup>, A. Wang<sup>29</sup>, C. Weaver<sup>12</sup>, P. Weigel<sup>38</sup>, A. Weindl<sup>26</sup>

J. Weldert<sup>40</sup>, A. Y. Wen<sup>10</sup> , C. Wendt<sup>4</sup> , J. Werthebach<sup>7</sup>, M. Weyrauch<sup>26</sup>, N. Whitehorn<sup>12</sup> , C. H. Wiebusch<sup>28</sup> ,  
 D. R. Williams<sup>25</sup>, L. Witthaus<sup>7</sup> , M. Wolf<sup>15</sup> , G. Wrede<sup>44</sup>, X. W. Xu<sup>49</sup>, J. P. Yanez<sup>45</sup>, E. Yildizci<sup>4</sup>, S. Yoshida<sup>48</sup> , R. Young<sup>24</sup>,  
 F. Yu<sup>10</sup> , S. Yu<sup>11</sup> , T. Yuan<sup>4</sup> , A. Zegarelli<sup>19</sup> , S. Zhang<sup>12</sup> , Z. Zhang<sup>35</sup>, P. Zhelnin<sup>10</sup> , P. Zilberman<sup>4</sup>,  
 M. Zimmerman<sup>4</sup>, and IceCube Collaboration

- <sup>1</sup> Department of Physics, Loyola University Chicago, Chicago, IL 60660, USA
- <sup>2</sup> Deutsches Elektronen-Synchrotron DESY, Platanenallee 6, D-15738 Zeuthen, Germany
- <sup>3</sup> Dept. of Physics and Astronomy, University of Canterbury, Private Bag 4800, Christchurch, New Zealand
- <sup>4</sup> Dept. of Physics and Wisconsin IceCube Particle Astrophysics Center, University of Wisconsin—Madison, Madison, WI 53706, USA
- <sup>5</sup> Université Libre de Bruxelles, Science Faculty CP230, B-1050 Brussels, Belgium
- <sup>6</sup> Niels Bohr Institute, University of Copenhagen, DK-2100 Copenhagen, Denmark
- <sup>7</sup> Dept. of Physics, TU Dortmund University, D-44221 Dortmund, Germany
- <sup>8</sup> Bartol Research Institute and Dept. of Physics and Astronomy, University of Delaware, Newark, DE 19716, USA
- <sup>9</sup> Department of Physics, Marquette University, Milwaukee, WI 53201, USA
- <sup>10</sup> Department of Physics and Laboratory for Particle Physics and Cosmology, Harvard University, Cambridge, MA 02138, USA
- <sup>11</sup> Department of Physics and Astronomy, University of Utah, Salt Lake City, UT 84112, USA
- <sup>12</sup> Dept. of Physics and Astronomy, Michigan State University, East Lansing, MI 48824, USA
- <sup>13</sup> Physics Department, South Dakota School of Mines and Technology, Rapid City, SD 57701, USA
- <sup>14</sup> Dept. of Physics and Astronomy, University of California, Irvine, CA 92697, USA
- <sup>15</sup> Physik-department, Technische Universität München, D-85748 Garching, Germany
- <sup>16</sup> Dept. of Physics, University of California, Berkeley, CA 94720, USA
- <sup>17</sup> Dept. of Astronomy, Ohio State University, Columbus, OH 43210, USA
- <sup>18</sup> Dept. of Physics and Center for Cosmology and Astro-Particle Physics, Ohio State University, Columbus, OH 43210, USA
- <sup>19</sup> Fakultät für Physik & Astronomie, Ruhr-Universität Bochum, D-44780 Bochum, Germany
- <sup>20</sup> Dept. of Physics and Astronomy, Uppsala University, Box 516, SE-75120 Uppsala, Sweden
- <sup>21</sup> Dept. of Physics and Astronomy, University of Rochester, Rochester, NY 14627, USA
- <sup>22</sup> Dept. of Physics, University of Maryland, College Park, MD 20742, USA
- <sup>23</sup> Dipartimento di Fisica e Astronomia Galileo Galilei, Università Degli Studi di Padova, I-35122 Padova PD, Italy
- <sup>24</sup> Dept. of Physics and Astronomy, University of Kansas, Lawrence, KS 66045, USA
- <sup>25</sup> Dept. of Physics and Astronomy, University of Alabama, Tuscaloosa, AL 35487, USA
- <sup>26</sup> Karlsruhe Institute of Technology, Institute for Astroparticle Physics, D-76021 Karlsruhe, Germany
- <sup>27</sup> Institute of Physics, University of Mainz, Staudinger Weg 7, D-55099 Mainz, Germany
- <sup>28</sup> III. Physikalisches Institut, RWTH Aachen University, D-52056 Aachen, Germany
- <sup>29</sup> School of Physics and Center for Relativistic Astrophysics, Georgia Institute of Technology, Atlanta, GA 30332, USA
- <sup>30</sup> Dept. of Physics, Engineering Physics, and Astronomy, Queen's University, Kingston, ON K7L 3N6, Canada
- <sup>31</sup> Department of Physics, University of Adelaide, Adelaide, 5005, Australia
- <sup>32</sup> Dept. of Physics, Drexel University, 3141 Chestnut St., Philadelphia, PA 19104, USA
- <sup>33</sup> Department of Physics & Astronomy, University of Nevada, Las Vegas, NV 89154, USA
- <sup>34</sup> Nevada Center for Astrophysics, University of Nevada, Las Vegas, NV 89154, USA
- <sup>35</sup> Dept. of Physics and Astronomy, Stony Brook University, Stony Brook, NY 11794-3800, USA
- <sup>36</sup> Dept. of Physics, Sungkyunkwan University, Suwon 16419, Republic of Korea
- <sup>37</sup> Institute of Basic Science, Sungkyunkwan University, Suwon 16419, Republic of Korea
- <sup>38</sup> Dept. of Physics, Massachusetts Institute of Technology, Cambridge, MA 02139, USA
- <sup>39</sup> Dept. of Astronomy and Astrophysics, Pennsylvania State University, University Park, PA 16802, USA
- <sup>40</sup> Dept. of Physics, Pennsylvania State University, University Park, PA 16802, USA
- <sup>41</sup> Vrije Universiteit Brussel (VUB), Dienst ELEM, B-1050 Brussels, Belgium
- <sup>42</sup> Centre for Cosmology, Particle Physics and Phenomenology - CP3, Université catholique de Louvain, Louvain-la-Neuve, Belgium
- <sup>43</sup> Institut für Kernphysik, Universität Münster, D-48149 Münster, Germany
- <sup>44</sup> Erlangen Centre for Astroparticle Physics, Friedrich-Alexander-Universität Erlangen-Nürnberg, D-91058 Erlangen, Germany
- <sup>45</sup> Dept. of Physics, University of Alberta, Edmonton, Alberta, T6G 2E1, Canada
- <sup>46</sup> Dept. of Physics, University of Wuppertal, D-42119 Wuppertal, Germany
- <sup>47</sup> Karlsruhe Institute of Technology, Institute of Experimental Particle Physics, D-76021 Karlsruhe, Germany
- <sup>48</sup> Dept. of Physics and The International Center for Hadron Astrophysics, Chiba University, Chiba 263-8522, Japan
- <sup>49</sup> Dept. of Physics, Southern University, Baton Rouge, LA 70813, USA
- <sup>50</sup> Institute of Physics, Academia Sinica, Taipei, 11529, Taiwan
- <sup>51</sup> Institut für Physik, Humboldt-Universität zu Berlin, D-12489 Berlin, Germany
- <sup>52</sup> Oskar Klein Centre and Dept. of Physics, Stockholm University, SE-10691 Stockholm, Sweden
- <sup>53</sup> Dept. of Astronomy, University of Wisconsin—Madison, Madison, WI 53706, USA
- <sup>54</sup> Lawrence Berkeley National Laboratory, Berkeley, CA 94720, USA
- <sup>55</sup> Dept. of Physics, Simon Fraser University, Burnaby, BC V5A 1S6, Canada
- <sup>56</sup> Dept. of Physics, Chung-Ang University, Seoul 06974, Republic of Korea
- <sup>57</sup> Dept. of Physics, University of Texas at Arlington, 502 Yates St., Science Hall Rm 108, Box 19059, Arlington, TX 76019, USA
- <sup>58</sup> Département de physique nucléaire et corpusculaire, Université de Genève, CH-1211 Genève, Switzerland
- <sup>59</sup> Columbia Astrophysics and Nevis Laboratories, Columbia University, New York, NY 10027, USA
- <sup>60</sup> Dept. of Physics, Yale University, New Haven, CT 06520, USA
- <sup>61</sup> Department of Physics, Mercer University, Macon, GA 31207-0001, USA
- <sup>62</sup> Dept. of Physics and Astronomy, University of Alaska Anchorage, 3211 Providence Dr., Anchorage, AK 99508, USA
- <sup>63</sup> Dept. of Physics and Astronomy, University of Gent, B-9000 Gent, Belgium
- <sup>64</sup> Dept. of Physics, University of Oxford, Parks Road, Oxford OX1 3PU, UK

<sup>65</sup> Dept. of Physics, University of Wisconsin, River Falls, WI 54022, USA

Received 2025 January 23; revised 2025 November 21; accepted 2025 December 12; published 2026 February 3

## Abstract

In the IceCube Neutrino Observatory, a signal of astrophysical neutrinos is obscured by backgrounds from atmospheric neutrinos and muons produced in cosmic-ray interactions. IceCube event selections used to isolate the astrophysical neutrino signal often focus on the morphology of the light patterns recorded by the detector. The analyses presented here use the new IceCube Enhanced Starting Track Event Selection (ESTES), which identifies events likely generated by muon–neutrino interactions within the detector geometry, focusing on neutrino energies of 1–500 TeV with a median angular resolution of  $1.4^\circ$ . Selecting for starting-track events filters out not only the atmospheric-muon background but also the atmospheric-neutrino background in the southern sky. This improves IceCube’s muon–neutrino sensitivity to southern-sky neutrino sources, especially for Galactic sources that are not expected to produce a substantial flux of neutrinos above 100 TeV. In this work, the ESTES sample was applied for the first time to search for astrophysical sources of neutrinos, including a search for diffuse neutrino emission from the Galactic plane. No significant excesses were identified from any of the analyses; however, constraining limits are set on the hadronic emission from TeV gamma-ray Galactic plane objects and models of the diffuse Galactic plane neutrino flux.

*Unified Astronomy Thesaurus concepts:* [Neutrino astronomy \(1100\)](#); [High energy astrophysics \(739\)](#)

## 1. Introduction

A main aim of the IceCube Neutrino Observatory is to locate the origins of astrophysical neutrinos in our Universe. Neutrino emission from astrophysical sources gives insight into the production mechanisms of high-energy cosmic rays in the source environment. In 2013, IceCube announced its discovery of a flux of astrophysical neutrinos observed on Earth; however, the arrival direction of these astrophysical candidate events was found to be consistent with an isotropic flux (M. G. Aartsen et al. 2013a). In the decade since this discovery, two active galactic nuclei (AGN) have been identified by IceCube as likely cosmic-ray accelerators: TXS 0506+056 and NGC 1068 (M. G. Aartsen et al. 2018a, 2018b; R. Abbasi et al. 2022a). Recently, IceCube published evidence for neutrinos from the Galactic plane (R. Abbasi et al. 2023a). Each of these discoveries used different selections of IceCube events curated to take advantage of the various morphological structures that are created in the light produced by neutrinos interacting in or around the IceCube detector, which instruments a cubic kilometer of deep Antarctic ice. Here, we present the results of the neutrino source searches from a new IceCube event selection: the Enhanced Starting Track Event Selection (ESTES; R. Abbasi et al. 2024). This novel selection improves the IceCube sensitivity to southern-sky neutrino sources by finding astrophysical neutrino candidate events previously missed in other IceCube data selections.

ESTES was developed to select starting-track events created by muon neutrinos, which undergo a charged-current (CC) weak-force interaction inside the IceCube detector geometry, producing a hadronic cascade in addition to the muon track. A final level ESTES starting-track event is shown in Figure 1 to

exemplify the morphological structure in data. Identifying starting events allows for the distinction between astrophysical neutrinos and both atmospheric muons produced in cosmic-ray air showers, which deposit light as they enter into the detector, and atmospheric neutrinos that are accompanied into the detector by muons from the same air shower (S. Schönert et al. 2009). Additionally, observing the interaction point of a muon neutrino also allows the energy reconstruction to estimate the muon energy at its conception and the energy deposited via inelastic scattering into the hadronic cascade, resulting in a more accurate energy reconstruction of the neutrino than for samples that use incoming muon events (R. Abbasi et al. 2024). The final ESTES sample has a higher astrophysical neutrino proportion in the southern sky compared to other track selections, improving IceCube’s sensitivity to some of the most active regions of the Galactic plane, such as the Galactic center, with track events.

The detailed motivation for a new starting-track event selection and a brief description of the selection and reconstruction techniques used by ESTES are provided in Section 2. Four neutrino source searches were performed using ESTES, each testing different hypotheses of neutrino production using a maximum-likelihood approach. The first test, detailed in Section 3, is an all-sky scan that searched for statistically significant clusters of neutrinos using only the ESTES data by testing pixels across the whole sky. The locations of Galactic and extragalactic gamma-ray bright sources were tested individually, as reported in Section 4. Classes of Galactic sources—supernova remnants (SNR), pulsar wind nebulae (PWN), unidentified (UID) TeV objects, and TeV binaries (BIN)—were also evaluated using a stacking analysis described in Section 5. Finally in Section 6, an analysis was performed to look for diffuse neutrino emission from the Galactic plane created by the ambient cosmic-ray flux interacting with Galactic matter using the same template technique as in R. Abbasi et al. (2023a). A discussion of the results is given in Section 7 and conclusions and future prospects of this work are presented in Section 8.

## 2. Starting-track Event Selection

### 2.1. The Starting-track Event Morphology in IceCube

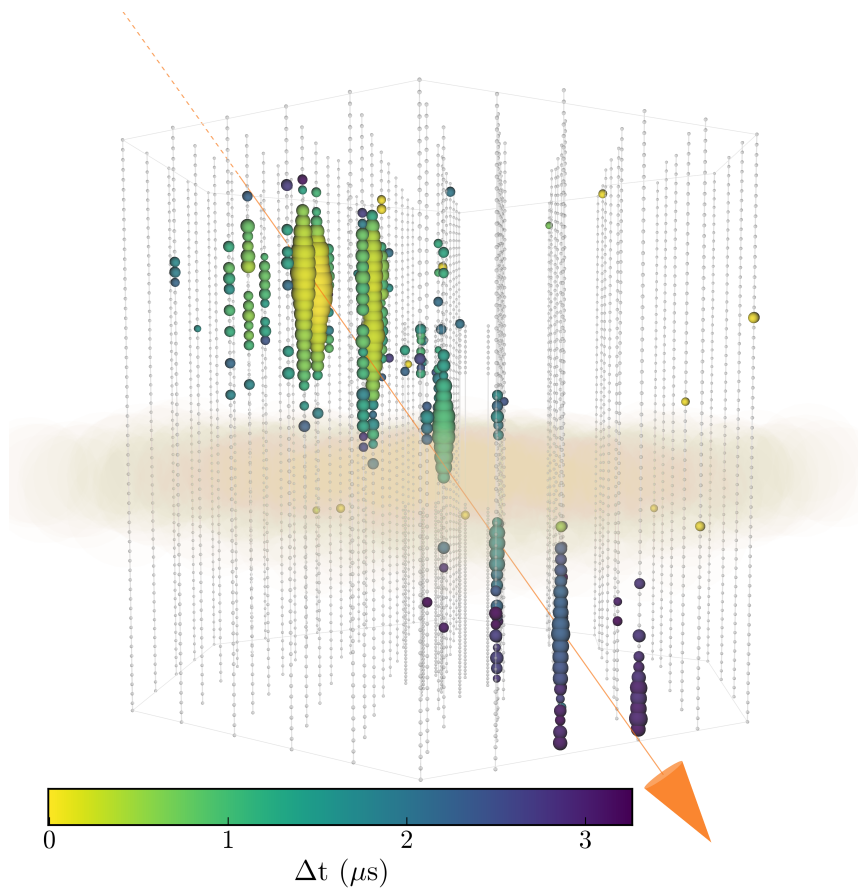
To be observed, neutrinos must interact and produce relativistic charged particles, which deposit light via the

<sup>66</sup> also at Institute of Physics, Sachivalaya Marg, Sainik School Post, Bhubaneswar 751005, India.

<sup>67</sup> also at Department of Space, Earth and Environment, Chalmers University of Technology, 412 96 Gothenburg, Sweden.

<sup>68</sup> also at Earthquake Research Institute, University of Tokyo, Bunkyo, Tokyo 113-0032, Japan.





**Figure 1.** Visualization of a starting-track event from the final level of the enhanced starting-track event selection (ESTES) data. The strings (gray lines) and optical modules (spheres) depict the IceCube detector geometry buried in the glacial ice at the South Pole. The colored spheres represent modules that detected Cherenkov light from the charged particles of the event where the color signifies the relative time, and the size of the sphere represents the relative light intensity seen by the module. In this event, the earliest correlated hits (in yellow) occur on an inner string of the IceCube detector, which is likely around the location of the neutrino interaction that produced a hadronic cascade and muon. The orange arrow shows the reconstructed path of the muon through the detector. The light brown region indicates the dust layer—a region of the ice where scattering and absorption of photons is increased.

Cherenkov effect. The IceCube Neutrino Observatory uses a cubic kilometer of the naturally occurring glacial ice at the South Pole to observe the Cherenkov light from particles created in such neutrino interactions (M. G. Aartsen et al. 2017a). The main IceCube detector is comprised of digital optical modules (DOMs), which house photomultiplier tubes (PMTs) that are sensitive to single photons (M. G. Aartsen et al. 2017a). The detector array consists of 86 strings, each spread approximately 125 m apart in a hexagonal grid and instrumented with 60 DOMs from depths of 1450 to 2450 m below the surface of the glacier (M. G. Aartsen et al. 2017a). Additionally, an array of light-tight tanks, known as IceTop, sits on the surface of the glacier and can detect light from cosmic-ray air showers using PMTs frozen in ice (M. G. Aartsen et al. 2017a).

The pattern of light deposition in the DOMs produced by the secondary charged particles, referred to as the event morphology, is used to infer information about the neutrino. In IceCube, we can classify the morphologies generated by the charged particles into two general categories: track and cascade. Tracks are produced by muons, which come from cosmic-ray air showers, muon–neutrino CC interactions, or tau neutrino CC interactions in which the tau decays to a muon with a 17.3% branching fraction (M. Tanabashi et al. 2018). A muon’s path is relatively unaltered as it passes through the ice.

Muons with an initial energy above 1 TeV can traverse through the whole IceCube detector, leaving a clear track pointing back to their source and resulting in a resolution of the neutrino direction of around  $0.5^\circ$ . The cascade morphology results from several types of neutrino interactions including neutral-current neutrino interactions, electron neutrino CC interactions, and tau neutrino CC interactions where the tau decays into an electron or hadrons. The particle showers produced from these interactions travel relatively short distances leading to a more isotropic distribution of light centered around the interaction vertex. Cascade events are more poorly resolved than track events with resolutions of around  $5^\circ$ – $15^\circ$  (R. Abbasi et al. 2023a). It should be noted that, in the energy range that IceCube is sensitive to, neutrinos interact through deep inelastic scattering meaning the neutrino will impart some energy to the nucleon producing a hadronic cascade even in muon–neutrino CC interactions (M. Tanabashi et al. 2018). Tau neutrinos can produce other unique morphologies; however, those are not relevant in the energy range of focus in this paper ( $\lesssim 100$  TeV).

In this work, we focus on a subclass of the track events, starting tracks (example event shown in Figure 1), which are generated by muon (and tau) neutrino CC interactions where the interaction vertex is contained inside the detector geometry, and light from a hadronic cascade and muon track

is observed. The other track events can be split into two other subclasses: incoming tracks, where a muon is generated outside of IceCube’s array, enters into the instrumented area, and most often leaves the detector, and skimming tracks, where a muon is generated outside of the detector and never enters the geometry but passes close enough to deposit light on the outer strings. For incoming and skimming tracks, the energy reconstruction can only estimate the muon’s energy at its first point of detection, effectively serving as a lower limit of the true neutrino energy. The muon–neutrino energy resolution is therefore improved when using starting tracks because the interaction vertex is observed, allowing the neutrino energy to be inferred from estimates of the muon’s energy at production as well as the energy transferred to the hadronic component (R. Abbasi et al. 2024).

The main motivating factor for a starting-track event selection is atmospheric muon and neutrino background rejection in the southern sky. A simple approach to reject the 3 kHz muon background is to use the Earth as a veto and look only at events coming from the northern sky (R. Abbasi et al. 2022a, 2022b). However, with this method, it is not possible to study the southern sky where many interesting astrophysical objects, such as the Galactic center, lie. The atmospheric-muon rate in the southern sky can be reduced by adding charge cuts to select only the highest-energy events, but the muon background contamination is still significant (M. G. Aartsen et al. 2017b, 2020a). Selecting for events that appear internally within the detector allows the rejection of atmospheric muons, which produce light as they enter the instrumented volume.

The power of using starting events lies in the ability to reject not only atmospheric muons but also atmospheric neutrinos, which are often accompanied into the detector by muons from the same cosmic-ray air shower. The probability of an atmospheric-neutrino event also including light from muons from the same air shower increases with the neutrino’s energy, producing a noticeable effect in the rate of atmospheric neutrino flux above a few TeV (S. Schönert et al. 2009). In C. A. Argüelles et al. (2018), the authors found that conventional atmospheric-muon neutrinos have a 50% chance of being vetoed by accompanying muons at energies of 7 TeV when entering the detector at a zenith angle of  $32^\circ$  relative to the vertical detector direction. This 50% threshold rises to 100 TeV for more inclined atmospheric neutrinos entering at zeniths around  $76^\circ$  as the muons must survive much longer to reach the detector (C. A. Argüelles et al. 2018). At a zenith of  $60^\circ$  (decl. of  $-30^\circ$ ), the atmospheric flux is reduced by a factor of 30 when the atmospheric self-veto is accounted for (R. Abbasi et al. 2024). Creating a specific starting-track sample has many uses, such as performing a fit of the diffuse astrophysical muon–neutrino flux, which is described in R. Abbasi et al. (2024). Here, we use the advantages of starting tracks in IceCube to investigate the southern sky for potential neutrino sources, achieving a higher astrophysical neutrino purity between 1 and 500 TeV through the rejection of atmospheric muons and neutrinos.

## 2.2. Design of the Enhanced Starting Track Event Selection

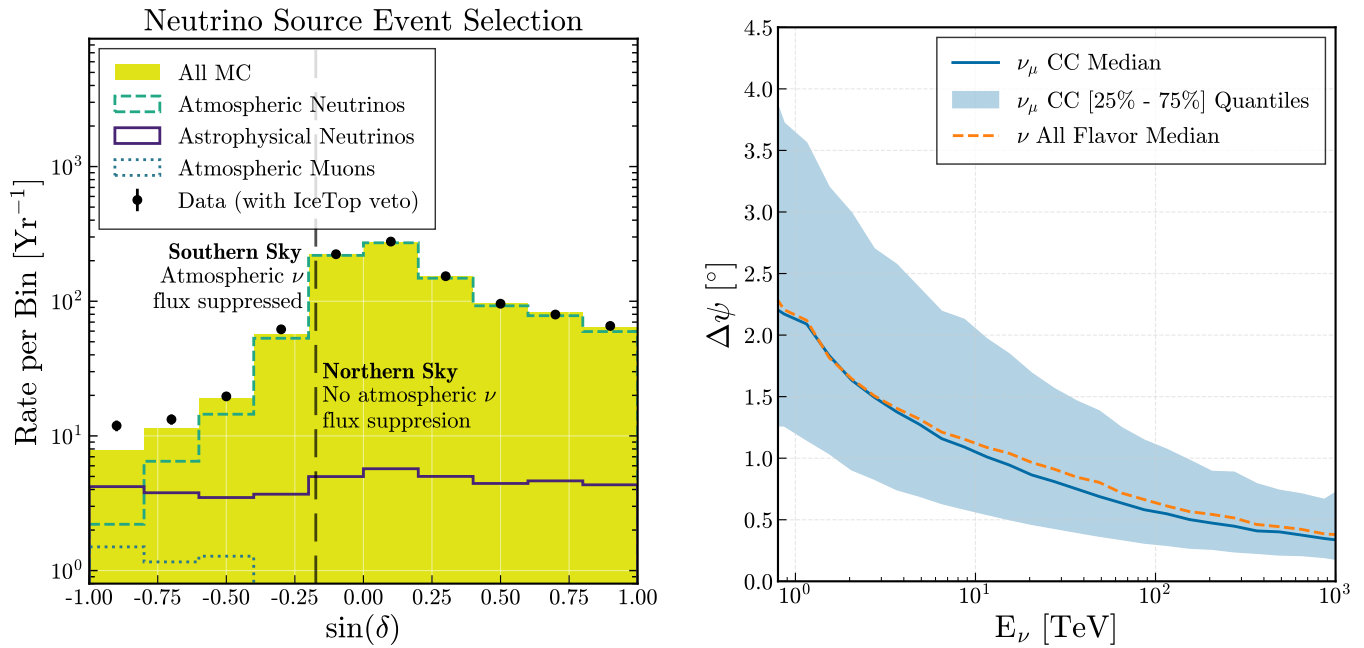
ESTES looks to gather starting-tracks events, particularly in the southern sky with neutrino energies below 100 TeV, where many of IceCube’s astrophysical muon–neutrino events have not yet been selected by other methods. The previous starting event selections, such as the High-Energy Starting Events

(HESE) sample (R. Abbasi et al. 2021a), are optimized to find high-energy cascades by using a fixed layer of in-ice DOMs as a veto, rejecting events where light is observed in the veto layer. In order to use the same fixed-veto-layer method, but lower the energy threshold, the fiducial volume must be substantially restricted as the atmospheric-muon background is able to pass farther through the detector without leaving a signal (M. G. Aartsen et al. 2015).

To identify starting-track events while maximizing the fiducial volume, a new tool was developed to reconstruct the position of the interaction vertex and calculate the probability that the event was an incoming or skimming muon, referred to as the starting-track veto “miss probability” (S. Mancina et al. 2019; R. Abbasi et al. 2024). This tool uses parameterized simulations to estimate the light expected at IceCube DOMs based on a track hypothesis. From the light estimates, the interaction vertex is reconstructed. The “miss probability” quantifies the probability that the DOMs lying along the track before the vertex would have missed light from an incoming or skimming muon. For example, an event that passes through an uninstrumented corridor of the detector will have a larger “miss probability” than one that passes close to the DOMs. Similarly, for an event with a reconstructed vertex near the edge where it enters the detector, the “miss probability” will be much greater than for one where the vertex is located well within the detector. The “miss probability” and other input variables are used to train a boosted decision tree (BDT) to distinguish starting neutrino events from incoming or skimming atmospheric muons described in more detail in R. Abbasi et al. (2024).

An observable unique to this work is the novel angular error estimator, which is required for the unbinned likelihood technique explained in Section 3. This parameter is used to model the spread of the reconstructed directions from the true event direction in the signal probability distribution function (PDF) of the likelihood test statistic (TS). Older IceCube analyses have estimated this parameter by using the shape of the directional reconstruction likelihood around the best-fit point (M. G. Aartsen et al. 2020a); however, we found the traditional estimators to be unreliable for the starting-track morphology. Therefore, we took an approach similar to the one used in R. Abbasi et al. (2022a). We used a LightGBM regressor (G. Ke et al. 2017) to estimate the angular distance between the true and reconstructed directions ( $\Delta\psi$ ) with inputs such as the reconstructed zenith, the reconstructed length, the weighted location of light intensity registered by the DOMs of an event, and the reconstructed cascade and muon-deposited energies. The BDT output was then mapped to an angular error estimate by fitting the  $\Delta\psi$  distribution with a von Mises function for events with a similar BDT output. This estimator also worked well for distinguishing track-like events from cascade-like events.

For this work, a cut was applied to remove the cascade-like events by removing events with an angular error estimate greater than  $10^\circ$ . Removing the cascade-like events improves the bias on the number of signal events fit by the likelihood described in Section 3 by removing the poorly reconstructed events. Finally, the selection removes events where the IceTop surface array measures a voltage in the PMTs greater than or equal to two photoelectron equivalents around the time the reconstructed track would have passed by the surface array. This final IceTop cut adds an additional rejection for events



**Figure 2.** (Left) The final level event rates from data and simulation (MC) versus the sine of the reconstructed decl. of the events of the neutrino source searches version of ESTES. The MC is weighted according to the results of the ESTES diffuse astrophysical flux fit (R. Abbasi et al. 2024). The data is shown as black circles with small error bars. A disagreement between the data and MC was observed and is discussed in Section 2.3. A suppression of the atmospheric-neutrino event rate is seen in the southern sky due to the atmospheric-neutrino veto that occurs when selecting starting events. Additionally, events that see in-time light in the IceTop array at the time the event would have passed through the surface detectors are removed from the data (for the simulation, joint IceCube-IceTop simulation was not used, so the cut could not be applied). (Right) The angular distance between the reconstructed and true direction,  $\Delta\psi$ , for MC events. The blue line represents the median angular distance as a function of true neutrino energy for muon-neutrino charged-current interactions weighted by the astrophysical neutrino flux and the blue-shaded regions contains the region of errors between the 25% and 75% weighted quantiles. The orange-dashed line shows the weighted-median angular error for all neutrino flavors and interaction types contained in the final level of the selection.

likely to be atmospheric in origin and has been implemented in the other southern-sky track samples from IceCube (M. G. Aartsen et al. 2013b, 2020a).

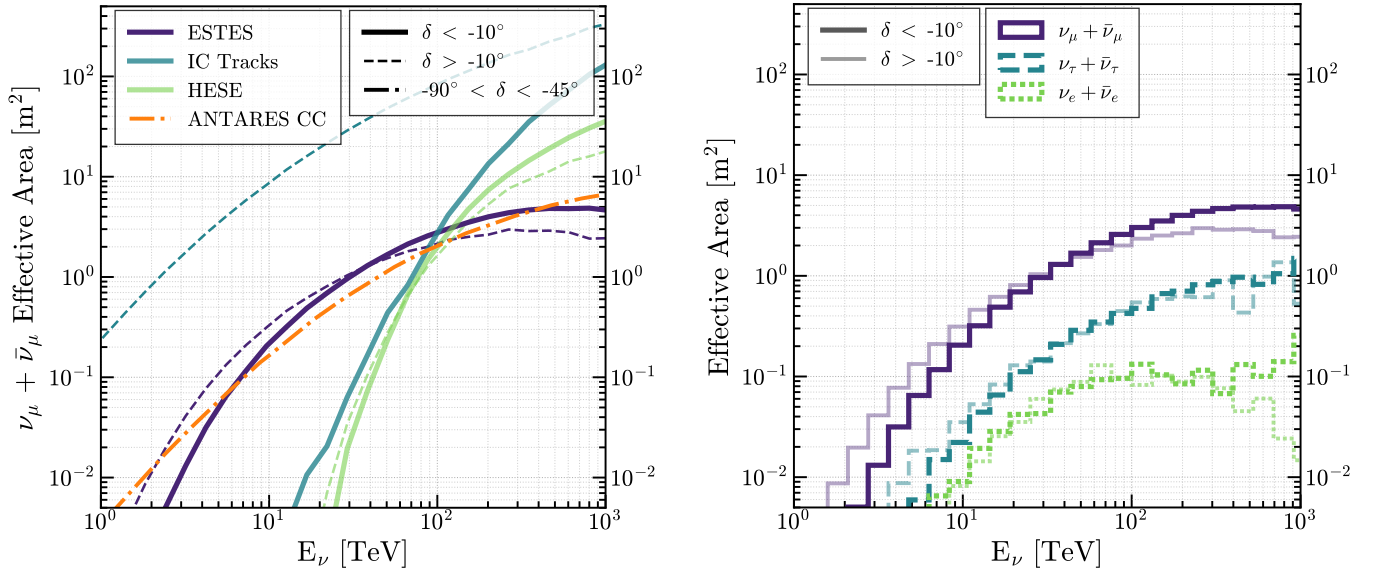
### 2.3. Final ESTES Event Rate and Comparison to Other IceCube Event Selections

The ESTES event selection was applied to 10.3 yr of IceCube data and resulted in a sample of 10,350 total events, with 9061 reconstructed as coming from the northern sky ( $\delta \geq -10^\circ$ ) and 1289 from the southern sky ( $\delta < -10^\circ$ ) where the atmospheric-neutrino veto effect occurs. For most of this work, we differentiate the sky at  $\delta = -10^\circ$  as, around this decl. and for more negative declinations, the atmospheric-neutrino flux is suppressed by the veto effect. The atmospheric-neutrino veto effect can be seen in the left plot of Figure 2, where the atmospheric rate is suppressed for “down-going” events ( $\sin(\delta) < -0.17$  where  $\delta$  is the equatorial decl. of the event). This effect also pushes down the 1:1 atmospheric to astrophysical neutrino crossover point in reconstructed neutrino energy from around 100 TeV to tens of TeV depending on the zenith angle (R. Abbasi et al. 2021b). More detailed discussions of the modeling of the atmospheric-neutrino veto effect are given in R. Abbasi et al. (2024), which derived methods from T. K. Gaissner et al. (2014) and C. A. Argüelles et al. (2018).

From the simulation, around 72 astrophysical neutrinos, 930–950 low-energy atmospheric neutrinos, and 35–60 atmospheric muons were expected in the southern sky assuming the best-fit parameters from the diffuse neutrino flux fit by ESTES in R. Abbasi et al. (2024; ranges given account for simulation statistics). Figure 2 shows the event rate, binned in sine decl.,

and the distribution of the angular distance between the true and reconstructed directions as a function of the true neutrino energy (right) for the final event selection. In the rate plot (left of Figure 2), a disagreement between data and simulation rates is seen in the most southern bins. Studies of the data excess in this region found events that appear by-eye to be clear starting tracks and not incoming or skimming muons. The discrepancy occurs only when the angular error estimate cut is applied, which was not used in R. Abbasi et al. (2024). The cause of this disagreement was traced to differences in the distribution of reconstructions of cascade-like events between the simulation and data, which may be due to incomplete modeling of ice properties and pulse timing. More events are removed by the  $\sigma_{\text{BDT}}$  cut in the simulation than in the data resulting in an underestimation of events. This data–simulation disagreement motivated the use of data-driven background estimation when creating the background PDFs of the point-source likelihood and the background trials used in hypothesis testing. The median angular distance between the true and reconstructed direction for all final level ESTES events is found to be  $1.4^\circ$ . At the final level of the ESTES selection, the energy resolution was found to be within 25%–30% of the true neutrino energy in logarithmic space (R. Abbasi et al. 2024).

To compare ESTES to other samples of astrophysical neutrinos, we include their effective areas in the left plot of Figure 3. The IceCube track sample in M. G. Aartsen et al. (2020a) selects for all track-like events and in the southern sky requires events to be high energy ( $\gtrsim 100$  TeV) to reduce the atmospheric-muon background rate. The effect of this energy cut is the effective area for the track sample drops below the ESTES effective area around 100 TeV for the southern



**Figure 3.** (Left) The muon–neutrino and antineutrino effective area—a measurement of signal capture efficiency—versus the true neutrino energy for the ESTES selection (dark purple) compared to other selections including the IceCube tracks selection (M. G. Aartsen et al. 2020a), the High-Energy Starting Events (HESE; R. Abbasi et al. 2021a), and the ANTARES charged-current (CC) tracks (A. Albert et al. 2017). The different line styles represent different regions of the sky, motivating the focus of the ESTES analysis on the southern-equatorial sky. (Right) The ESTES neutrino and antineutrino effective areas for the three neutrino flavors—distinguished by color and line-style—in two regions of the sky—denoted by line brightness and thickness.

declinations. It should also be noted that the atmospheric-muon background still dominates the other track sample’s astrophysical neutrinos by around 4 orders of magnitude in the southern sky at the final level selection compared to ESTES where the ratio is around 1 order of magnitude due to the atmospheric-neutrino veto (M. G. Aartsen et al. 2017b, 2020a).

A search for overlapping events between ESTES and other IceCube samples was performed to understand how unique the events in the sample were. In the northern sky (events with reconstructed decl.,  $\delta \geq -10^\circ$ ), a large overlap is found with the previous IceCube track samples (M. G. Aartsen et al. 2020a; R. Abbasi et al. 2022a), which select for all track morphology subclasses, as 75% of the ESTES events are also captured in these samples. On the contrary, in the southern sky ( $\delta < -10^\circ$ ), only 1.5% of ESTES events are found in the other track samples. ESTES was found to have little overlap with HESE due to differences in the energy range and incoming-event veto method (R. Abbasi et al. 2021a), with 0.14% of ESTES events found in the HESE sample across the whole sky. Only 2.2% of the ESTES events were in the Deep Neural Network (DNN) cascade sample, responsible for the discovery of an excess of neutrinos from the galactic plane (R. Abbasi et al. 2023a). From the overlap studies, it can be concluded that the ESTES sample provides a complementary and unique selection of relatively high-purity astrophysical muon–neutrino candidates from the southern sky.

### 3. All-sky Search

#### 3.1. Method

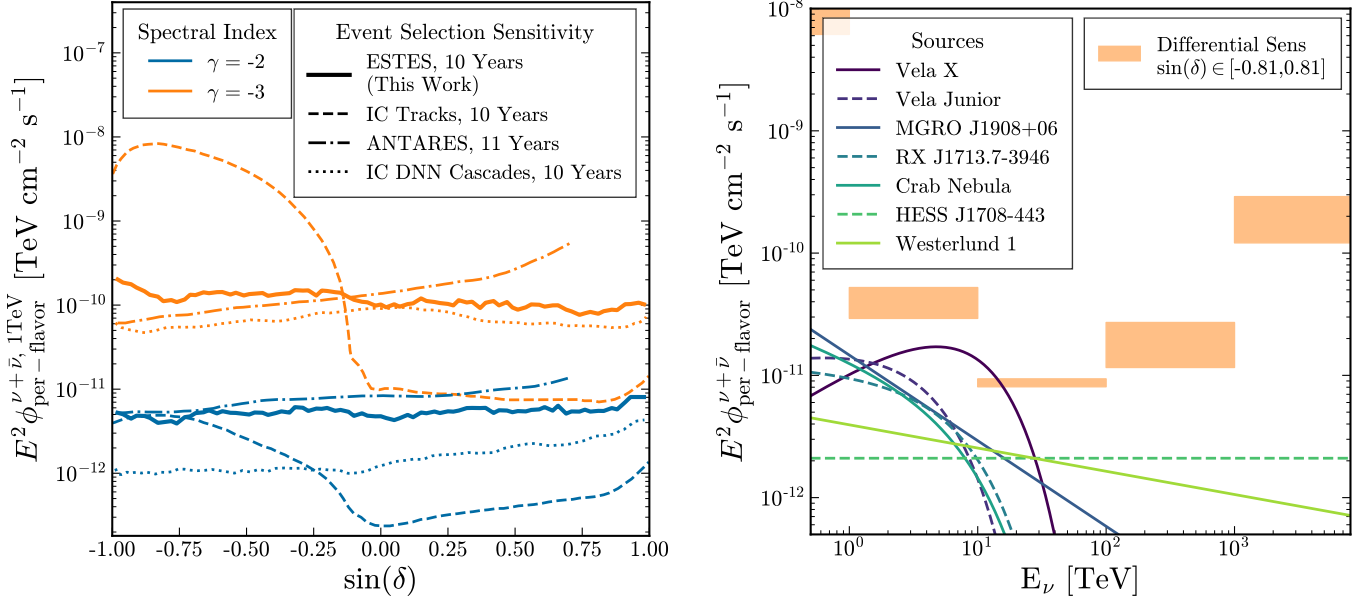
The four analyses presented in this paper all use an unbinned maximum-likelihood method (J. Braun et al. 2008) to generate test statistics (TS) to test against the same null hypothesis: that the ESTES events are produced by a flux isotropic in R.A. (accounting for the decl. dependence of the atmospheric-muon and neutrino backgrounds, the detector, and

the ESTES selection). For point-source searches, the likelihood uses the angular separation between the sample’s events and tested source location as well as the reconstructed energy of the events to estimate the number of signal neutrinos,  $n_s$ , and power-law spectral index,  $\gamma$ , of the hypothesized source, assuming a neutrino flux that follows a single power law:  $\phi(E) = \phi_0 \left(\frac{E}{E_0}\right)^\gamma$ . The likelihood is formulated as

$$\begin{aligned} \mathcal{L}(n_s, \gamma; \mathbf{x}_s, \mathbf{x}_i, \sigma_i, E_i) \\ = \prod_i \left( \frac{n_s}{N} \mathcal{S}(\mathbf{x}_s, \mathbf{x}_i, \sigma_i, E_i, \gamma) + \left(1 - \frac{n_s}{N}\right) \mathcal{B}(\delta_i, E_i) \right) \end{aligned} \quad (1)$$

where  $\mathbf{x}_s$  is the hypothetical source location,  $\vec{x}_i$  is the reconstructed best-fit event locations comprised of declinations and right ascensions ( $\delta_i, \alpha_i$ ),  $\sigma_i$  is the reconstructed angular errors of each event,  $E_i$  is the reconstructed energies of each event, and  $N$  is the total number of events. For the likelihood’s signal PDF,  $\mathcal{S}$ , we used a von Mises distribution to model the point-spread function and used a simulation to estimate the signal energy PDF as a function of decl. The background spatial and energy PDFs, encompassed by  $\mathcal{B}$  in Equation (1), were generated by taking the spline of the data binned in reconstructed decl. and energy.

A TS is generated using the log of the ratio of the likelihood maximized over  $n_s$  and  $\gamma$ ,  $\mathcal{L}(\hat{n}_s, \hat{\gamma})$ , to the likelihood assuming the null hypothesis,  $\mathcal{L}(n_s = 0)$ . Background simulations are produced by randomizing our data events in R.A., under the assumption that most of our data events are from atmospheric-muon and neutrino backgrounds. The background TS distributions are calculated for several declinations ranging from  $-78.5^\circ$  to  $78.5^\circ$ . For the all-sky search, we test uniformly distributed pixels produced using healpy (K. M. Górski et al. 2005; A. Zonca et al. 2019) across the sky, scanning for a pointlike neutrino source. This provides an unbiased search for a pointlike neutrino source; however, it is subject to significant



**Figure 4.** (Left) The sensitivity neutrino and antineutrino per-flavor flux at 1 TeV for a steadily emitting point source versus decl. for the ESTES analysis (solid lines) for two different source spectral index hypotheses ( $\gamma = -2.0$  in blue and  $\gamma = -3.0$  in orange). The ESTES sensitivity is compared to that from the 10 yr IceCube track sample (M. G. Aartsen et al. 2020a), 11 yr ANTARES (J. Aublin et al. 2019), and the IceCube Deep Neural Network (DNN)-selected cascade search (R. Abbasi et al. 2023a). (Right) The maximum and minimum ESTES differential neutrino and antineutrino per-flavor sensitivity flux from simulations produced across a range of declinations (orange bands). The lines represent the models compiled from sources in gamma-cat<sup>69</sup> and TeVCat (S. P. Wakely & D. Horan 2008) converted into a neutrino flux under an optimistic scenario (F. Halzen 2006) for the seven Galactic plane sources chosen in the source-list analysis.

trial corrections due to the nearly 800,000 pixels tested ( $\text{healpy} \text{nside} = 256$  with a mean spacing of 0.23<sup>6</sup>).

To evaluate and compare the performance of ESTES for point-source searches, we produce sensitivities as a function of source location and energy range as shown in Figure 4. The sensitivity is defined as the median expected 90% upper limit flux found in the absence of an observed signal from simulations. In the left plot of Figure 4, the sum of the neutrino and antineutrino per-flavor sensitivity flux at 1 TeV is given for ESTES and other neutrino samples that performed sources in the southern sky. We find that ESTES improves on the track sample (M. G. Aartsen et al. 2020a) in the southern sky for sources with a spectral index more negative than 2, and it is comparable to the ANTARES 11 yr sensitivity (J. Aublin et al. 2019) and the IceCube cascade (R. Abbasi et al. 2023a) samples. The ESTES sample also improves on the previous sensitivity using IceCube tracks for sources with exponential energy cutoffs in the southern sky, which is expected from the effective area comparison in Figure 3.

In the right plot of Figure 4, the differential sensitivity flux for different decades in energy is given assuming a spectral index of  $\gamma = 2.0$  in a limited energy range. The differential sensitivity is represented by an orange band, which gives the minimum and maximum flux values calculated over a range of declinations for each decade across the whole sky. The ESTES sample is most sensitive between 10 and 100 TeV. The sensitivity is compared to models of the neutrino flux from TeV gamma-ray bright Galactic plane sources, which will be discussed further in Section 4.

### 3.2. Results

The resulting pretrial  $p$ -values for each pixel are shown in Figure 5. We highlight the brightest points found in the northern and southern sky at  $(\delta, \alpha) = (+1.79, 195.64)$  and  $(\delta, \alpha) = (-23.97, 155.21)$ , respectively. The northern hotspot

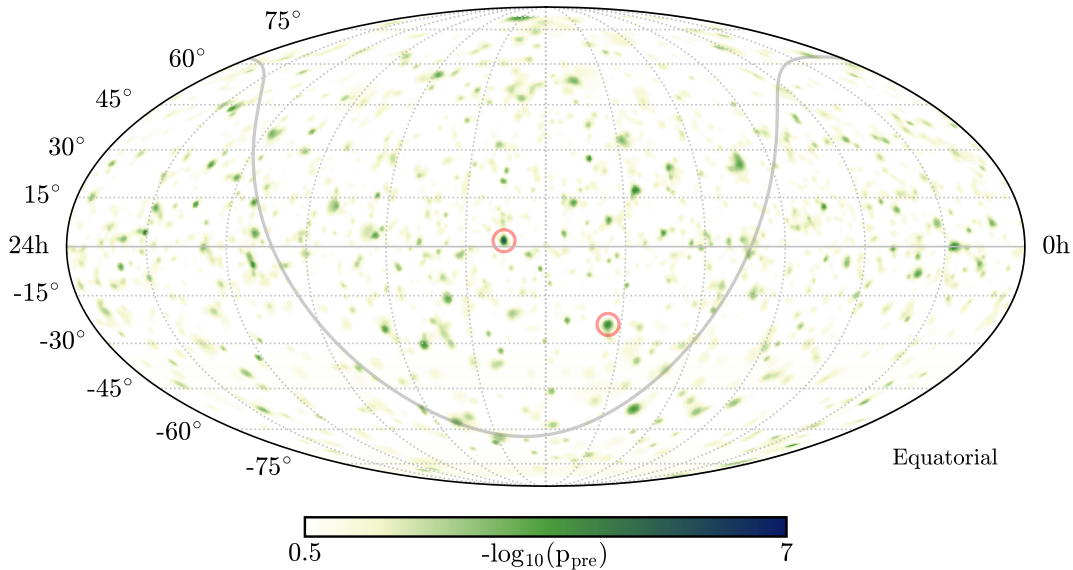
results in the most significant pretrial  $p$ -value of  $1.4 \times 10^{-6}$ . However, after correcting for the number of pixels tested by performing full scans of background simulation maps to account for the angular error radius being larger than the map resolution, the posttrial  $p$ -value is found to be 0.22. Therefore, the null hypothesis cannot be rejected in the all-sky scan analysis.

## 4. Search for Neutrinos from Selected Gamma-Ray Sources

### 4.1. Method

The second analysis performed looked for significant clusters of neutrinos from the direction of known gamma-ray sources. The source selection was based on the hypothesis that astrophysical neutrinos and gamma rays are both produced via pion decay from high-energy cosmic rays interacting with the matter or radiation in the vicinity of the cosmic-ray accelerator. Under this assumption, we selected sources that are bright in gamma rays taking into account the decl. dependence of the ESTES sample, which was the same approach detailed in M. G. Aartsen et al. (2020a). To select extragalactic sources, we used the flux at 1 GeV for sources reported in the Fermi-LAT (Large Area Telescope) 4FGL catalog to select the 103 brightest sources (S. Abdollahi et al. 2020). For Galactic sources, the TeV gamma-ray data compiled in TeVCat (S. P. Wakely & D. Horan 2008) and gamma-cat<sup>69</sup> were converted to a neutrino flux assuming the optimistic scenario that all gamma-ray emission was produced by proton-proton collisions at the source (F. Halzen 2006). The Galactic sources were then compared to the ESTES differential sensitivity at the sources' decl., illustrated in the right plot of Figure 4, and the top seven sources were selected.

<sup>69</sup> gamma-cat: <https://github.com/gammapy/gamma-cat>



**Figure 5.** Map of the ESTES all-sky scan pretrial  $p$ -values,  $p_{\text{pre}}$ . The hottest spots found in the northern and southern equatorial sky are highlighted by the red circles.

The final number of 110 Galactic and extragalactic sources was chosen such that a  $5\sigma$  pretrial significance would drop to a  $4\sigma$  posttrial significance.

All sources were treated as pointlike neutrino emitters in the likelihood evaluation. This assumption has a negligible effect on the sensitivity and discovery potential for all of the sources tested except for Westerlund 1 and Vela Junior (RX J0852.0-4622), which both have measured extensions of around  $1^\circ$  (F. Aharonian et al. 2007; A. Abramowski et al. 2012). Comparing the point-source hypothesis to an extended source hypothesis, we found that the median flux required for a  $3\sigma$  observation using the point-source likelihood was 20% (10%) higher than that of the extended source hypothesis when injecting a  $1.1^\circ$  ( $1.0^\circ$ ) source at the position of Westerlund 1 (Vela Junior). The simulated sources in this test used a two-dimensional Gaussian shape; however, in reality, the shapes of the emission regions for these two sources likely have more complicated morphologies. While treating the sources as pointlike reduces sensitivity, the resulting likelihood outputs and  $p$ -values for these sources are still valid. This approximation, however, affects the upper limits presented in Table 2 in Appendix, which are calculated relative to models derived from the differential gamma-ray flux measurements, which are integrated over the source region (H. Abdalla et al. 2018). For Vela Jr, the upper limit value rises from 1000% of the model flux to 1200% when injecting the extended emission model and fitting with the point-source hypothesis, and for Westerlund 1, it rises from 260% to 360%.

#### 4.2. Results

The most significant source identified by the ESTES sample was 1ES 0647+250, a BL Lac object located in the northern sky. The pretrial correction  $p$ -value is  $4.1 \times 10^{-4}$ , and, after correcting for the 110 source locations tested, the posttrial  $p$ -value was found to be 0.044. Because this source is located in the northern sky, the previous track analyses have a greater sensitivity to 1ES 0647+250 and also tested the location. In the previous analysis with IceCube tracks, no signal neutrinos were found from the direction of the source, and a  $p$ -value of

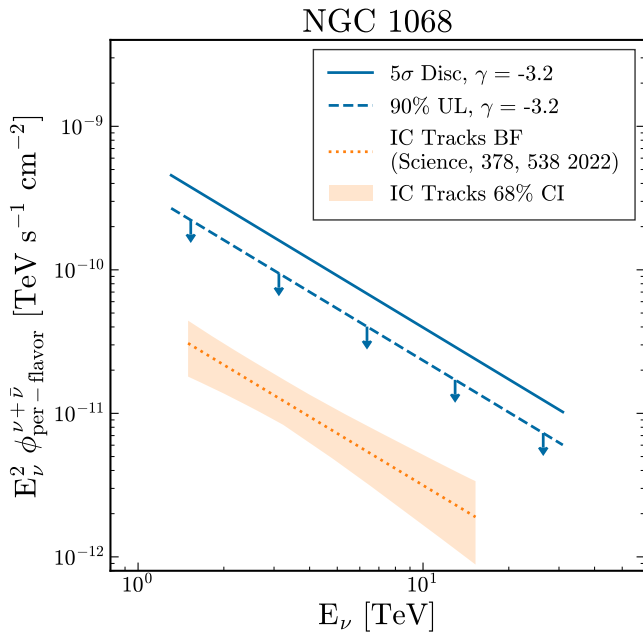
0.53 was reported (M. G. Aartsen et al. 2020a). The previous analysis sets a more constraining 90% upper limit of  $3.0 \times 10^{-13} \text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$  assuming a power-law energy spectrum with  $\gamma = -2$  (M. G. Aartsen et al. 2020a; for this work, the upper limit calculated was  $1.7 \times 10^{-11} \text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$ ).

The full results for all sources are given in Table 4 in Appendix. The most significant ESTES source in the southern region of the sky ( $\delta < -10^\circ$ ), where ESTES provides the most unique events, was the flat spectrum radio quasar, PKS 0727-11, which reports a pretrial  $p$ -value of 0.005. NGC 1068 and TXS 0506+056, two sources with evidence of neutrino emission (M. G. Aartsen et al. 2018a, 2018b; R. Abbasi et al. 2022a), are the 10th and 15th most significant sources in our catalog. Both sources lie above a decl. of  $-10^\circ$ , where the previous track analyses are more sensitive than ESTES (illustrated in the right plot of Figure 4). In Figure 6, the ESTES  $5\sigma$  discovery potential and 90% upper limits assuming the best-fit spectral index for NGC 1068 from R. Abbasi et al. (2022a) analysis are shown to both lie above the 68% confidence interval fit by the track sample. The  $5\sigma$  discovery potential is defined as the median flux required for a discovery at a significance of  $5\sigma$ . Additionally, many northern sky ( $\geq 10^\circ$ ) ESTES events are also captured in the other track samples; in fact, of the eight most significant events contributing to the ESTES NGC 1068 fit, six are also found in the sample from R. Abbasi et al. (2022a). It is concluded that no new neutrino sources were identified by this work.

## 5. Galactic Plane Source Stacking Catalogs

### 5.1. Method

To test the correlation of the ESTES events with specific classes of known sources, a stacking analysis was performed. To “stack” a catalog, the log-likelihood values at the locations of similarly classified sources are summed to produce a single TS. For ESTES, Galactic plane sources were of a particular focus because many lie in the southern sky and have been observed to have energy spectra with power-law indices less than  $-2.0$  and energy cutoffs in the single to tens of TeV when



**Figure 6.** The  $5\sigma$  discovery potential (disk) and 90% upper limit (UL) neutrino and antineutrino per-flavor flux from ESTES, this work, (blue) compared to the best-fit (BF) and confidence interval (CI) fluxes from the NGC 1068 discovery from R. Abbasi et al. (2022a; orange). The energy range shown for this work was selected by finding the central range in which 90% of simulated source events lie assuming the track best-fit power-law spectrum of  $\gamma = -3.2$  (R. Abbasi et al. 2022a). The track energy range was selected by defining the central range where 68% of the top contributing data events lie.

measured in TeV gamma rays (S. P. Wakely & D. Horan 2008). Sources from the TeVCat and gamma-cat<sup>69</sup> databases (S. P. Wakely & D. Horan 2008) were placed into four catalogs: SNR, PWN, TeV BIN, and UID TeV gamma-ray sources. The UID sources are bright sources in TeV gamma rays for which the associated source at other wavelengths has not been concretely identified, but many of the sources are likely to be either SNR or PWN. For the SNR, PWN, and UID catalogs, the top 12 brightest gamma-ray sources of each type were included; whereas, the BIN catalog included all 7 TeV BIN that had been identified at the time.

Although the sources were selected based on gamma-ray brightness, the log-likelihood values evaluated at each source location were summed together without applying any weights to favor the brighter sources, and we fit a single floating spectral index,  $\gamma$ , for each catalog. However, in deriving a sensitivity, discovery potential, and upper limits, a model for the neutrino emission was used for the PWN, SNR, and UID catalogs. The models for each source were derived using the same optimistic approach as outlined for the Galactic sources in Section 4. The parameters for the neutrino flux models from gamma-ray data assuming a proton-proton model for the PWN, SNR, and UID sources are given in Table 2 in Appendix and are shown as the faint orange-dotted lines in Figure 7. The ESTES analysis was found to be slightly sensitive to the PWN and UID catalogs according to these models, but not sensitive to the SNR catalog under the model hypothesis. No models were derived for the BIN catalog because of the time-variability of the objects' flux. Note that, due to the location of the sources of these catalogs, the signal hypothesis tested is correlated with the diffuse Galactic plane neutrino hypothesis test described in Section 6. As discussed in Section 4, the

sources are being treated as pointlike in the likelihood construction; however, here, we found no difference in the sensitivity, discovery potential, or upper limit fluxes when testing the pointlike and extended source hypotheses in the stacking analysis.

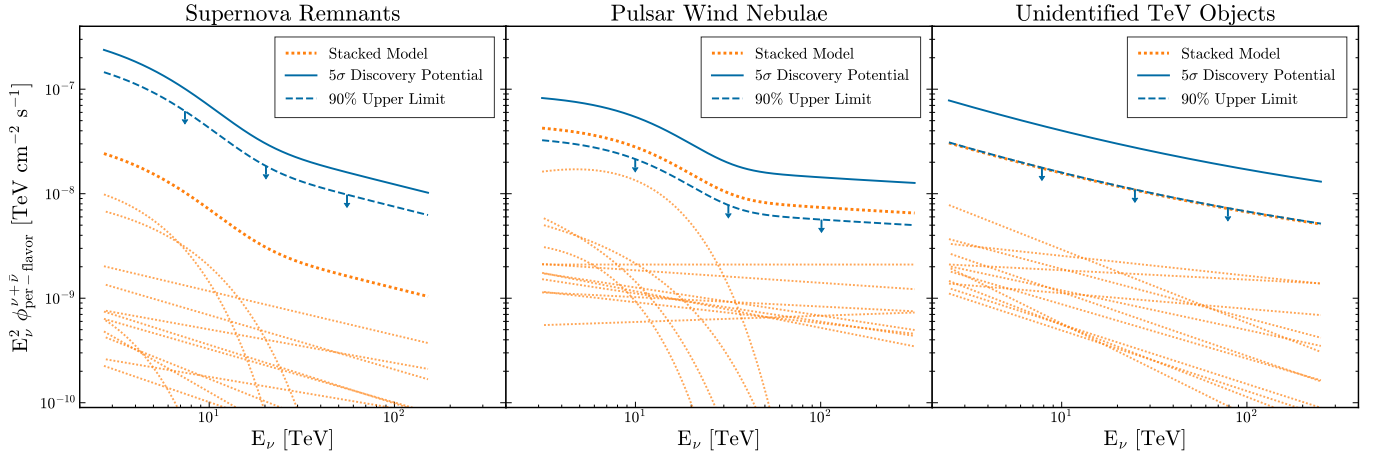
## 5.2. Results

The most significant catalog evaluated in the stacking analysis was the SNR catalog with a pre-trial  $p$ -value of 0.040. To correct for trials, we collected the best-fit TS from background simulations where the four stacking catalog analyses were run on the same map to account for correlations between the fits to each catalog. The post-trial  $p$ -value was found to be 0.056 using this technique; therefore, the null hypothesis could not be rejected. The results and upper limits for all four catalogs are given in Tables 2 and 3 of Appendix. Using the gamma-ray-derived models, the 90% upper limits were calculated for the SNR, PWN, and UID catalogs and are compared to the model expectation fluxes in Figure 7. For the PWN and UID catalog, the 90% upper limits are 77% and 101% of the model flux. For the SNR catalog, the upper limit lies at 601% of the model flux and is much less constraining due to the weaker predicted flux from the models and the fitted result. The BIN catalog returned a TS of 0.0, fitting no signal neutrinos from the direction of the TeV BIN sources, and therefore, competitive upper limits can be set on emission from TeV BIN. Assuming equivalent fluxes from each source and a power-law spectral energy distribution of  $\gamma = -2.0$  and  $-3.0$ , the 90% statistics-only upper limits are, respectively,  $3.0 \times 10^{-12} \text{ TeV}^{-2} \text{ cm}^{-2} \text{ s}^{-1}$  and  $9.2 \times 10^{-13} \text{ TeV}^{-2} \text{ cm}^{-2} \text{ s}^{-1}$ .

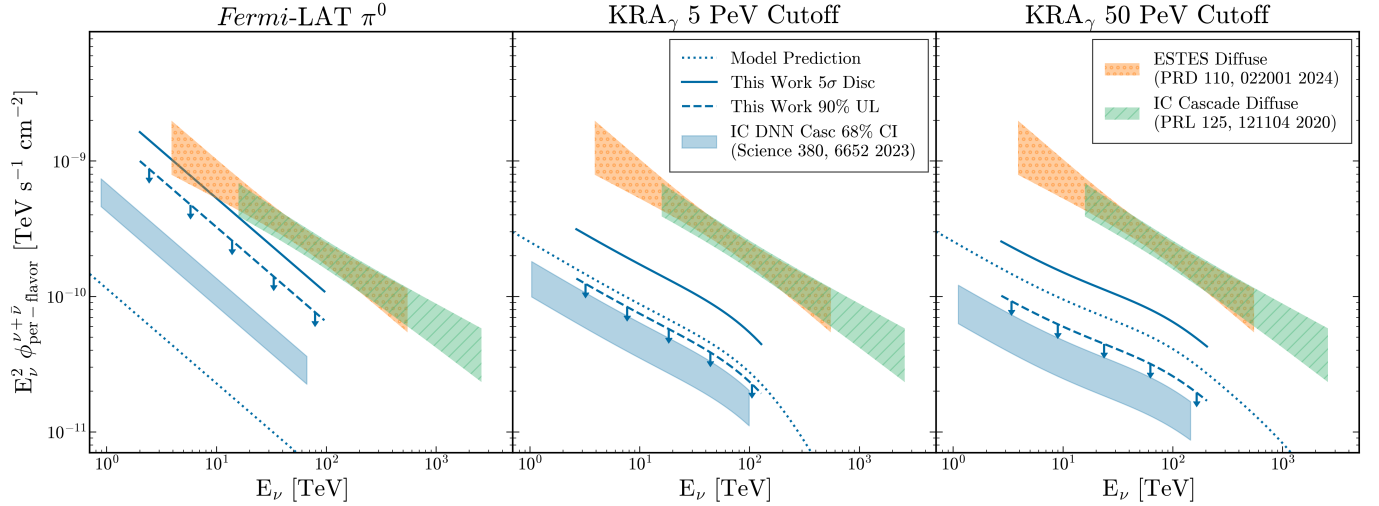
## 6. Diffuse Galactic Plane Neutrino Emission

### 6.1. Method

A flux of cosmic rays permeates our Galaxy and interacts with the interstellar medium (ISM), producing a diffuse flux of neutrinos and gamma rays from the Galactic plane (D. Gaggero et al. 2015; R. Abbasi et al. 2023a). Two models of diffuse Galactic plane neutrinos are tested here: a model extrapolated from the Fermi-LAT  $\pi^0$  model for diffuse galactic gamma rays (M. Ackermann et al. 2012) and the  $\text{KRA}_\gamma$  model (D. Gaggero et al. 2015; P. de la Torre et al. 2022). Both models were produced by injecting a cosmic-ray flux into propagation tools, which interact with the ISM, which is modeled by hydrogen I and CO emission line observations (M. Ackermann et al. 2012). The two models' parameters are fitted using the locally observed cosmic-ray data and are validated by gamma-ray observations (M. Ackermann et al. 2012; D. Gaggero et al. 2015). The  $\text{KRA}_\gamma$  model includes a Galactic radial dependence in the cosmic-ray diffusion coefficient, resulting in an emission intensity map more peaked around the Galactic center than the Fermi-LAT  $\pi^0$  model (D. Gaggero et al. 2015; P. de la Torre et al. 2022). Additionally, a per nucleon exponential energy cutoff is applied to the cosmic-ray spectra in the  $\text{KRA}_\gamma$  model, and the authors produce a prediction for the neutrino spectrum (D. Gaggero et al. 2015; P. de la Torre et al. 2022). Besides the shape differences in the energy spectrum and relative morphological shape of the models, the  $\text{KRA}_\gamma$  model also predicts a larger expected diffuse Galactic plane neutrino flux at TeV energies than the extrapolation of the



**Figure 7.** The  $5\sigma$  discovery potential and 90% statistics-only upper limit neutrino and antineutrino per-flavor fluxes for the stacking catalogs for the three catalogs compared to the model neutrino fluxes (blue lines). The source models (orange, dotted lines) were generated by compiling the TeV gamma-ray data from `gammapcat`<sup>63</sup> and TeVCat (S. P. Wakely & D. Horan 2008) and converting the data into a neutrino flux assuming an optimistic scenario where all TeV gamma rays are produced from pions originating from proton–proton interactions (F. Halzen 2006). The 90% upper limits set for the pulsar wind nebulae rule out 100% of the model flux. The energy range reported represents the central range where 90% of simulated source neutrinos lie assuming the model flux hypothesis.



**Figure 8.** The  $5\sigma$  discovery potential (disk) and 90% statistics-only upper limit (UL) neutrino and antineutrino per-flavor flux for the ESTES analyses compared to the confidence intervals (CI) from the IceCube DNN Cascades (IC DNN Casc) analysis (R. Abbasi et al. 2023a) for the three Galactic plane template models: Fermi-LAT  $\pi^0$  model (M. Ackermann et al. 2012; left) and the  $KRA_\gamma$  (P. de la Torre et al. 2022) with a 5 PeV (middle) and 50 PeV (right) cosmic-ray exponential cutoff. The ESTES UL rule out the  $KRA_\gamma$  model prediction at 90% confidence (P. de la Torre et al. 2022). For added comparison, the astrophysical diffuse fits for ESTES (R. Abbasi et al. 2024) and the diffuse cascade selection (M. G. Aartsen et al. 2020b) are included. The energy range shown for this work is defined by simulating the source hypotheses and finding the central range where 90% of the simulated events lie. The energy range for the DNN Cascades is calculated from the central range where 90% of the contributing data events lie.

Fermi-LAT  $\pi^0$  model (which is observed by comparing the dotted lines in the plots of Figure 8; D. Gaggero et al. 2015; P. de la Torre et al. 2022).

To test these models, the maximum-likelihood approach was again used, but to account for the large size of the emission region, a “signal-subtraction” component was added to remove signal contamination from the background PDFs generated from data. This same likelihood technique and galactic plane models were used in previous searches for Galactic diffuse neutrino emission using different IceCube event selections (M. G. Aartsen et al. 2017c; R. Abbasi et al. 2023a). Unlike the point-source tests described above, the energy spectrum was fixed. In the analysis of the Fermi-LAT  $\pi^0$  model, a single power-law distribution with a spectral index of  $\gamma = -2.7$  was used to describe the energy spectrum. Meanwhile, the likelihood of the  $KRA_\gamma$  model incorporated

the model’s predicted spectra, testing two exponential energy cutoffs applied to the cosmic-ray flux at 5 PeV and at 50 PeV (P. de la Torre et al. 2022). The sensitivity flux for these models is shown for ESTES and compared to previously published results using the same analysis method in Table 1. ESTES is comparable to or more sensitive than all previous analyses using muon track events in the southern sky in IceCube and ANTARES (M. G. Aartsen et al. 2017c; A. Albert et al. 2018; T. Cartraud et al. 2025), but it is less sensitive than the DNN cascade sample, which found an  $4.5\sigma$  excess of neutrinos correlated with the Fermi-LAT  $\pi^0$  model (R. Abbasi et al. 2023a).

## 6.2. Results

The analysis returned a most significant pretrial  $p$ -value of 0.040 for the Fermi-LAT  $\pi^0$  model, and, after correcting for

**Table 1**

Comparison of the ESTES Galactic Plane Results to the IceCube Deep Neural Network (DNN)-selected Cascades (R. Abbasi et al. 2023a), 7 yr of IceCube Tracks (M. G. Aartsen et al. 2017c), the Joint ANTARES and IceCube Track Analysis (A. Albert et al. 2018), and the ANTARES 2023 Results (T. Cartraud et al. 2025)

	Fermi-LAT $\pi^0$ , $\gamma = -2.7$			KRA $_{\gamma}$ 5 PeV			KRA $_{\gamma}$ 50 PeV		
				Cutoff			Cutoff		
			( $p$ )			( $p$ )			( $p$ )
Event Selection	$\phi_{90\%}^{\text{sens}}$	$\phi_{90\%}^{\text{UL}}$	$-\log_{10}(p)$	MF $_{90\%}^{\text{sens}}$	MF $_{90\%}^{\text{UL}}$	$-\log_{10}(p)$	MF $_{90\%}^{\text{sens}}$	MF $_{90\%}^{\text{UL}}$	$-\log_{10}(p)$
ESTES (This Work)	2.6	6.5	1.4	53%	85%	0.78	47%	63%	0.75
DNN Cascades	0.598	...	5.9	16%	...	5.2	11%	...	4.4
IC Tracks, 7 Yr	...	...	...	...	...	...	79%	120%	0.54
ANTARES and IC Tracks	...	...	...	81%	119%	0.54	57%	90%	0.58
ANTARES 2025	2.5	4.2	0.69	70%	150%	1.0	...	...	...

**Note.** The three models tested were the Fermi-LAT  $\pi^0$  model (M. Ackermann et al. 2012) assuming a power-law spectral index of  $\gamma = -2.7$  and the KRA $_{\gamma}$  model (P. de la Torre et al. 2022) assuming a 5 PeV and 50 PeV exponential energy cutoff on the cosmic-ray spectrum. For the Fermi-LAT  $\pi^0$  model, the 90% sensitivity and statistics-only upper limits,  $\phi_{90\%}^{\text{sens}}$  and  $\phi_{90\%}^{\text{UL}}$ , are given as  $E^2 \phi_{\text{per-flavor}}^{\nu+\bar{\nu}}$  at 100 TeV reported in units of  $\times 10^{-11}$  TeV cm $^{-2}$  s $^{-1}$ . For the KRA $_{\gamma}$  models, the sensitivity and upper limits are given as a percentage of the model flux (MF) given by the reference (P. de la Torre et al. 2022). The  $p$ -values reported here are the pretrial's corrected  $p$ -values.

three maps, found a posttrial  $p$ -value of 0.057. Therefore, the null hypothesis cannot be rejected by the ESTES result alone. The preference for the Fermi-LAT  $\pi^0$  model appears to be motivated by the spatial distribution of events as no neutrinos with energies greater than 10 TeV are found in the direction of the Galactic center in the ESTES sample. However, this preference is not statistically significant. The 90% upper limits for ESTES are given in Table 1 and shown in Figure 8. The KRA $_{\gamma}$  90% statistics-only upper limits rule out 100% of the model flux for both choice of cutoff; however, they lie above the 68% confidence intervals fit by the DNN Cascades analysis, and the upper limit calculations do not include modeling of systematic uncertainties (R. Abbasi et al. 2023a).

## 7. Discussion

We performed four analyses to search for sources of astrophysical neutrinos. The all-sky scan, source-list search, Galactic plane source stacking catalogs, and diffuse Galactic plane templates all returned significances of less than  $2\sigma$  when accounting for the trials performed in each of the four analysis, so no new hints of the origin of neutrino sources were identified. A global trials-corrected  $p$ -value was not calculated accounting for the four types of analyses because none of the values were statistically significant after accounting for the trials within each analysis type. However, ESTES was able to provide the most constraining upper limits from IceCube for many Galactic sources of neutrinos.

In 2023, IceCube announced the discovery of Galactic plane neutrinos with the DNN cascade sample with  $4.5\sigma$  posttrial significance (R. Abbasi et al. 2023a). The DNN cascade result had a preference for the Fermi-LAT  $\pi^0$  template as a model for the emission; however, the other models of Galactic plane emission tested—the KRA $_{\gamma}$  template and stacking catalogs—all rejected the null hypothesis with a significance greater than  $3\sigma$ . Discerning between different models of Galactic plane neutrino emission is difficult in this early era of neutrino astronomy, especially for cascade events with a median angular resolution between  $5^\circ$  and  $15^\circ$ . The ESTES sample has an improved resolution of around  $1.4^\circ$  but a smaller

effective area, meaning that, on its own, this sample could not reject the null hypothesis. The 90% statistics-only upper limits presented here limit the KRA $_{\gamma}$  to be less than 100% of the predicted model flux for both cutoff energy hypotheses. The ESTES results also rule out an optimistic scenario for neutrino production from PWN assuming that all TeV gamma-ray emission is from pions created in proton–proton interactions. For ESTES, the Fermi-LAT  $\pi^0$  and the SNR analyses returned similar  $p$ -values, which are both consistent with the background-only hypothesis. One of the next steps in Galactic neutrino astronomy will be disentangling diffuse Galactic plane neutrinos from neutrinos produced at Galactic accelerators.

A near-future solution that can better resolve the Galactic plane is to combine the multiple event selection streams within IceCube to take advantage of all the astrophysical neutrino candidate events recorded by the detector. An analysis has been proposed (P. Savina et al. 2023; M. Thiesmeyer et al. 2025), combining the ESTES starting tracks with the DNN cascade events (R. Abbasi et al. 2023a) and the Earth through-going track events from a track sample (R. Abbasi et al. 2022a). The track sample was used to search for galactic neutrino excesses using a binned style analysis, which rejected the isotropic null hypothesis at the  $2.7\sigma$  level (P. M. Fuerst et al. 2024). However, this track sample does not have a view of the central Galactic region where most emission is expected nor does it benefit from the atmospheric-neutrino veto effect that ESTES and DNN Cascades experience. Using the combined sample, IceCube can further increase the significance of its detection of Galactic plane neutrinos and potentially provide some insight on the exact origin and energy spectrum of Galactic plane neutrinos. The combined sample can also look for neutrino emission from extragalactic sources as more information is learned from the recent NGC 1068 discovery (R. Abbasi et al. 2022a). Many x-ray bright Seyfert galaxies like NGC 1068 are found in the southern sky, and the combination of the pointing resolution of tracks and event rate of the cascades could help further identify these potential neutrino sources (R. Abbasi et al. 2025).

## 8. Conclusions

ESTES is a new selection of IceCube starting-track events, which provides a unique sample of astrophysical muon–neutrino candidates, focusing on neutrinos with energies between 1 and 500 TeV. The advantages of this event selection come from the improved properties of energy resolution, pointing resolution, and background rejection for starting-track events. In particular, selecting for starting events removes not only atmospheric muons created by cosmic-ray air showers but also atmospheric neutrinos, which are accompanied into the detector by muons from the same shower. This increases the purity of the astrophysical neutrino sample in the southern equatorial sky ( $\delta < -10^\circ$ ), resulting in an atmospheric background rate approximately 1 order of magnitude larger than the astrophysical neutrino rate compared to 4 orders of magnitude in the case of previous track selections (M. G. Aartsen et al. 2017b, 2020a). The ESTES analyses improve upon IceCube’s sensitivity to southern-sky neutrino sources for sources with spectral indices less than  $-2.0$  or energy cutoffs. Although no new sources of astrophysical neutrinos were identified in the analysis of the ESTES data, we provide some of the strongest constraints for potential neutrino sources in the southern sky, including a 90% statistics-only upper limit of 85% of the  $KRA_\gamma$  5 PeV model prediction for the diffuse Galactic plane neutrino emission.

The ESTES sample also has other future uses for both multimessenger astronomy and particle physics. IceCube sends neutrino alerts to the astrophysical community for follow-up for events with a high likelihood of being astrophysical in origin (R. Abbasi et al. 2023b). Due to the atmospheric-neutrino background, most of these events must have energies of around 100 TeV or greater to be sent as an alert. However, with ESTES, this energy threshold can be lowered due to the atmospheric-neutrino veto in the southern sky and the improved energy resolution of starting tracks over incoming events. By sending out starting-track alert events using ESTES, IceCube can provide neutrino alert events with a 50% or greater likelihood of astrophysical origin with energies in tens of TeV. These events are of particular interest for real-time follow-up as transient sources located within our own Galaxy or nearby galaxies, such as tidal disruption events, may not have the power to produce 100 TeV neutrinos due to size limitations. Nearby transients provide a good opportunity for gamma-ray follow-up as the gamma rays are less likely to be attenuated. The ESTES-based real-time alerts would add an estimated 5.4 “gold” events per year to the current IceCube alerts (J. Osborn et al. 2023).

Finally, for particle physics, the ESTES starting-track events provide a sample of over 10,000 muon–neutrino CC interactions from 1 to 500 TeV contained inside the detector geometry. These events can be used to measure the inelasticity—the ratio of energy transferred to the hadronic shower to the initial neutrino energy—which gives insight into the atmospheric neutrino-to-antineutrino ratio. Most future proposed cubic kilometer neutrino experiments will be located in the northern hemisphere and should therefore have a good resolution of the southern sky and Galactic center. However, for IceCube-Gen2, developing a starting-track event selection has been considered a goal in order to generate a sample of muon–neutrino events with atmospheric-neutrino suppression for studying the properties of the astrophysical neutrino flux below 100 TeV.

## Acknowledgments

The IceCube collaboration recognizes the significant contributions to this work by Sarah Mancina and Manuel Silva. The authors gratefully acknowledge the support from the following agencies and institutions: USA—U.S. National Science Foundation, Office of Polar Programs, U.S. National Science Foundation, Physics Division, U.S. National Science Foundation, EPSCoR, U.S. National Science Foundation, Office of Advanced Cyberinfrastructure, Wisconsin Alumni Research Foundation, Center for High Throughput Computing (CHTC) at the University of Wisconsin, Madison, Open Science Grid (OSG), Partnership to Advance Throughput Computing (PATH), Advanced Cyberinfrastructure Coordination Ecosystem: Services & Support (ACCESS), Frontera computing project at the Texas Advanced Computing Center, U.S. Department of Energy, National Energy Research Scientific Computing Center, Particle astrophysics research computing center at the University of Maryland, Institute for Cyber-Enabled Research at Michigan State University, Astroparticle physics computational facility at Marquette University, NVIDIA Corporation, and Google Cloud Platform; Belgium—Funds for Scientific Research (FRS-FNRS and FWO), FWO Odysseus and Big Science programmes, and Belgian Federal Science Policy Office (Belspo); Germany—Bundesministerium für Bildung und Forschung (BMBF), Deutsche Forschungsgemeinschaft (DFG), Helmholtz Alliance for Astroparticle Physics (HAP), Initiative and Networking Fund of the Helmholtz Association, Deutsches Elektronen Synchrotron (DESY), and High Performance Computing cluster of the RWTH Aachen; Sweden—Swedish Research Council, Swedish Polar Research Secretariat, Swedish National Infrastructure for Computing (SNIC), and Knut and Alice Wallenberg Foundation; European Union—EGI Advanced Computing for research; Australia—Australian Research Council; Canada—Natural Sciences and Engineering Research Council of Canada, Calcul Québec, Compute Ontario, Canada Foundation for Innovation, WestGrid, and Digital Research Alliance of Canada; Denmark—Villum Fonden, Carlsberg Foundation, and European Commission; New Zealand—Marsden Fund; Japan—Japan Society for Promotion of Science (JSPS) and Institute for Global Prominent Research (IGPR) of Chiba University; Korea—National Research Foundation of Korea (NRF); Switzerland—Swiss National Science Foundation (SNSF). This research made use of gamma-cat (<https://github.com/gammapy/gamma-cat>), an open data collection and source catalog for gamma-ray astronomy.

*Facility:* IceCube.

*Software:* Astropy (T. P. Robitaille et al. 2013; Astropy Collaboration et al. 2018, 2022), Healpy (K. M. Górski et al. 2005; A. Zonca et al. 2019), LightGBM (G. Ke et al. 2017).

## Appendix Full Source-list Results and Upper Limits

Here, we present the individual results for each source and catalog tested by the gamma-ray-informed source-list analyses. Tables 2 and 3 list the sources tested in each stacking catalog along with the analysis results for the full catalog and the sources that were tested individually. Table 4 lists all the single sources tested ordered by the log of the pretrial  $p$ -value returned by the analysis.

**Table 2**  
Results from Searches for Galactic Plane Pointlike Sources of Neutrinos

Source Name	$\delta$ (deg)	$\alpha$ (deg)	$\phi_{\nu+\bar{\nu}}^{\text{1TeV}}$	$\gamma$	$E_{\text{cut}}$	TS	$-\log_{10}(p_{\text{pre}})$	90% Sens (%)	90% UL (%)
SNR Catalog	...	...	...	...	...	3.77	1.40	220	600
Vela Junior	-46.33	133.00	$1.4 \times 10^{-11}$	-1.81	3.34	0.00	0.46	1000	1000
RX J1713.7-3946	-39.77	258.36	$9.4 \times 10^{-12}$	-2.06	6.45	0.00	0.61	960	1000
HESS J1614-518	-51.82	243.56	$3.1 \times 10^{-12}$	-2.42	...	...	...	...	...
HESS J1457-593	-59.07	223.70	$2.3 \times 10^{-12}$	-2.52	...	...	...	...	...
SNR G323.7-01.0	-57.20	233.63	$1.1 \times 10^{-12}$	-2.51	...	...	...	...	...
HESS J1731-347	-34.71	262.98	$1.1 \times 10^{-12}$	-2.32	...	...	...	...	...
Gamma Cygni	+40.52	305.27	$1.3 \times 10^{-12}$	-3.02	...	...	...	...	...
RCW 86	-62.65	220.12	$1.1 \times 10^{-12}$	-1.60	1.75	...	...	...	...
HESS J1912+101	+10.19	288.33	$1.3 \times 10^{-12}$	-2.56	...	...	...	...	...
HESS J1745-303	-30.20	266.30	$8.7 \times 10^{-13}$	-2.71	...	...	...	...	...
Cassiopeia A	+58.81	350.85	$4.1 \times 10^{-13}$	-2.61	...	...	...	...	...
CTB 37A	-38.55	258.64	$3.5 \times 10^{-13}$	-2.30	...	...	...	...	...
PWN Catalog	...	...	...	...	...	0.67	0.54	54	77
Vela X	-45.19	128.29	$1.0 \times 10^{-11}$	-1.32	7.00	0.00	0.53	170	170
Crab Nebula	+22.01	83.63	$1.2 \times 10^{-11}$	-2.39	7.15	0.00	0.39	780	820
HESS J1708-443	-44.30	257.00	$2.1 \times 10^{-12}$	-2.00	...	0.00	0.57	250	250
HESS J1825-137	-13.58	276.55	$8.1 \times 10^{-12}$	-2.26	12.40	...	...	...	...
HESS J1632-478	-47.87	248.01	$2.4 \times 10^{-12}$	-2.12	...	...	...	...	...
MSH 15-52	-59.16	228.53	$2.4 \times 10^{-12}$	-2.27	...	...	...	...	...
HESS J1813-178	-17.85	273.36	$1.3 \times 10^{-12}$	-2.09	...	...	...	...	...
HESS J1303-631	-63.20	195.75	$3.1 \times 10^{-12}$	-1.50	3.85	...	...	...	...
HESS J1616-508	-50.91	244.06	$2.6 \times 10^{-12}$	-2.35	...	...	...	...	...
HESS J1420-607	-60.98	214.69	$1.5 \times 10^{-12}$	-2.20	...	...	...	...	...
HESS J1837-069	-6.93	279.43	$2.1 \times 10^{-12}$	-2.27	...	...	...	...	...
HESS J1026-582	-58.29	157.17	$5.2 \times 10^{-13}$	-1.94	...	...	...	...	...
UID Catalog	...	...	...	...	...	0.29	0.51	75	100
MGRO J1908+06	+6.32	286.91	$1.5 \times 10^{-11}$	-2.70	...	0.00	0.36	330	350
Westerlund 1	-45.80	251.50	$3.9 \times 10^{-12}$	-2.19	...	0.00	0.49	260	260
HESS J1702-420	-42.02	255.68	$2.3 \times 10^{-12}$	-2.09	...	...	...	...	...
2HWC J1814-173	-17.31	273.52	$4.7 \times 10^{-12}$	-2.61	...	...	...	...	...
HESS J1841-055	-5.55	280.23	$5.6 \times 10^{-12}$	-2.47	...	...	...	...	...
2HWC J1819-150	-15.06	274.83	$4.3 \times 10^{-12}$	-2.88	...	...	...	...	...
HESS J1804-216	-21.73	271.12	$2.7 \times 10^{-12}$	-2.69	...	...	...	...	...
HESS J1809-193	-19.30	272.63	$2.9 \times 10^{-12}$	-2.38	...	...	...	...	...
HESS J1843-033	-3.30	280.75	$1.6 \times 10^{-12}$	-2.15	...	...	...	...	...
TeV J2032+4130	+41.51	307.93	$2.9 \times 10^{-12}$	-2.52	...	...	...	...	...
HESS J1708-410	-41.09	257.10	$1.9 \times 10^{-12}$	-2.58	...	...	...	...	...
HESS J1857+026	+2.67	284.30	$2.1 \times 10^{-12}$	-2.57	...	...	...	...	...

**Note.** The results for the stacking catalog searches for each source classification are given in the top lines of each table: supernova remnants (SNR), pulsar wind nebulae (PWN), and unidentified TeV Galactic plane objects (UID). Below, the stacking results are the individual sources included in the stacking catalogs with the flux assumptions used to model each source given. The results for the seven individual Galactic plane sources are also given in line with the source information. The flux reported on the left,  $\phi_{\nu+\bar{\nu}}^{\text{1TeV}}$ , is the neutrino and antineutrino per-flavor flux at 1 TeV in units of  $\text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$ . Some source models include an exponential energy cutoff, which is given in units of TeV. The maximized test statistic, TS, and pretrial  $p$ -values are reported. The sensitivity and statistics-only upper limits are given as a percentage of the injected model flux given by the model parameters. All sources were treated as pointlike, and therefore, the limits for Vela Jr. and Westerlund 1 (which are both extended  $1^\circ$ ) become less constraining (1000% to 1200% for Vela Jr., and 260% to 360% for Westerlund 1) when injecting an extended emission model as discussed in Section 4.

**Table 3**  
Results from the TeV Binary Stacking Analysis

Source Name	$\delta$ (deg)	$\alpha$ (deg)	TS	$-\log_{10}(p_{\text{pre}})$	$\phi_{1\text{TeV}}^{\gamma=-2.0}$ 90% UL	$\phi_{1\text{TeV}}^{\gamma=-3.0}$ 90% UL
BIN Catalog	...	...	0.00	0.31	$3.0 \times 10^{-12}$	$9.2 \times 10^{-13}$
LS I +61 303	+61.23	40.13	...	...	...	...
LS 5039	-14.84	276.56	...	...	...	...
PSR B1259-63	-63.84	195.70	...	...	...	...
LMC P3	-67.59	84.00	...	...	...	...
HESS J0632+057	+5.80	98.25	...	...	...	...
Eta Carinae	-59.68	161.27	...	...	...	...
PSR J2032+4127	+41.46	308.05	...	...	...	...
HESS J1018-589 A	-58.93	154.75	...	...	...	...
HESS J1832-093	-9.37	278.19	...	...	...	...

**Note.** This catalog does not include models due to the variability of the gamma-ray flux from binary systems. The sources included in the catalog are listed with their position under the catalog name. The maximized test statistic, TS, and pretrial  $p$ -values are reported. The upper limit fluxes from the analysis are given assuming equal emission from each source and a spectral index of  $\gamma = -2.0$  and  $\gamma = -3.0$ . The statistics-only upper limit results are given as the neutrino and antineutrino per-flavor flux in units of  $\text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$ .

**Table 4**  
Results for Each Source from the Gamma-Ray-influenced Single Source List Given in Order of Pretrial  $p$ -value

Name	Type	$\delta$ (deg)	$-\log_{10}(p_{\text{pre}})$	TS	$\hat{n}_s$	$\hat{\gamma}$	$\phi_{1\text{TeV}}^{\gamma=-2.0}$ 90% UL	$\phi_{1\text{TeV}}^{\gamma=-3.0}$ 90% UL
IES 0647+250	BLL	+25.05	3.39	10.6	11.3	-4.0	$1.7 \times 10^{-11}$	$3.9 \times 10^{-10}$
1H 1720+117	BLL	+11.87	2.32	6.1	1.6	-1.6	$1.3 \times 10^{-11}$	$3.3 \times 10^{-10}$
PKS 0727-11	FSRQ	-11.69	2.31	5.8	6.5	-3.6	$1.3 \times 10^{-11}$	$4.2 \times 10^{-10}$
IES 1959+650	BLL	+65.15	1.89	4.8	5.0	-4.0	$1.7 \times 10^{-11}$	$2.6 \times 10^{-10}$
OT 081	BLL	+9.65	1.78	4.0	6.5	-3.7	$1.1 \times 10^{-11}$	$2.9 \times 10^{-10}$
SMC	GAL	-72.75	1.52	2.8	1.6	-2.1	$7.2 \times 10^{-12}$	$3.2 \times 10^{-10}$
B3 0133+388	BLL	+39.10	1.47	2.9	4.3	-3.2	$1.1 \times 10^{-11}$	$2.2 \times 10^{-10}$
B3 1343+451	FSRQ	+44.88	1.45	3.6	1.8	-2.2	$1.2 \times 10^{-11}$	$2.1 \times 10^{-10}$
PKS 0332-403	BLL	-40.15	1.35	2.0	1.5	-3.1	$7.3 \times 10^{-12}$	$2.2 \times 10^{-10}$
NGC 1068	SBG	-0.01	1.28	3.3	8.1	-4.0	$9.6 \times 10^{-12}$	$2.5 \times 10^{-10}$
PKS 2247-131	BCU	-12.85	1.21	1.9	3.3	-3.2	$9.6 \times 10^{-12}$	$2.8 \times 10^{-10}$
OJ 287	BLL	+20.12	1.12	1.6	3.7	-3.9	$9.5 \times 10^{-12}$	$2.0 \times 10^{-10}$
PKS 0700-661	BLL	-66.18	0.99	1.4	1.1	-2.5	$6.1 \times 10^{-12}$	$2.3 \times 10^{-10}$
B3 0609+413	BLL	+41.37	0.97	1.3	2.1	-2.7	$9.3 \times 10^{-12}$	$1.7 \times 10^{-10}$
TXS 0506+056	BLL	+5.70	0.92	1.3	5.2	-4.0	$8.0 \times 10^{-12}$	$2.0 \times 10^{-10}$
M 82	SBG	+69.67	0.91	1.2	3.0	-3.9	$1.2 \times 10^{-11}$	$1.8 \times 10^{-10}$
Cen B	RDG	-60.45	0.83	1.1	1.9	-3.5	$5.7 \times 10^{-12}$	$1.9 \times 10^{-10}$
PKS 0521-36	AGN	-36.47	0.80	0.4	0.7	-4.0	$6.1 \times 10^{-12}$	$1.7 \times 10^{-10}$
PMN J1918-4111	BLL	-41.19	0.80	0.3	0.7	-2.5	$6.1 \times 10^{-12}$	$1.7 \times 10^{-10}$
PMN J0531-4827	BLL	-48.46	0.76	1.2	1.3	-4.0	$5.8 \times 10^{-12}$	$1.7 \times 10^{-10}$
NGC 3424	SBG	+32.89	0.76	0.9	2.4	...	$8.1 \times 10^{-12}$	$1.6 \times 10^{-10}$
PKS 0208-512	FSRQ	-51.02	0.75	2.0	2.1	...	$6.2 \times 10^{-12}$	$1.8 \times 10^{-10}$
PKS 0235+164	BLL	+16.62	0.75	0.7	3.0	...	$7.8 \times 10^{-12}$	$1.7 \times 10^{-10}$
PKS 0447-439	BLL	-43.84	0.70	0.1	0.4	...	$5.6 \times 10^{-12}$	$1.5 \times 10^{-10}$
PKS 0502+049	FSRQ	+5.00	0.69	0.7	3.3	...	$7.0 \times 10^{-12}$	$1.7 \times 10^{-10}$
PKS 0735+17	BLL	+17.71	0.68	0.5	1.2	...	$7.6 \times 10^{-12}$	$1.6 \times 10^{-10}$
OJ 014	BLL	+1.78	0.66	1.0	4.7	...	$6.8 \times 10^{-12}$	$1.6 \times 10^{-10}$
PKS 2142-75	FSRQ	-75.60	0.66	0.0	0.0	...	$5.0 \times 10^{-12}$	$2.0 \times 10^{-10}$
PKS 1502+106	FSRQ	+10.50	0.66	0.4	2.0	...	$7.2 \times 10^{-12}$	$1.6 \times 10^{-10}$
PKS 1424-41	FSRQ	-42.11	0.61	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.4 \times 10^{-10}$
RX J1713.7-3946	SNR	-39.77	0.61	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 1440-389	BLL	-39.15	0.61	0.0	0.0	...	$5.5 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 0426-380	BLL	-37.94	0.61	0.0	0.0	...	$5.4 \times 10^{-12}$	$1.4 \times 10^{-10}$
Cen A	RDG	-43.02	0.61	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.4 \times 10^{-10}$
MH 2136-428	BLL	-42.59	0.61	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.4 \times 10^{-10}$
4C +55.17	FSRQ	+55.38	0.60	0.9	0.7	...	$8.3 \times 10^{-12}$	$1.3 \times 10^{-10}$
LMC	GAL	-68.75	0.59	0.0	0.0	...	$4.8 \times 10^{-12}$	$1.7 \times 10^{-10}$
1H 1013+498	BLL	+49.43	0.58	0.8	0.7	...	$7.9 \times 10^{-12}$	$1.3 \times 10^{-10}$
4C +01.28	BLL	+1.56	0.58	0.7	2.7	...	$6.6 \times 10^{-12}$	$1.5 \times 10^{-10}$
PKS 0537-441	BLL	-44.09	0.58	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 0301-243	BLL	-24.12	0.58	0.2	0.6	...	$6.0 \times 10^{-12}$	$1.6 \times 10^{-10}$

**Table 4**  
(Continued)

Name	Type	$\delta$ (deg)	$-\log_{10}(P_{\text{pre}})$	TS	$\hat{n}_s$	$\hat{\gamma}$	$\phi_{1\text{TeV}}^{\gamma=-2.0}$ 90% UL	$\phi_{1\text{TeV}}^{\gamma=-3.0}$ 90% UL
3C 279	FSRQ	-5.79	0.58	0.3	1.4	...	$6.4 \times 10^{-12}$	$1.6 \times 10^{-10}$
HESS J1708-443	PWN	-44.30	0.57	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.4 \times 10^{-10}$
PMN J1610-6649	BLL	-66.81	0.57	0.0	0.0	...	$4.7 \times 10^{-12}$	$1.6 \times 10^{-10}$
1H 1914-194	BLL	-19.36	0.54	0.1	0.7	...	$6.2 \times 10^{-12}$	$1.6 \times 10^{-10}$
4C +01.02	FSRQ	+1.58	0.54	0.6	3.1	...	$6.2 \times 10^{-12}$	$1.5 \times 10^{-10}$
Vela X	PWN	-45.19	0.53	0.0	0.0	...	$5.0 \times 10^{-12}$	$1.3 \times 10^{-10}$
Mkn 421	BLL	+38.21	0.52	0.1	1.0	...	$6.8 \times 10^{-12}$	$1.2 \times 10^{-10}$
Arp 299	SBG	+58.52	0.51	0.4	1.8	...	$7.7 \times 10^{-12}$	$1.2 \times 10^{-10}$
Westerlund 1	UID	-45.80	0.49	0.0	0.0	...	$4.8 \times 10^{-12}$	$1.3 \times 10^{-10}$
Ton 599	FSRQ	+29.24	0.49	0.1	0.7	...	$6.6 \times 10^{-12}$	$1.2 \times 10^{-10}$
NGC 1275	RDG	+41.51	0.48	0.1	0.9	...	$6.6 \times 10^{-12}$	$1.2 \times 10^{-10}$
PKS 1936-623	BLL	-62.18	0.48	0.0	0.0	...	$4.4 \times 10^{-12}$	$1.4 \times 10^{-10}$
BL Lac object	BLL	+42.28	0.48	0.1	0.8	...	$6.7 \times 10^{-12}$	$1.1 \times 10^{-10}$
PKS 2155-304	BLL	-30.23	0.47	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 1101-536	BLL	-53.96	0.47	0.6	1.3	...	$4.9 \times 10^{-12}$	$1.4 \times 10^{-10}$
Vela Junior	SNR	-46.33	0.46	0.0	0.0	...	$4.8 \times 10^{-12}$	$1.2 \times 10^{-10}$
B2 0218+357	FSRQ	+35.94	0.45	0.1	0.2	...	$6.4 \times 10^{-12}$	$1.1 \times 10^{-10}$
Galactic Centre	BCU	-29.00	0.45	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 1124-186	FSRQ	-18.96	0.44	0.0	0.0	...	$5.7 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 1830-211	FSRQ	-21.06	0.44	0.0	0.0	...	$5.4 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 0823-223	BLL	-22.51	0.44	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.3 \times 10^{-10}$
PKS 0454-234	FSRQ	-23.41	0.43	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.4 \times 10^{-10}$
TXS 0628-240	BLL	-24.11	0.43	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.3 \times 10^{-10}$
PKS 0118-272	BLL	-27.02	0.43	0.0	0.0	...	$5.1 \times 10^{-12}$	$1.3 \times 10^{-10}$
NGC 253	SBG	-25.29	0.43	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.4 \times 10^{-10}$
S5 0716+71	BLL	+71.34	0.43	0.0	0.4	...	$8.8 \times 10^{-12}$	$1.2 \times 10^{-10}$
PKS 1244-255	FSRQ	-25.80	0.43	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.3 \times 10^{-10}$
PMN J2345-1555	FSRQ	-15.92	0.43	0.0	0.0	...	$5.9 \times 10^{-12}$	$1.4 \times 10^{-10}$
AP Librae	BLL	-24.37	0.43	0.0	0.0	...	$5.1 \times 10^{-12}$	$1.3 \times 10^{-10}$
PKS 0805-07	FSRQ	-7.86	0.43	0.0	0.2	...	$5.6 \times 10^{-12}$	$1.4 \times 10^{-10}$
PKS 1510-089	FSRQ	-9.11	0.42	0.0	0.0	...	$5.6 \times 10^{-12}$	$1.5 \times 10^{-10}$
PKS 0048-09	BLL	-9.49	0.42	0.0	0.0	...	$5.7 \times 10^{-12}$	$1.5 \times 10^{-10}$
PKS 2233-148	BLL	-14.56	0.42	0.0	0.0	...	$5.7 \times 10^{-12}$	$1.5 \times 10^{-10}$
PKS 1730-13	FSRQ	-13.09	0.42	0.0	0.0	...	$5.8 \times 10^{-12}$	$1.5 \times 10^{-10}$
PKS 2023-07	FSRQ	-7.59	0.41	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.3 \times 10^{-10}$
S2 0109+22	BLL	+22.75	0.40	0.0	0.2	...	$5.9 \times 10^{-12}$	$1.1 \times 10^{-10}$
NGC 5380	GAL	+37.50	0.40	0.0	0.0	...	$6.0 \times 10^{-12}$	$1.0 \times 10^{-10}$
RGB J2243+203	BLL	+20.36	0.40	0.0	0.0	...	$5.7 \times 10^{-12}$	$1.1 \times 10^{-10}$
4C +38.41	FSRQ	+38.14	0.39	0.0	0.0	...	$5.9 \times 10^{-12}$	$9.9 \times 10^{-11}$
4C +21.35	FSRQ	+21.38	0.39	0.0	0.0	...	$5.6 \times 10^{-12}$	$1.1 \times 10^{-10}$
Crab Nebula	PWN	+22.01	0.39	0.0	0.0	...	$5.6 \times 10^{-12}$	$1.0 \times 10^{-10}$
MG2 J201534+3710	FSRQ	+37.18	0.39	0.0	0.0	...	$5.9 \times 10^{-12}$	$1.0 \times 10^{-10}$
NGC 2146	SBG	+78.33	0.39	0.0	0.0	...	$8.3 \times 10^{-12}$	$1.1 \times 10^{-10}$
TXS 0518+211	BLL	+21.21	0.39	0.0	0.0	...	$5.5 \times 10^{-12}$	$1.1 \times 10^{-10}$
Mkn 501	BLL	+39.76	0.39	0.0	0.0	...	$5.9 \times 10^{-12}$	$1.0 \times 10^{-10}$
MG1 J021114+1051	BLL	+10.86	0.39	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.1 \times 10^{-10}$
PG 1553+113	BLL	+11.19	0.38	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.1 \times 10^{-10}$
PKS 1424+240	BLL	+23.80	0.38	0.0	0.0	...	$5.6 \times 10^{-12}$	$1.0 \times 10^{-10}$
Arp 220	SBG	+23.53	0.38	0.0	0.0	...	$5.5 \times 10^{-12}$	$1.0 \times 10^{-10}$
ON 246	BLL	+25.30	0.38	0.0	0.0	...	$5.6 \times 10^{-12}$	$1.0 \times 10^{-10}$
TXS 2241+406	FSRQ	+40.96	0.38	0.0	0.1	...	$6.0 \times 10^{-12}$	$1.0 \times 10^{-10}$
CTA 102	FSRQ	+11.73	0.38	0.0	0.0	...	$5.3 \times 10^{-12}$	$1.1 \times 10^{-10}$
M 31	GAL	+41.24	0.37	0.0	0.0	...	$5.7 \times 10^{-12}$	$9.6 \times 10^{-11}$
TXS 0141+268	BLL	+27.09	0.37	0.0	0.0	...	$5.5 \times 10^{-12}$	$1.0 \times 10^{-10}$
IC 678	GAL	+6.63	0.37	0.0	0.0	...	$4.8 \times 10^{-12}$	$1.1 \times 10^{-10}$
B2 1215+30	BLL	+30.12	0.36	0.0	0.0	...	$5.4 \times 10^{-12}$	$1.0 \times 10^{-10}$
MGRO J1908+06	UID	+6.32	0.36	0.0	0.0	...	$4.8 \times 10^{-12}$	$1.0 \times 10^{-10}$
4C +28.07	FSRQ	+28.80	0.36	0.0	0.0	...	$5.4 \times 10^{-12}$	$1.0 \times 10^{-10}$
S4 0814+42	BLL	+42.38	0.36	0.0	0.0	...	$5.5 \times 10^{-12}$	$8.9 \times 10^{-11}$
3C 454.3	FSRQ	+16.15	0.35	0.0	0.0	...	$5.2 \times 10^{-12}$	$1.0 \times 10^{-10}$
B2 1520+31	FSRQ	+31.74	0.35	0.0	0.0	...	$5.3 \times 10^{-12}$	$9.6 \times 10^{-11}$
3C 66A	BLL	+43.04	0.35	0.0	0.0	...	$5.4 \times 10^{-12}$	$8.8 \times 10^{-11}$
PMN J1329-5608	BLL	-56.12	0.34	0.0	0.0	...	$4.0 \times 10^{-12}$	$1.1 \times 10^{-10}$

**Table 4**  
(Continued)










































Name	Type	$\delta$ (deg)	$-\log_{10}(P_{\text{pre}})$	TS	$\hat{n}_s$	$\hat{\gamma}$	$\phi_{1\text{TeV}}^{\hat{\gamma}=-2.0}$ 90% UL	$\phi_{1\text{TeV}}^{\hat{\gamma}=-3.0}$ 90% UL
PMN J1603-4904	BLL	-49.06	0.33	0.0	0.0	...	$4.1 \times 10^{-12}$	$1.1 \times 10^{-10}$
NGC 4945	SBG	-49.47	0.32	0.0	0.0	...	$4.1 \times 10^{-12}$	$1.1 \times 10^{-10}$
PKS 0336-01	FSRQ	-1.78	0.32	0.0	0.7	...	$4.8 \times 10^{-12}$	$1.1 \times 10^{-10}$
GB6 J1542+6129	BLL	+61.50	0.32	0.0	0.0	...	$6.1 \times 10^{-12}$	$8.5 \times 10^{-11}$
PKS 2326-502	FSRQ	-49.93	0.30	0.0	0.0	...	$4.0 \times 10^{-12}$	$1.0 \times 10^{-10}$
PMN J1650-5044	BLL	-50.75	0.29	0.0	0.0	...	$3.9 \times 10^{-12}$	$1.0 \times 10^{-10}$

**Note.** The classifications of each source are given by the catalogs used (S. Abdollahi et al. 2020; S. P. Wakely & D. Horan 2008). The maximized test statistic, TS, and best-fit values for the number of signal neutrinos,  $\hat{n}_s$ , are given for all sources, and the best-fit values for the spectral index,  $\hat{\gamma}$ , are given for the top 20 sources in the list. When maximizing the likelihood, the spectral index allowed it to float between  $-1.0$  and  $-4.0$ . The 90% statistics-only upper limits are reported assuming two different spectral indices,  $\gamma = -2.0$  and  $\gamma = -3.0$ , and are given as the neutrino and antineutrino per-flavor flux upper limit (UL) at 1 TeV in units of  $\text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$ .

**ORCID iDs**

- R. Abbasi <https://orcid.org/0000-0001-6141-4205>  
M. Ackermann <https://orcid.org/0000-0001-8952-588X>  
S. K. Agarwalla <https://orcid.org/0000-0002-9714-8866>  
J. A. Aguilar <https://orcid.org/0000-0003-2252-9514>  
M. Ahlers <https://orcid.org/0000-0003-0709-5631>  
J.M. Alameddine <https://orcid.org/0000-0002-9534-9189>  
K. Andeen <https://orcid.org/0000-0001-9394-0007>  
C. Argüelles <https://orcid.org/0000-0003-4186-4182>  
S. N. Axani <https://orcid.org/0000-0001-8866-3826>  
X. Bai <https://orcid.org/0000-0002-1827-9121>  
A. Balagopal V. <https://orcid.org/0000-0001-5367-8876>  
S. W. Barwick <https://orcid.org/0000-0003-2050-6714>  
V. Basu <https://orcid.org/0000-0002-9528-2009>  
J. J. Beatty <https://orcid.org/0000-0003-0481-4952>  
J. Becker Tjus <https://orcid.org/0000-0002-1748-7367>  
J. Beise <https://orcid.org/0000-0002-7448-4189>  
C. Bellenghi <https://orcid.org/0000-0001-8525-7515>  
S. BenZvi <https://orcid.org/0000-0001-5537-4710>  
E. Bernardini <https://orcid.org/0000-0003-3108-1141>  
E. Blaufuss <https://orcid.org/0000-0001-5450-1757>  
L. Bloom <https://orcid.org/0009-0005-9938-3164>  
S. Blot <https://orcid.org/0000-0003-1089-3001>  
J. Y. Book Motzkin <https://orcid.org/0000-0001-6687-5959>  
C. Boscolo Meneguolo <https://orcid.org/0000-0001-8325-4329>  
S. Böser <https://orcid.org/0000-0002-5918-4890>  
O. Botner <https://orcid.org/0000-0001-8588-7306>  
J. Böttcher <https://orcid.org/0000-0002-3387-4236>  
B. Brinson <https://orcid.org/0000-0001-9128-1159>  
M. A. Campana <https://orcid.org/0000-0003-4162-5739>  
K. Carloni <https://orcid.org/0000-0003-3859-3748>  
J. Carpio <https://orcid.org/0000-0003-0667-6557>  
S. Chattopadhyay <https://orcid.org/0009-0006-1352-2248>  
D. Chirkin <https://orcid.org/0000-0003-4911-1345>  
B. A. Clark <https://orcid.org/0000-0003-4089-2245>  
A. Coleman <https://orcid.org/0000-0003-1510-1712>  
J. M. Conrad <https://orcid.org/0000-0002-6393-0438>  
D. F. Cowen <https://orcid.org/0000-0003-4738-0787>  
C. De Clercq <https://orcid.org/0000-0001-5266-7059>  
J. J. DeLaunay <https://orcid.org/0000-0001-5229-1995>  
D. Delgado <https://orcid.org/0000-0002-4306-8828>  
A. Desai <https://orcid.org/0000-0001-7405-9994>  
P. Desiati <https://orcid.org/0000-0001-9768-1858>  
K. D. de Vries <https://orcid.org/0000-0002-9842-4068>  
G. de Wasseige <https://orcid.org/0000-0002-1010-5100>  
T. DeYoung <https://orcid.org/0000-0003-4873-3783>  
A. Diaz <https://orcid.org/0000-0001-7206-8336>  
J. C. Díaz-Vélez <https://orcid.org/0000-0002-0087-0693>  
H. Dujmovic <https://orcid.org/0000-0003-1891-0718>  
D. Durnford <https://orcid.org/0000-0002-6608-7650>  
M. A. DuVernois <https://orcid.org/0000-0002-2987-9691>  
A. Eimer <https://orcid.org/0009-0002-6308-0258>  
P. Eller <https://orcid.org/0000-0001-6354-5209>  
D. Elsässer <https://orcid.org/0000-0001-6796-3205>  
H. Erpenbeck <https://orcid.org/0000-0001-6319-2108>  
A. R. Fazely <https://orcid.org/0000-0002-6907-8020>  
A. Fedynitch <https://orcid.org/0000-0003-2837-3477>  
C. Finley <https://orcid.org/0000-0003-3350-390X>  
L. Fischer <https://orcid.org/0000-0002-7645-8048>  
D. Fox <https://orcid.org/0000-0002-3714-672X>  
A. Franckowiak <https://orcid.org/0000-0002-5605-2219>  
P. Fürst <https://orcid.org/0000-0002-7951-8042>  
J. Gallagher <https://orcid.org/0000-0001-8608-0408>  
E. Ganster <https://orcid.org/0000-0003-4393-6944>  
A. Garcia <https://orcid.org/0000-0002-8186-2459>  
E. Genton <https://orcid.org/0009-0003-5263-972X>  
A. Ghadimi <https://orcid.org/0000-0002-6350-6485>  
C. Glaser <https://orcid.org/0000-0001-5998-2553>  
T. Glüsenskamp <https://orcid.org/0000-0002-2268-9297>  
S. J. Gray <https://orcid.org/0000-0003-2907-8306>  
S. Griffin <https://orcid.org/0000-0002-0779-9623>  
S. Griswold <https://orcid.org/0000-0002-7321-7513>  
K. M. Groth <https://orcid.org/0000-0002-1581-9049>  
D. Guevel <https://orcid.org/0000-0002-0870-2328>  
C. Günther <https://orcid.org/0009-0007-5644-8559>  
P. Gutjahr <https://orcid.org/0000-0001-7980-7285>  
C. Haack <https://orcid.org/0000-0003-3932-2448>  
A. Hallgren <https://orcid.org/0000-0001-7751-4489>  
L. Halve <https://orcid.org/0000-0003-2237-6714>  
F. Halzen <https://orcid.org/0000-0001-6224-2417>  
H. Hamdaoui <https://orcid.org/0000-0001-5709-2100>  
A. Haungs <https://orcid.org/0000-0002-9638-7574>  
K. Helbing <https://orcid.org/0000-0003-2072-4172>  
J. Hellrung <https://orcid.org/0009-0006-7300-8961>  
N. Heyer <https://orcid.org/0000-0001-9036-8623>  
C. Hill <https://orcid.org/0000-0003-0647-9174>  
S. Hori <https://orcid.org/0009-0007-2644-5955>

- M. Hostert  <https://orcid.org/0000-0002-9584-8877>  
W. Hou  <https://orcid.org/0000-0003-3422-7185>  
T. Huber  <https://orcid.org/0000-0002-6515-1673>  
K. Hultqvist  <https://orcid.org/0000-0003-0602-9472>  
M. Hünnefeld  <https://orcid.org/0000-0002-2827-6522>  
W. Iwakiri  <https://orcid.org/0000-0002-0207-9010>  
S. Jain  <https://orcid.org/0009-0000-7455-782X>  
O. Janik  <https://orcid.org/0009-0007-3121-2486>  
M. Jeong  <https://orcid.org/0000-0003-2420-6639>  
M. Jin  <https://orcid.org/0000-0003-0487-5595>  
B. J. P. Jones  <https://orcid.org/0000-0003-3400-8986>  
N. Kamp  <https://orcid.org/0000-0001-9232-259X>  
D. Kang  <https://orcid.org/0000-0002-5149-9767>  
W. Kang  <https://orcid.org/0000-0003-3980-3778>  
A. Kappes  <https://orcid.org/0000-0003-1315-3711>  
T. Karg  <https://orcid.org/0000-0003-3251-2126>  
M. Karl  <https://orcid.org/0000-0003-2475-8951>  
A. Karle  <https://orcid.org/0000-0001-9889-5161>  
U. Katz  <https://orcid.org/0000-0002-7063-4418>  
M. Kauer  <https://orcid.org/0000-0003-1830-9076>  
J. L. Kelley  <https://orcid.org/0000-0002-0846-4542>  
A. Khathe Zathul  <https://orcid.org/0000-0002-8735-8579>  
A. Kheirandish  <https://orcid.org/0000-0001-7074-0539>  
J. Kiryluk  <https://orcid.org/0000-0003-0264-3133>  
S. R. Klein  <https://orcid.org/0000-0003-2841-6553>  
Y. Kobayashi  <https://orcid.org/0009-0005-5680-6614>  
A. Kochocki  <https://orcid.org/0000-0003-3782-0128>  
R. Koirala  <https://orcid.org/0000-0002-7735-7169>  
H. Kolanoski  <https://orcid.org/0000-0003-0435-2524>  
T. Kontrimas  <https://orcid.org/0000-0001-8585-0933>  
C. Kopper  <https://orcid.org/0000-0001-6288-7637>  
D. J. Koskinen  <https://orcid.org/0000-0002-0514-5917>  
P. Koundal  <https://orcid.org/0000-0002-5917-5230>  
M. Kowalski  <https://orcid.org/0000-0001-8594-8666>  
T. Krishnan  <https://orcid.org/0000-0002-3237-3114>  
K. Kruiswijk  <https://orcid.org/0009-0002-9261-0537>  
A. Kumar  <https://orcid.org/0000-0002-8367-8401>  
N. Kurahashi  <https://orcid.org/0000-0003-1047-8094>  
N. Lad  <https://orcid.org/0000-0001-9302-5140>  
C. Lagunas Gualda  <https://orcid.org/0000-0002-9040-7191>  
M. Lamoureux  <https://orcid.org/0000-0002-8860-5826>  
M. J. Larson  <https://orcid.org/0000-0002-6996-1155>  
F. Lauber  <https://orcid.org/0000-0001-5648-5930>  
J. P. Lazar  <https://orcid.org/0000-0003-0928-5025>  
K. Leonard DeHolton  <https://orcid.org/0000-0002-8795-0601>  
A. Leszczyńska  <https://orcid.org/0000-0003-0935-6313>  
J. Liao  <https://orcid.org/0009-0008-8086-586X>  
M. Lincetto  <https://orcid.org/0000-0002-1460-3369>  
Y. T. Liu  <https://orcid.org/0009-0007-5418-1301>  
L. Lu  <https://orcid.org/0000-0003-3175-7770>  
F. Lucarelli  <https://orcid.org/0000-0002-9558-8788>  
W. Luszczak  <https://orcid.org/0000-0003-3085-0674>  
Y. Lyu  <https://orcid.org/0000-0002-2333-4383>  
J. Madsen  <https://orcid.org/0000-0003-2415-9959>  
E. Magnus  <https://orcid.org/0009-0008-8111-1154>  
E. Manao  <https://orcid.org/0009-0002-6197-8574>  
S. Mancina  <https://orcid.org/0009-0003-9879-3896>  
A. Mand  <https://orcid.org/0009-0005-9697-1702>  
I. C. Mariş  <https://orcid.org/0000-0002-5771-1124>  
S. Marka  <https://orcid.org/0000-0002-3957-1324>  
Z. Marka  <https://orcid.org/0000-0003-1306-5260>  
I. Martinez-Soler  <https://orcid.org/0000-0002-0308-3003>  
R. Maruyama  <https://orcid.org/0000-0003-2794-512X>  
F. Mayhew  <https://orcid.org/0000-0001-7609-403X>  
F. McNally  <https://orcid.org/0000-0002-0785-2244>  
K. Meagher  <https://orcid.org/0000-0003-3967-1533>  
M. Meier  <https://orcid.org/0000-0002-9483-9450>  
L. Merten  <https://orcid.org/0000-0003-1332-9895>  
T. Montaruli  <https://orcid.org/0000-0001-5014-2152>  
R. W. Moore  <https://orcid.org/0000-0003-4160-4700>  
M. Moulai  <https://orcid.org/0000-0001-7909-5812>  
T. Mukherjee  <https://orcid.org/0000-0002-0962-4878>  
R. Naab  <https://orcid.org/0000-0003-2512-466X>  
J. Necker  <https://orcid.org/0000-0003-0280-7484>  
L. Neste  <https://orcid.org/0000-0002-4829-3469>  
H. Niederhausen  <https://orcid.org/0000-0002-9566-4904>  
M. U. Nisa  <https://orcid.org/0000-0002-6859-3944>  
K. Noda  <https://orcid.org/0000-0003-1397-6478>  
A. Obertacke Pollmann  <https://orcid.org/0000-0002-2492-043X>  
V. O'Dell  <https://orcid.org/0000-0003-0903-543X>  
E. O'Sullivan  <https://orcid.org/0000-0003-1882-8802>  
H. Pandya  <https://orcid.org/0000-0002-6138-4808>  
N. Park  <https://orcid.org/0000-0002-4282-736X>  
E. N. Paudel  <https://orcid.org/0000-0001-9276-7994>  
L. Paul  <https://orcid.org/0000-0003-4007-2829>  
C. Pérez de los Heros  <https://orcid.org/0000-0002-2084-5866>  
A. Pizzuto  <https://orcid.org/0000-0002-8466-8168>  
M. Plum  <https://orcid.org/0000-0001-8691-242X>  
B. Pries  <https://orcid.org/0000-0003-4811-9863>  
L. Pyras  <https://orcid.org/0000-0003-1146-9659>  
C. Raab  <https://orcid.org/0000-0001-9921-2668>  
N. Rad  <https://orcid.org/0000-0002-5204-0851>  
A. Rehman  <https://orcid.org/0000-0001-7616-5790>  
E. Resconi  <https://orcid.org/0000-0003-0705-2770>  
W. Rhode  <https://orcid.org/0000-0003-2636-5000>  
B. Riedel  <https://orcid.org/0000-0002-9524-8943>  
M. Rongen  <https://orcid.org/0000-0002-7057-1007>  
A. Rosted  <https://orcid.org/0000-0003-2410-400X>  
C. Rott  <https://orcid.org/0000-0002-6958-6033>  
T. Ruhe  <https://orcid.org/0000-0002-4080-9563>  
I. Safa  <https://orcid.org/0000-0001-8737-6825>  
J. Saffer  <https://orcid.org/0000-0002-0040-6129>  
D. Salazar-Gallegos  <https://orcid.org/0000-0002-9312-9684>  
A. Sandrock  <https://orcid.org/0000-0002-6779-1172>  
M. Santander  <https://orcid.org/0000-0001-7297-8217>  
S. Sarkar  <https://orcid.org/0000-0002-1206-4330>  
S. Sarkar  <https://orcid.org/0000-0002-3542-858X>  
H. Schieler  <https://orcid.org/0000-0002-2637-4778>  
S. Schindler  <https://orcid.org/0000-0001-5507-8890>  
L. Schlickmann  <https://orcid.org/0000-0002-9746-6872>  
F. Schlüter  <https://orcid.org/0000-0002-5545-4363>  
J. Schneider  <https://orcid.org/0000-0001-7752-5700>  
F. G. Schröder  <https://orcid.org/0000-0001-8495-7210>  
L. Schumacher  <https://orcid.org/0000-0001-8945-6722>  
S. Sclafani  <https://orcid.org/0000-0001-9446-1219>  
L. Seen  <https://orcid.org/0009-0004-9204-0241>  
M. Seikh  <https://orcid.org/0000-0002-4464-7354>  
S. Seunarine  <https://orcid.org/0000-0003-3272-6896>  
P. Seville Myhr  <https://orcid.org/0009-0005-9103-4410>  
R. Shah  <https://orcid.org/0000-0003-2829-1260>  
N. Shimizu  <https://orcid.org/0000-0001-6857-1772>  
M. Silva  <https://orcid.org/0000-0001-6940-8184>

- B. Skrzypek  <https://orcid.org/0000-0002-0910-1057>  
 B. Smithers  <https://orcid.org/0000-0003-1273-985X>  
 J. Soedingrekso  <https://orcid.org/0000-0003-1011-2797>  
 D. Soldin  <https://orcid.org/0000-0003-3005-7879>  
 P. Soldin  <https://orcid.org/0000-0003-1761-2495>  
 G. Sommani  <https://orcid.org/0000-0002-0094-826X>  
 G. M. Spiczak  <https://orcid.org/0000-0002-0030-0519>  
 C. Spiering  <https://orcid.org/0000-0001-7372-0074>  
 J. Stachurska  <https://orcid.org/0000-0002-0238-5608>  
 T. Stezelberger  <https://orcid.org/0000-0003-2676-9574>  
 T. Stuttard  <https://orcid.org/0000-0001-7944-279X>  
 G. W. Sullivan  <https://orcid.org/0000-0002-2585-2352>  
 I. Taboada  <https://orcid.org/0000-0003-3509-3457>  
 S. Ter-Antonyan  <https://orcid.org/0000-0002-5788-1369>  
 M. Thiesmeyer  <https://orcid.org/0009-0003-0005-4762>  
 W. G. Thompson  <https://orcid.org/0000-0003-2988-7998>  
 J. Thwaites  <https://orcid.org/0000-0001-9179-3760>  
 K. Tollefson  <https://orcid.org/0000-0001-9725-1479>  
 S. Toscano  <https://orcid.org/0000-0002-1860-2240>  
 M. A. Unland Elorrieta  <https://orcid.org/0000-0002-6124-3255>  
 A. K. Upadhyay  <https://orcid.org/0000-0003-1957-2626>  
 N. Valtonen-Mattila  <https://orcid.org/0000-0002-1830-098X>  
 J. Vandenbroucke  <https://orcid.org/0000-0002-9867-6548>  
 N. van Eijndhoven  <https://orcid.org/0000-0001-5558-3328>  
 J. van Santen  <https://orcid.org/0000-0002-2412-9728>  
 S. Verpoest  <https://orcid.org/0000-0002-3031-3206>  
 A. Wang  <https://orcid.org/0009-0006-9420-2667>  
 C. Weaver  <https://orcid.org/0000-0003-2385-2559>  
 A. Y. Wen  <https://orcid.org/0009-0009-4869-7867>  
 C. Wendt  <https://orcid.org/0000-0001-8076-8877>  
 N. Whitehorn  <https://orcid.org/0000-0002-3157-0407>  
 C. H. Wiebusch  <https://orcid.org/0000-0002-6418-3008>  
 L. Witthaus  <https://orcid.org/0009-0000-0666-3671>  
 M. Wolf  <https://orcid.org/0000-0001-9991-3923>  
 S. Yoshida  <https://orcid.org/0000-0003-2480-5105>  
 F. Yu  <https://orcid.org/0000-0002-5775-2452>  
 S. Yu  <https://orcid.org/0000-0003-0035-7766>  
 T. Yuan  <https://orcid.org/0000-0002-7041-5872>  
 A. Zegarelli  <https://orcid.org/0000-0003-1497-3826>  
 S. Zhang  <https://orcid.org/0000-0002-2967-790X>  
 P. Zhelmin  <https://orcid.org/0000-0003-1019-8375>

## References

- Aartsen, M. G., Abbasi, R., Abdou, Y., et al. 2013a, *Sci*, **342**, 1242856  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2013b, *ApJ*, **779**, 132  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2015, *PhRvD*, **91**, 022001  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2017a, *JInst*, **12**, P03012  
 Aartsen, M. G., Abraham, K., Ackermann, M., et al. 2017b, *ApJ*, **835**, 151  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2017c, *ApJ*, **849**, 67  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2018a, *Sci*, **361**, eaat1378  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2018b, *Sci*, **361**, 147  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2020a, *PhRvL*, **124**, 051103  
 Aartsen, M. G., Ackermann, M., Adams, J., et al. 2020b, *PhRvL*, **125**, 121104  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2021a, *PhRvD*, **104**, 022002  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2021b, *ICRC2021*, **395**, 1130  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2022a, *Sci*, **378**, 538  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2022b, *ApJ*, **928**, 50  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2023a, *Sci*, **380**, adc9818  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2023b, *ApJS*, **269**, 25  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2024, *PhRvD*, **110**, 022001  
 Abbasi, R., Ackermann, M., Adams, J., et al. 2025, *ICRC2025*, **501**, 1219  
 Abdalla, H., Abramowski, A., Aharonian, F., et al. 2018, *A&A*, **612**, A1  
 Abdollahi, S., Acero, F., Ackermann, M., et al. 2020, *ApJS*, **247**, 33  
 Abramowski, A., Acero, F., Aharonian, F., et al. 2012, *A&A*, **537**, A114  
 Ackermann, M., Ajello, M., Atwood, W. B., et al. 2012, *ApJ*, **750**, 3  
 Aharonian, F., Akhperianian, A. G., Bazer-Bachi, A. R., et al. 2007, *ApJ*, **661**, 236  
 Albert, A., André, M., Anghinolfi, M., et al. 2017, *PhRvD*, **96**, 082001  
 Albert, A., André, M., Anghinolfi, M., et al. 2018, *ApJL*, **868**, L20  
 Argüelles, C. A., Palomares-Ruiz, S., Schneider, A., Wille, L., & Yuan, T. 2018, *JCAP*, **1807**, 047  
 Astropy Collaboration, Price-Whelan, A. M., Lim, P. L., et al. 2022, *ApJ*, **935**, 167  
 Astropy Collaboration, Price-Whelan, A. M., Sipőcz, B. M., et al. 2018, *AJ*, **156**, 123  
 Aublin, J., Albert, A., André, M., et al. 2019, *ICRC2019*, **358**, 840  
 Braun, J., Dumm, J., De Palma, F., et al. 2008, *Aph*, **29**, 299  
 Cartraud, T., Albert, A., André, M., et al. 2025, *ICRC2025*, **501**, 1009  
 de la Torre, P., Gaggero, D., Grasso, D., et al. 2022, The Galactic diffuse gamma-ray emission meets the PeV frontier - All-sky distributions, v2, Zenodo, doi:10.5281/zenodo.6385205  
 Fuerst, P. M., Abbasi, R., Ackermann, M., et al. 2024, *ICRC2023*, **444**, 1046  
 Gaggero, D., Grasso, D., Marinelli, A., Urbano, A., & Valli, M. 2015, *ApJL*, **815**, L25  
 Gaisser, T. K., Jero, K., Karle, A., & van Santen, J. 2014, *PhRvD*, **90**, 023009  
 Górski, K. M., Hivon, E., Banday, A. J., et al. 2005, *ApJ*, **622**, 759  
 Halzen, F. 2006, *EPJC*, **46**, 669  
 Ke, G., Meng, Q., Finley, T., et al. 2017, Advances in Neural Information Processing Systems, **30**, 3146  
 Mancina, S., Silva, M., Aartsen, M. G., et al. 2019, *ICRC2019*, **358**, 954  
 Osborn, J., Mancina, S., Silva, M., et al. 2024, *ICRC2023*, **444**, 1464  
 Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33  
 Savina, P., Ackermann, M., Adams, J., et al. 2023, *ICRC2023*, **444**, 1010  
 Schöner, S., Gaisser, T. K., Resconi, E., & Schulz, O. 2009, *PhRvD*, **79**, 043009  
 Tanabashi, M., Hagiwara, K., Hikasa, K., et al. 2018, *PhRvD*, **98**, 030001  
 Thiesmeyer, M., Yuan, T., Seen, L., et al. 2025, *ICRC2025*, **501**, 1193  
 Wakely, S. P., & Horan, D. 2008, *ICRC2007*, **3**, 1341  
 Zonca, A., Singer, L., Lenz, D., et al. 2019, *JOSS*, **4**, 1298